Neutrinos from SN1987A
probing the inner properties of supernova

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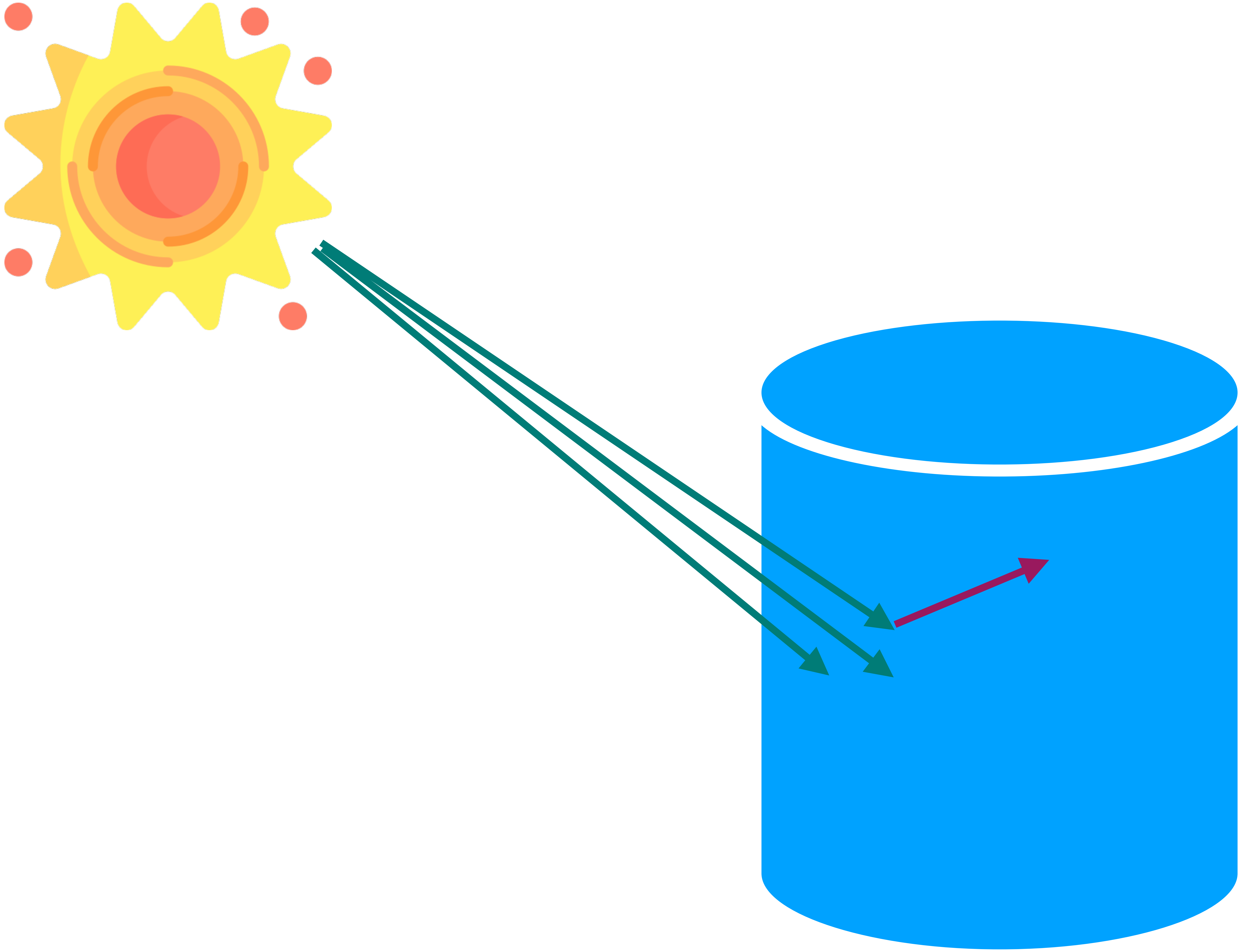
UNIVERSITY OF COPENHAGEN

VILLUM FONDEN

It is February 23 1987. At 49.59 ± 0.09stat ± 0.54syst kpc from Earth, blue supergiant Sanduleak -69 202 is going to explode…

Credits to E. Vitagliano

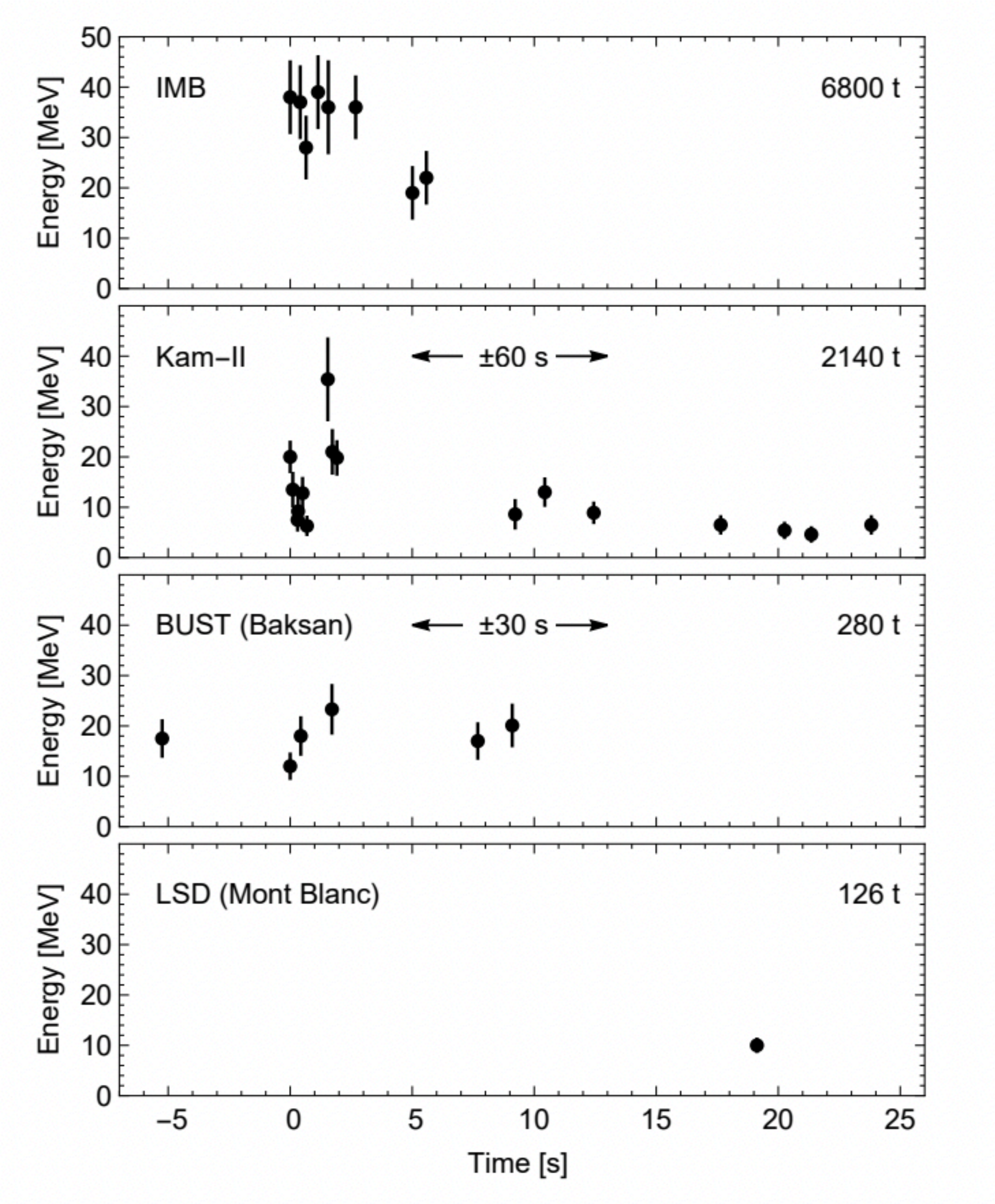
SN1987A neutrino observations



First time we had instruments to detect neutrinos

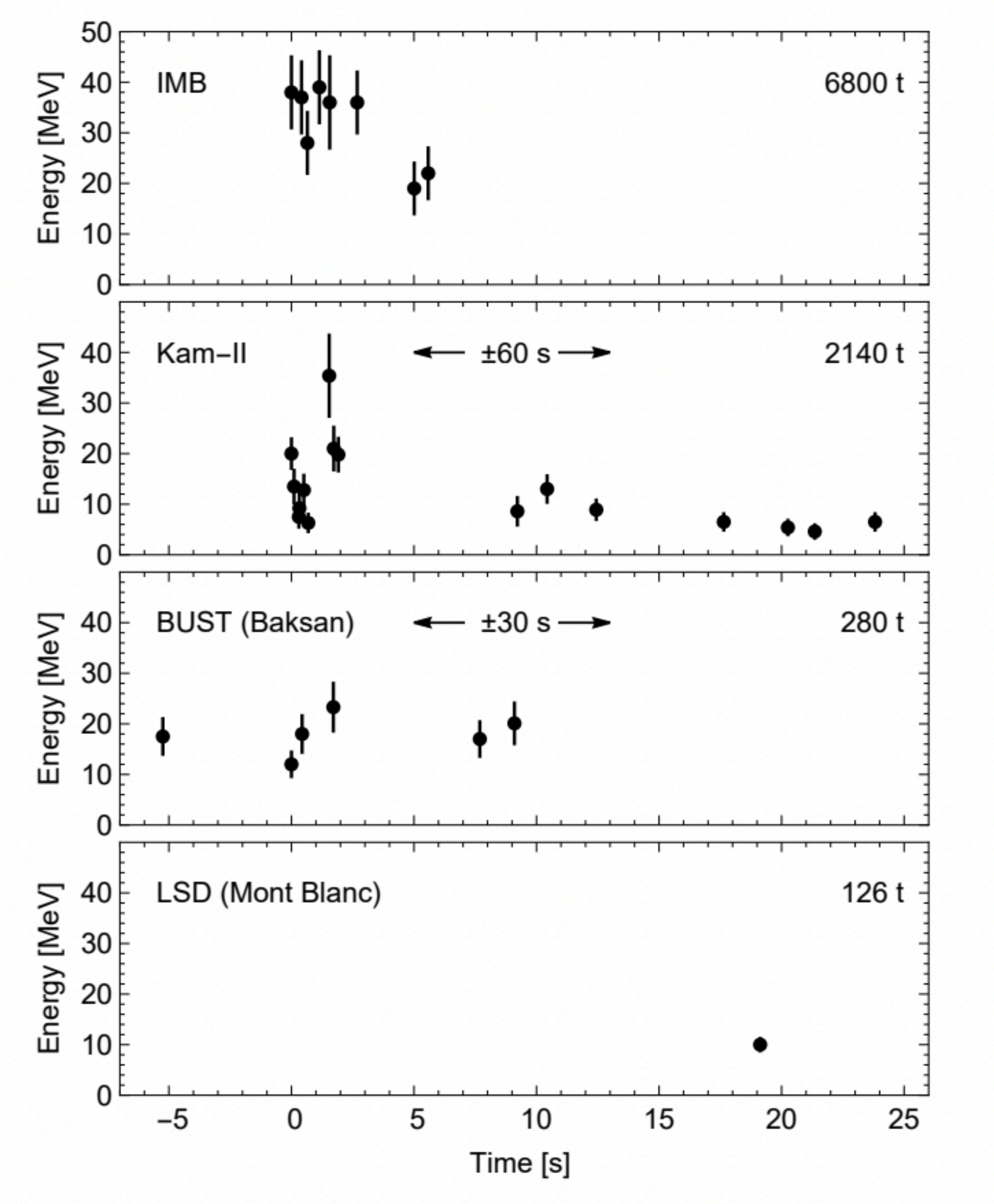
Last time we had a supernova close enough to Earth

SN1987A neutrino observations



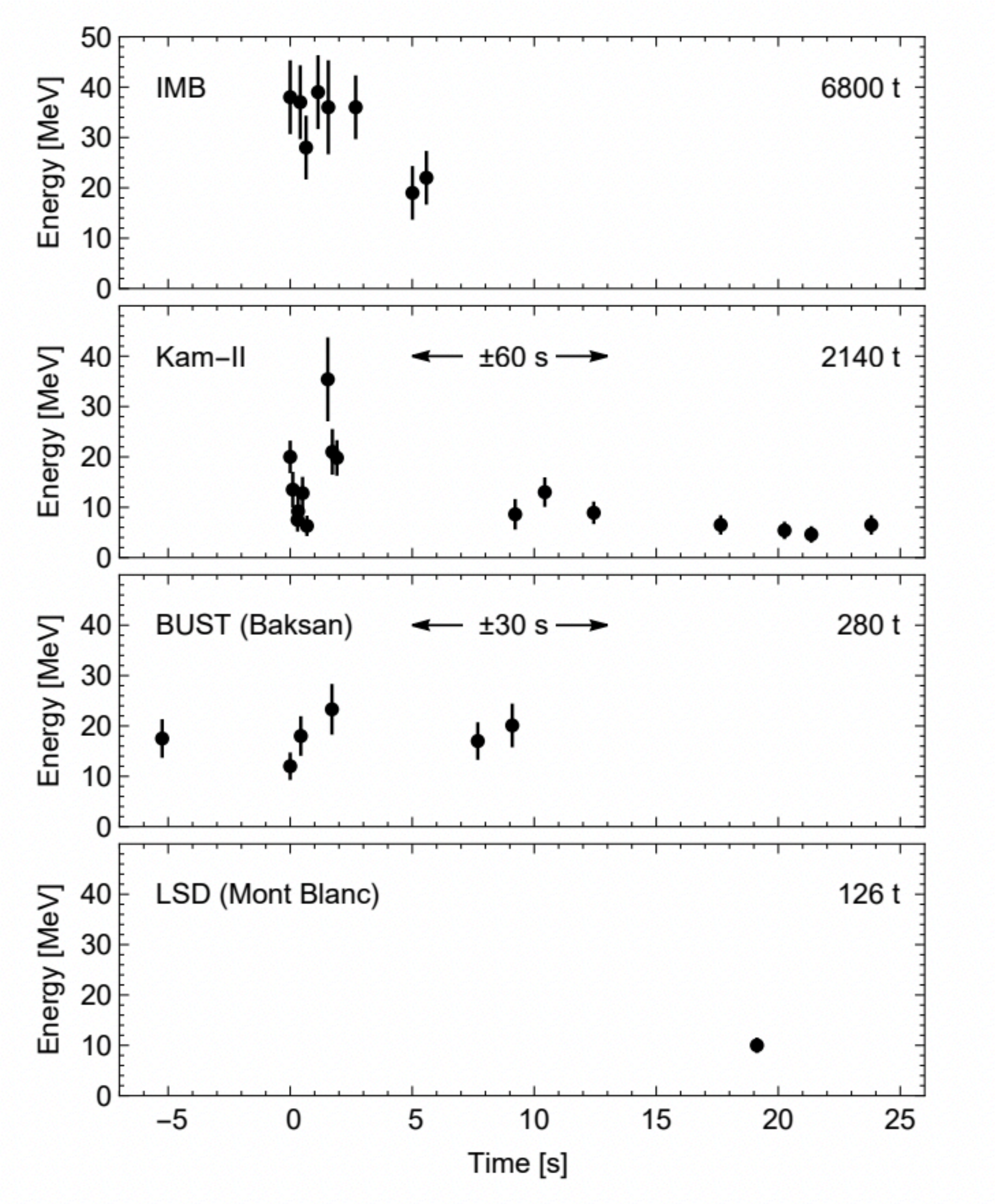
- ✦ About 15-20 neutrinos
- ✦ Typical energy 10-40 MeV
- ✦ Timescale of 5-7 seconds, with 4-5 events at 9-10 seconds

SN1987A neutrino observations



What do we learn on supernovae?

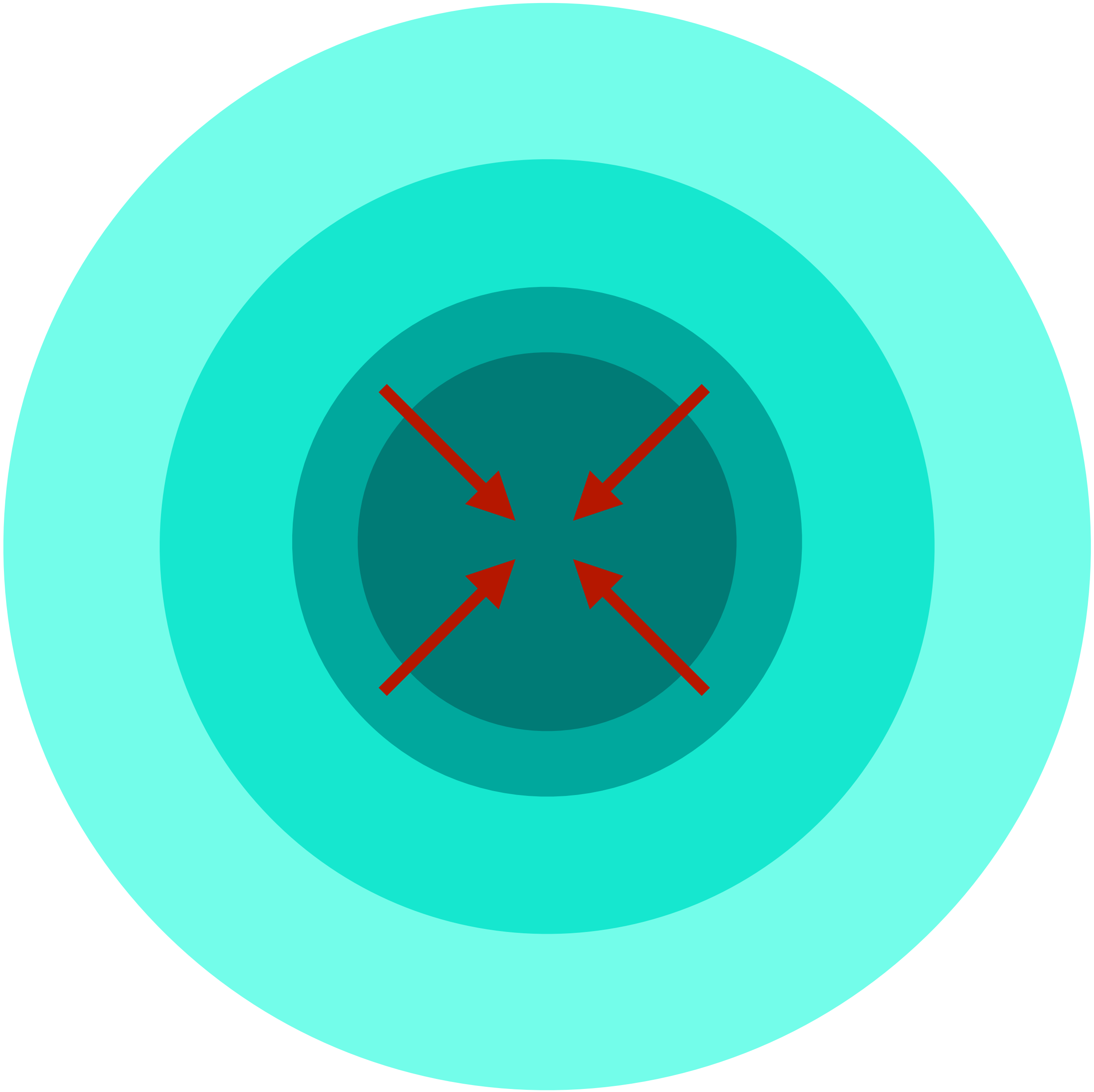
SN1987A neutrino observations



What do we learn by comparing SN 1987A with state-of-the-art SN models?

Supernova neutrinos

Core-Collapse Supernovae

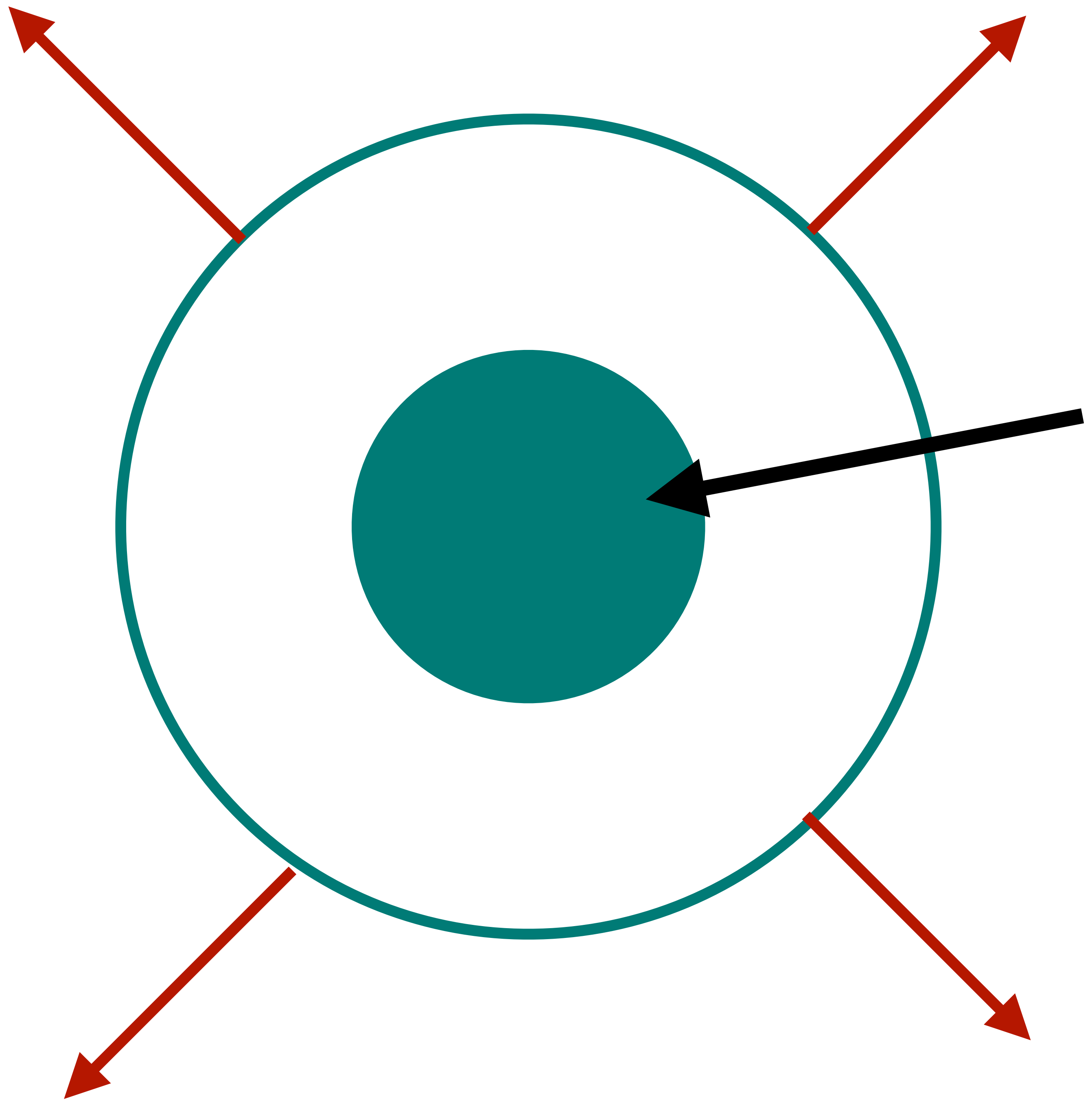


- Iron core with stopping of fusion
- Core collapses up to nuclear densities

$\rho \sim 10^{11}\ \mathrm{g\ cm^{-3}}$ Neutrinos are trapped

$\rho \sim 10^{14}\ \mathrm{g\ cm^{-3}}$ Rebounce

Core-Collapse Supernovae



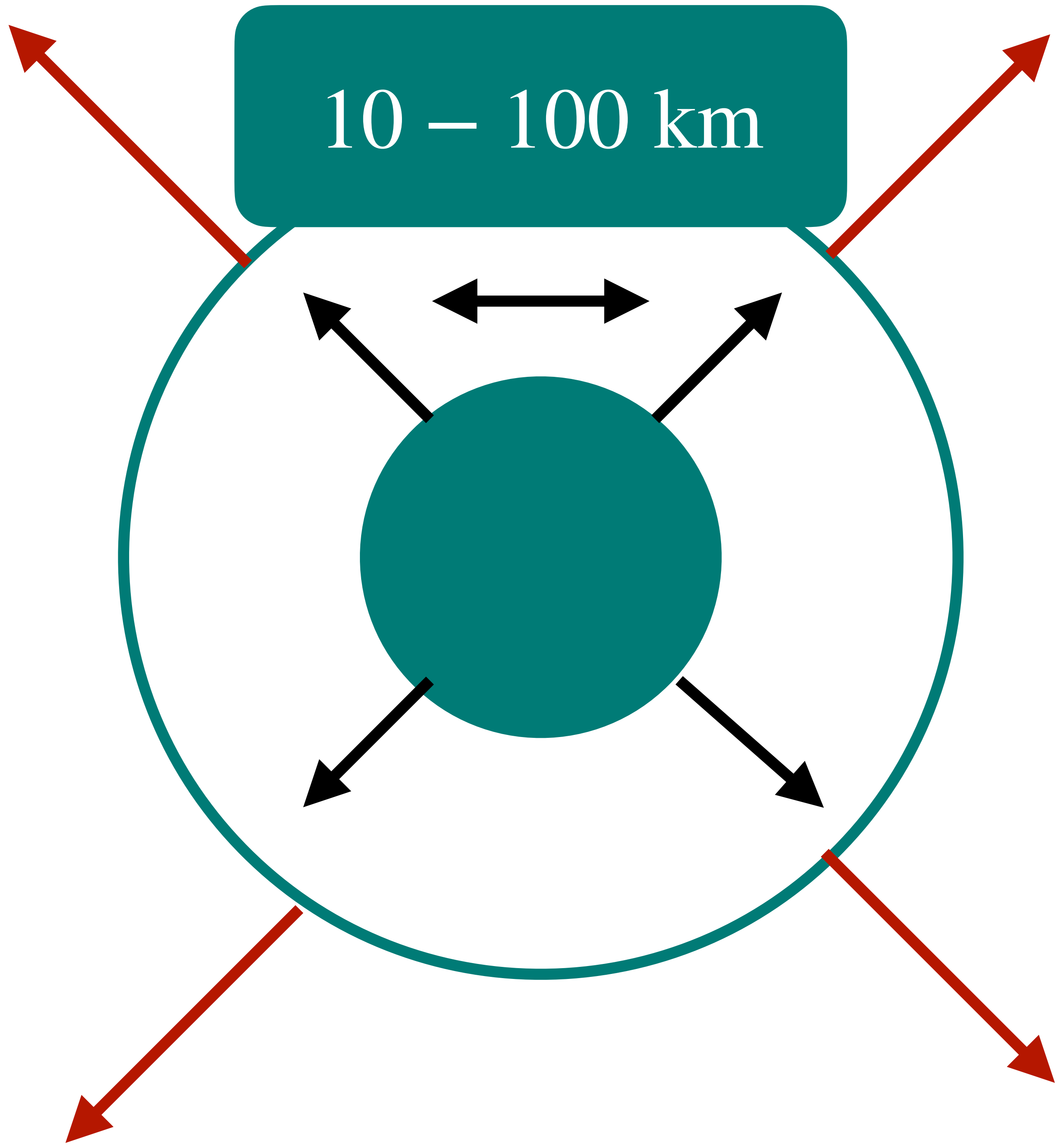
Neutrinos trapped and degenerate!

$$\tau_{\rm acc} \simeq 0.3 - 0.7 \text{ s}$$

- Iron core with stopping of fusion
- Core collapses up to nuclear densities
- Rebounce launches shock wave, leaving **accreting** proto-neutron star (PNS) in the center
- Shock wave stalls, revived by neutrinos depositing energy

Core-Collapse Supernovae

$E_\nu \simeq 10^{53}$ erg



$\tau_{\rm cool} \simeq 5 - 10$ s

- Iron core with stopping of fusion
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- Shock wave stalls, revived by neutrinos depositing energy
- Neutrino **cooling** of PNS

Core-Collapse Supernovae

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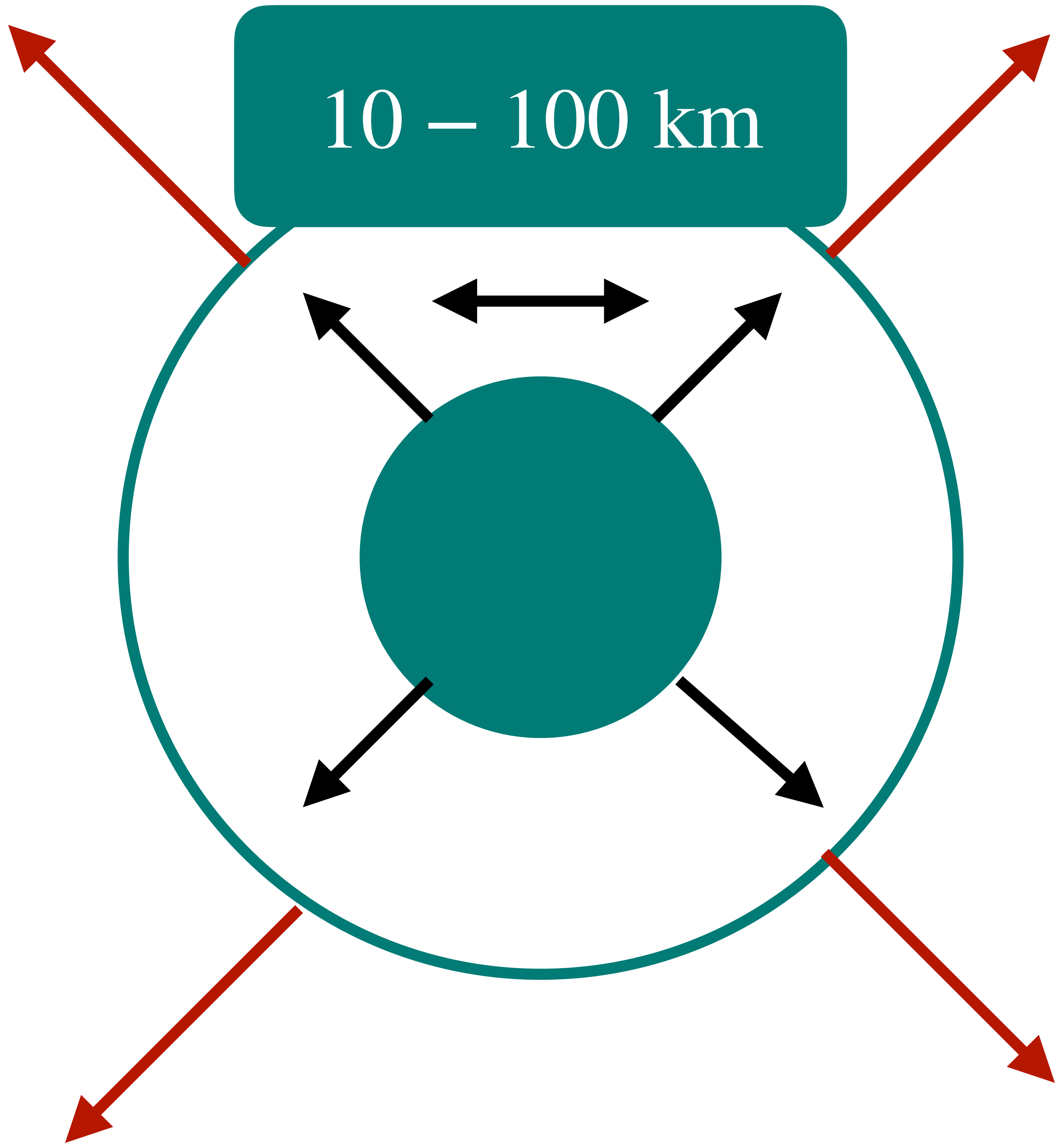

PNS mass main order parameter for neutrino flux!

- Iron core with stopping of fusion
- Core collapses up to nuclear densities
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Supernova simulations confront SN 1987A neutrinos

Motivation

- Increased confidence in the neutrino delayed explosion mechanism - 3D simulations show self-consistent explosions
- Significant updates to the simulations
    - **Convection**
    - **Updated neutrino-nucleon opacities**

Model choice

PNS mass ($M_\odot$)

1.36 1.44 1.62 1.77 1.93

Equation of state (EoS)

DD2

LS220

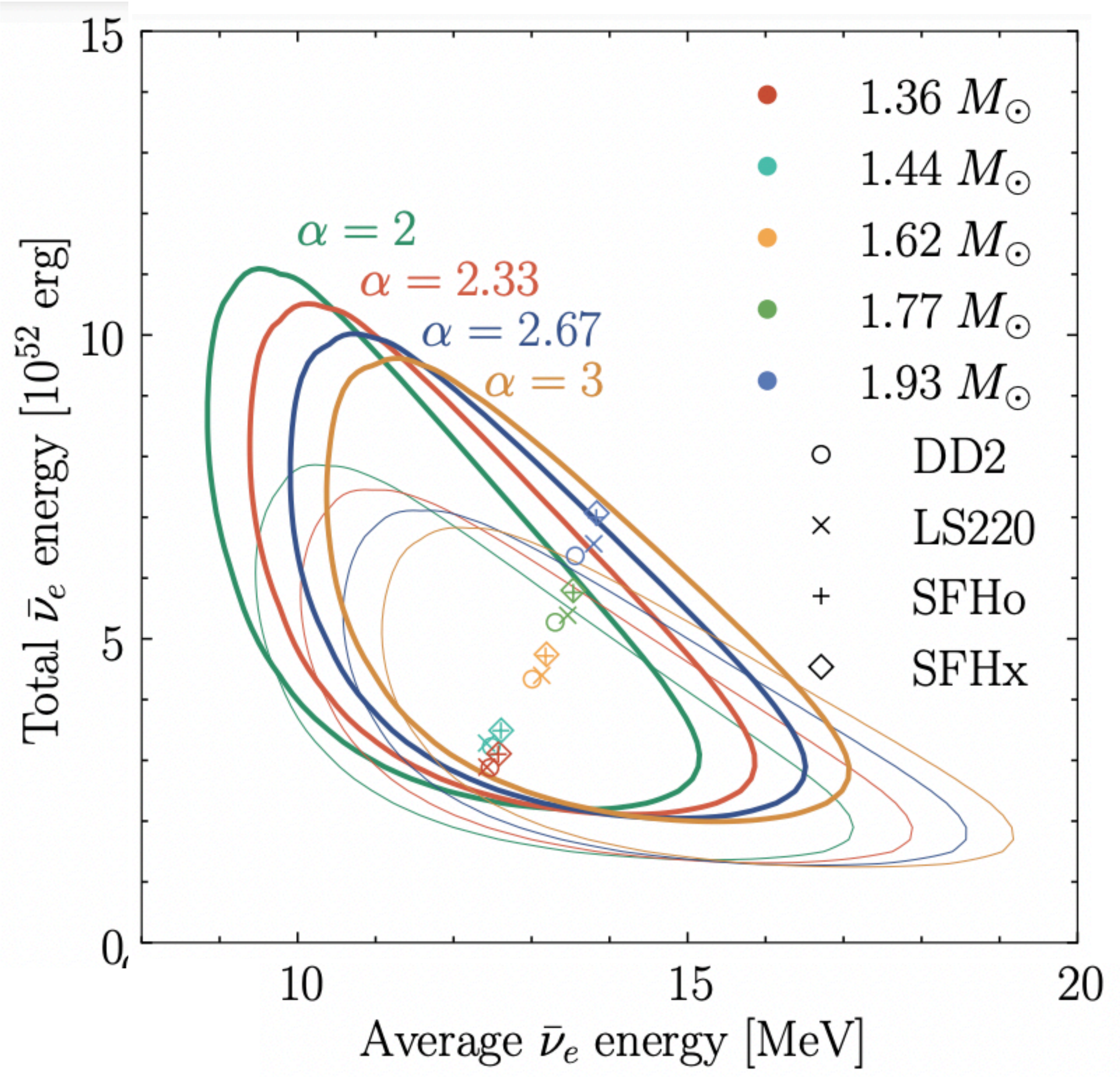
SFHo

SFHx

All 1D

(Reliable for the sparse data of SN 1987A!)

Time-integrated signal



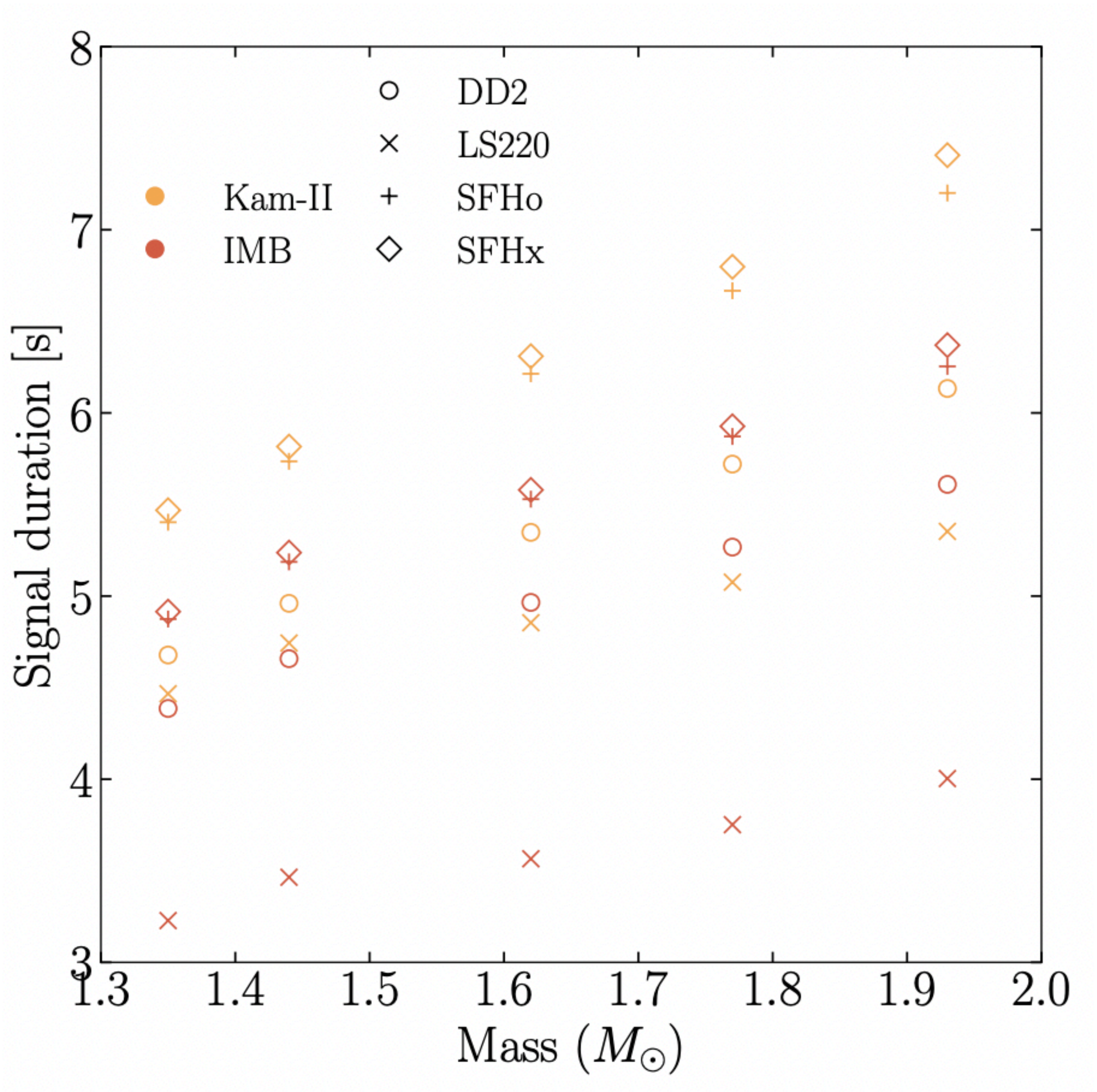
- ✦ Spectra can be pinched

$$\frac{d\mathcal{F}_{\bar{\nu}_e}}{d\epsilon_\nu} = \frac{E_{\rm tot}^{\bar{\nu}_e}}{\Gamma_{1+\alpha}\bar{\epsilon}^2}\frac{(1+\alpha)^{1+\alpha}}{4\pi d_{\rm SN}^2}\left(\frac{\epsilon_\nu}{\bar{\epsilon}}\right)^\alpha e^{-(1+\alpha)\epsilon_\nu/\bar{\epsilon}}$$

- ✦ Most SN models lie within 2σ regions — consistency with data

- ✦ Tension with heavy PNS

Time structure of the signal



- Signal duration less than 8 seconds for **all** models
- Tension with late-time Kam-II events
- Key role played by convection and updated neutrino-nucleon opacities

Conclusions

- SN 1987A ideal laboratory for astrophysics
- Generally consistent with modern simulations, both all-duration and first second
- Requires light PNS $\lesssim 1.8\, M_\odot$
- Origin of late-time events?

Backup slides

Model choice

- 3D models more reliable on the detailed time structure during first second
- 1D and 3D models consistent on the time-integrated signal during first second

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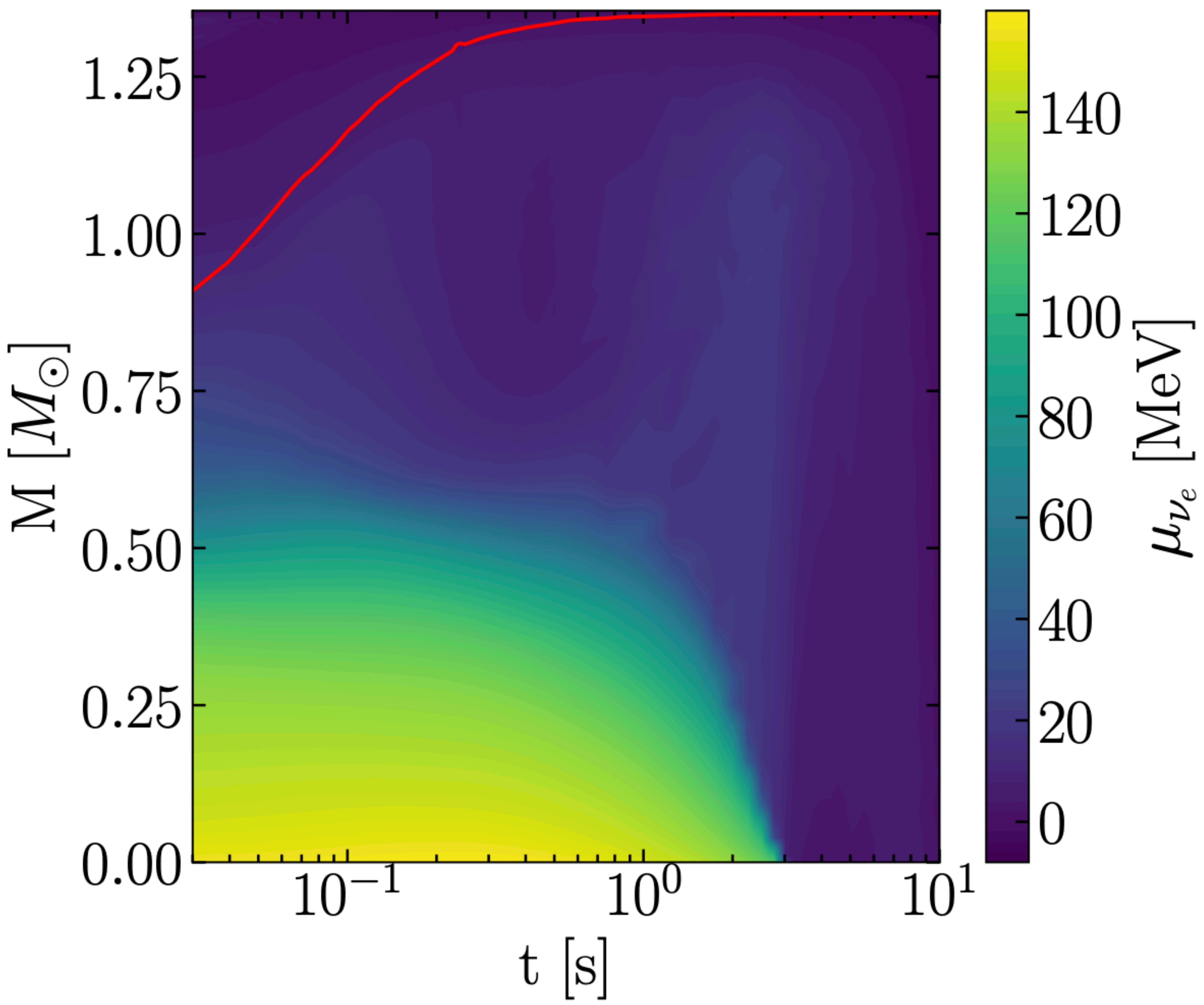
- 3D models have severe limitations
  - Cannot systematically scan parameter space (PNS mass)

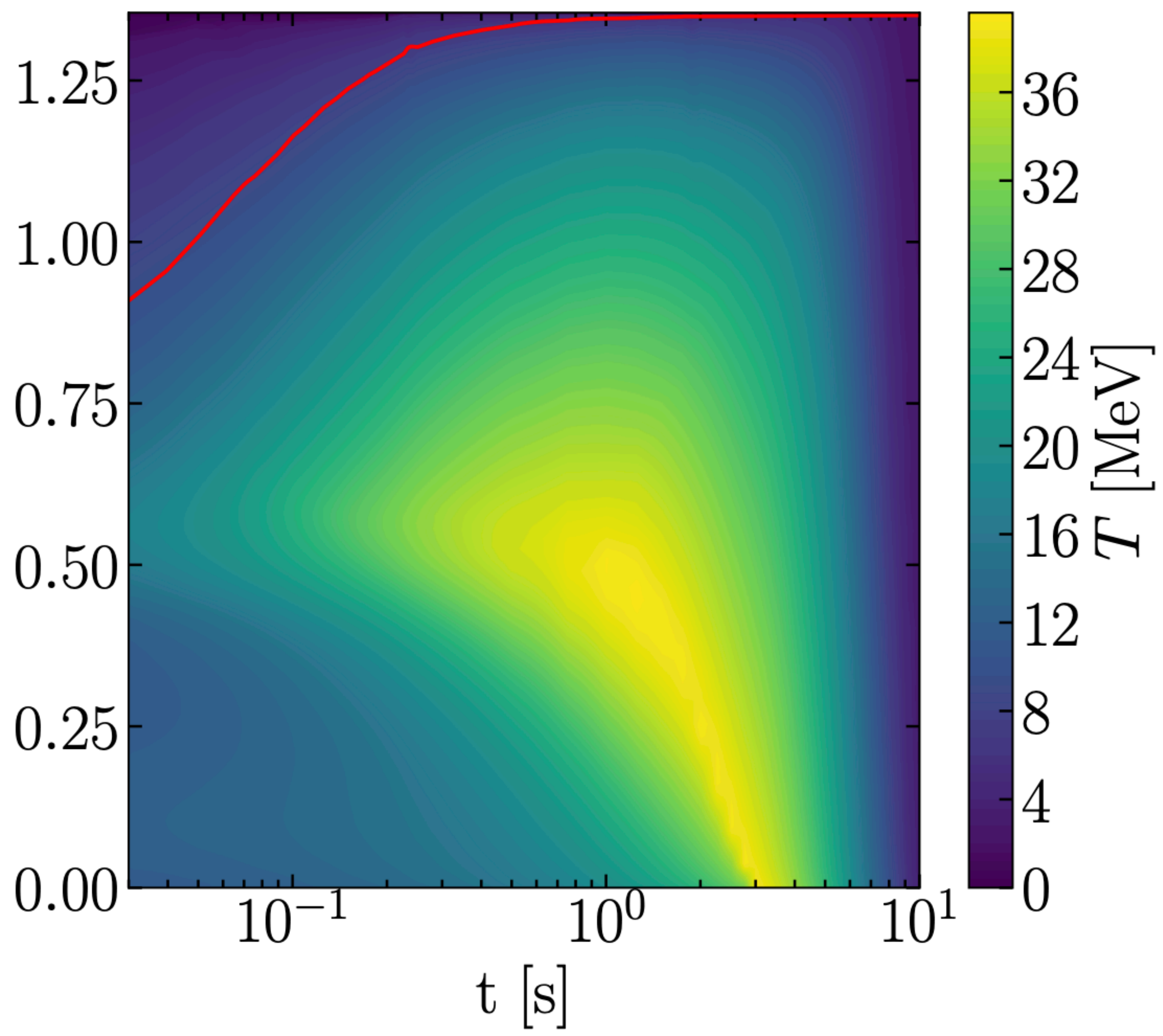
Model choice

- 3D models more reliable on the detailed time structure during first second
- 1D and 3D models consistent on the time-integrated signal during first second
- 3D models have severe limitations
    - Cannot systematically scan parameter space (PNS mass)
    - Cannot extend to more than 1 second (statistical pitfalls?)

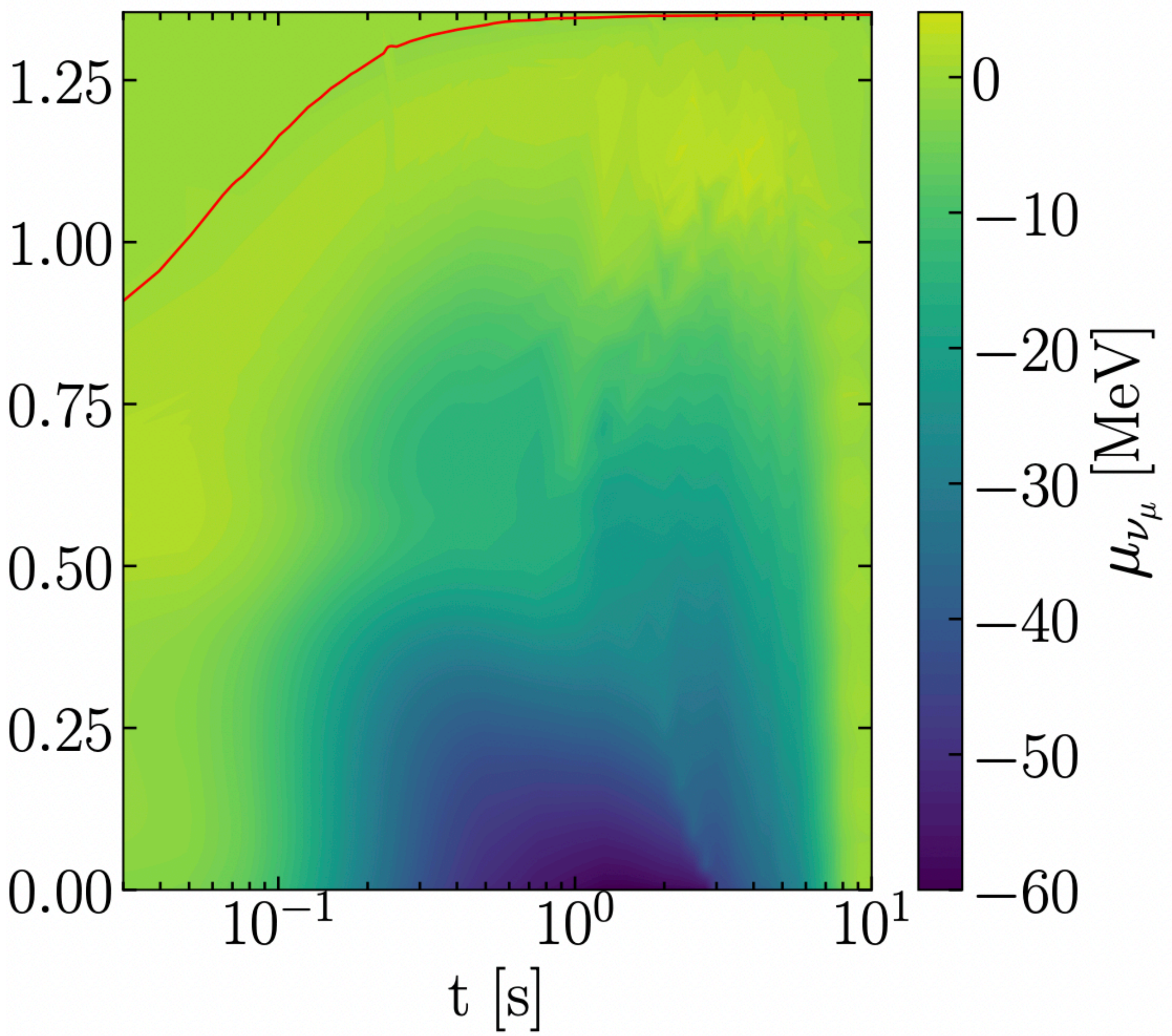
Core-Collapse Supernovae



PNS deleptonizes and cools

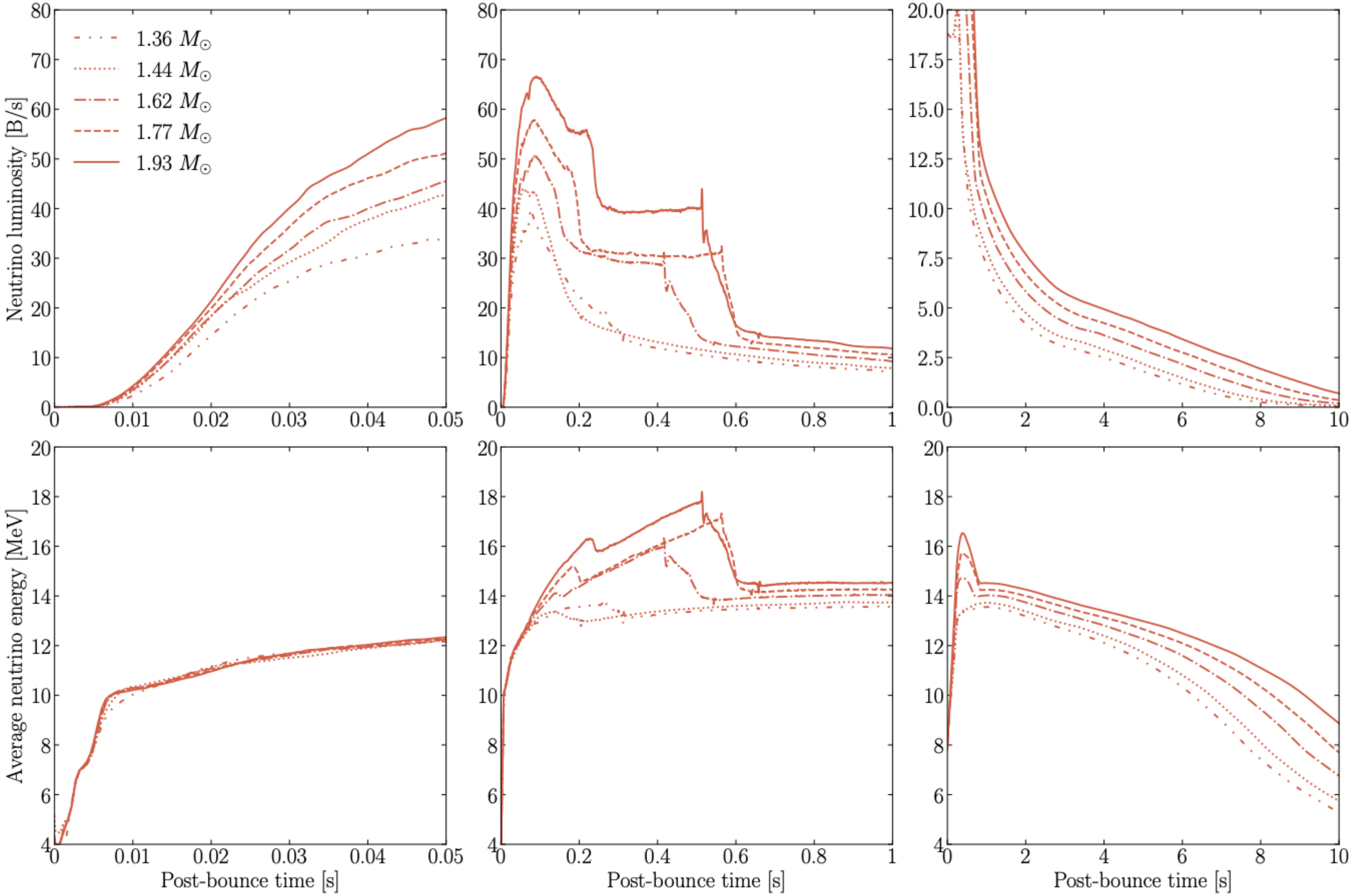


Heats up the external material

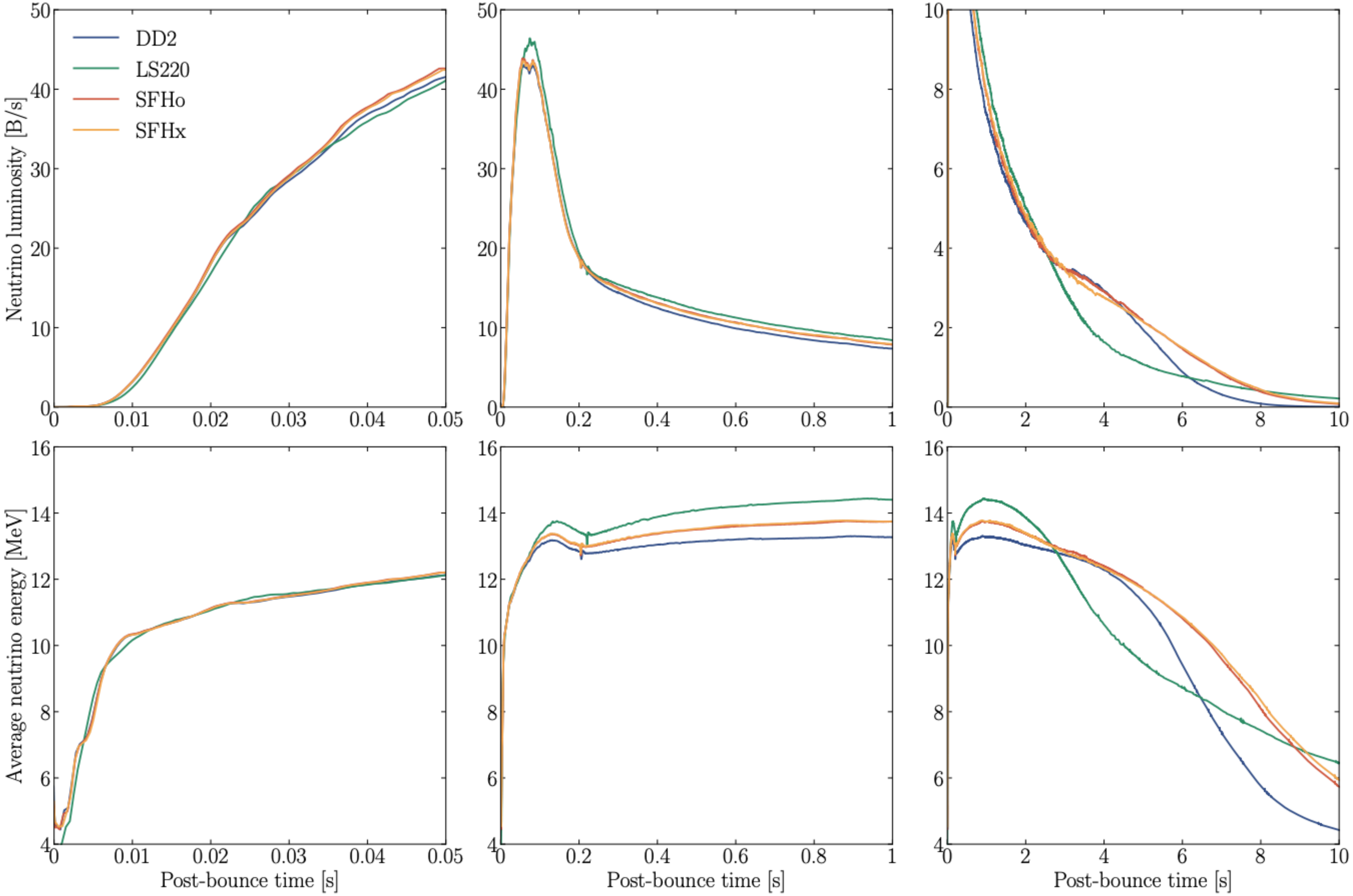


Produces muons and muon neutrinos

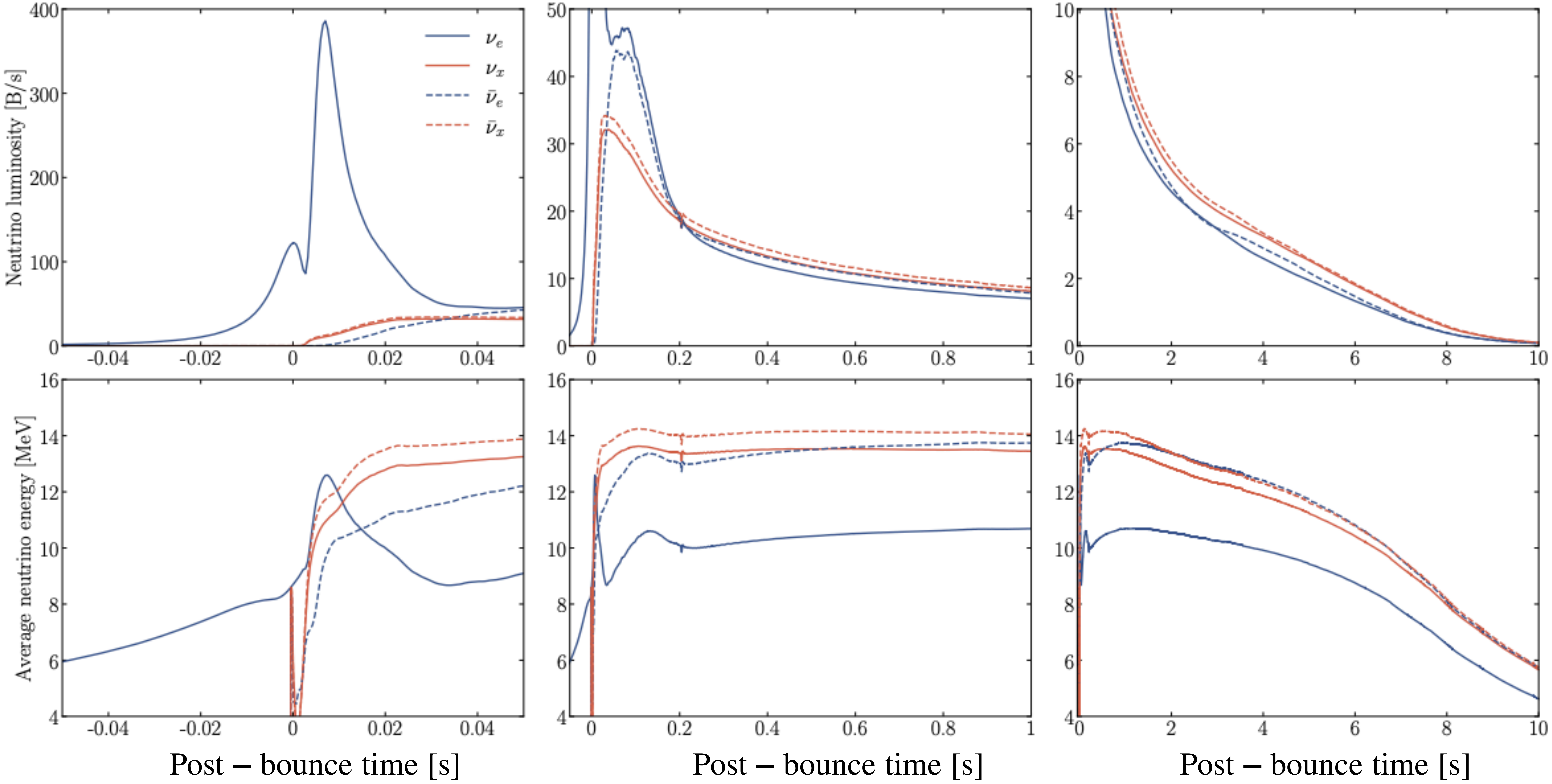
SN models - neutrino signal



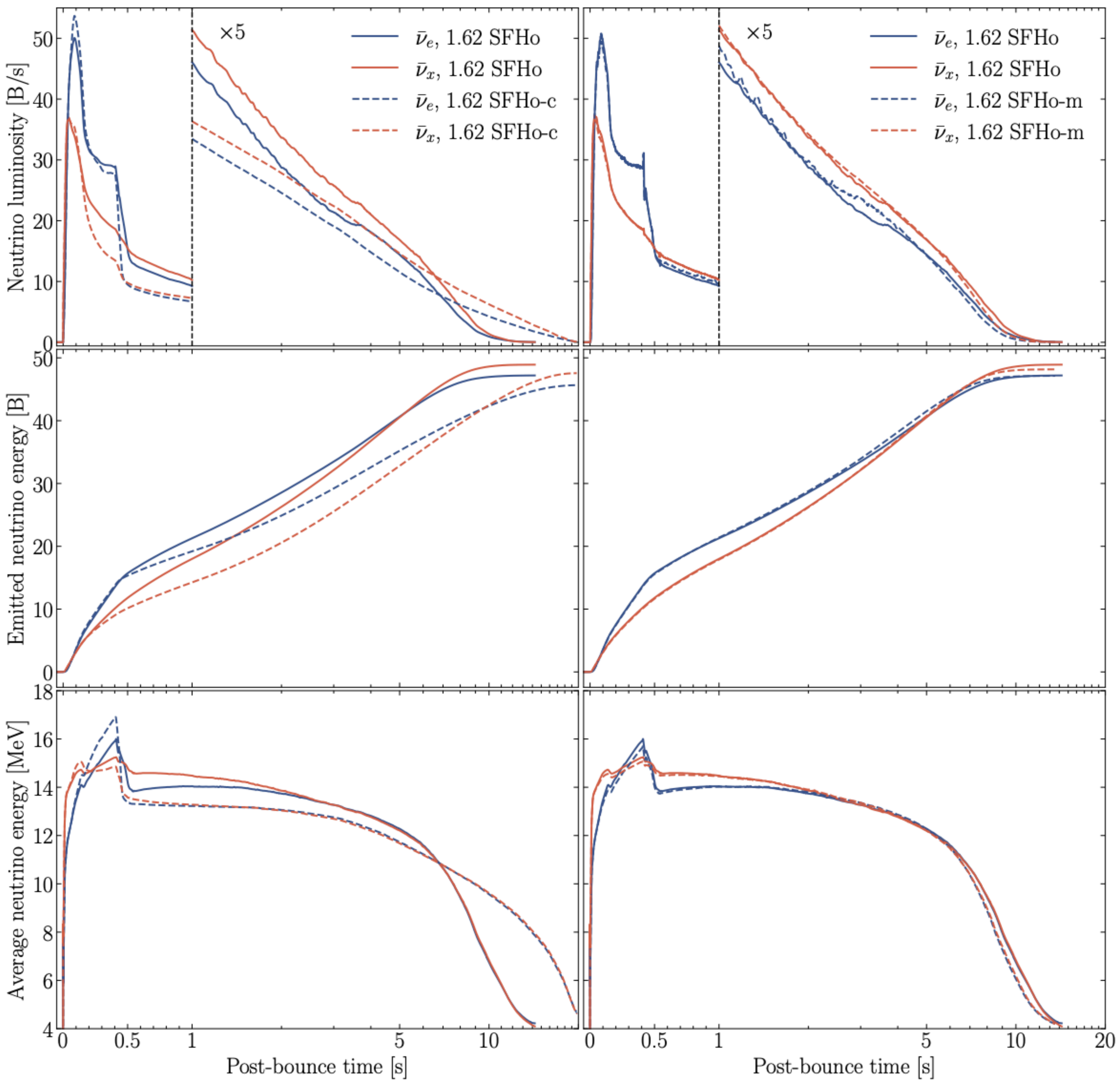
SN models - neutrino signal



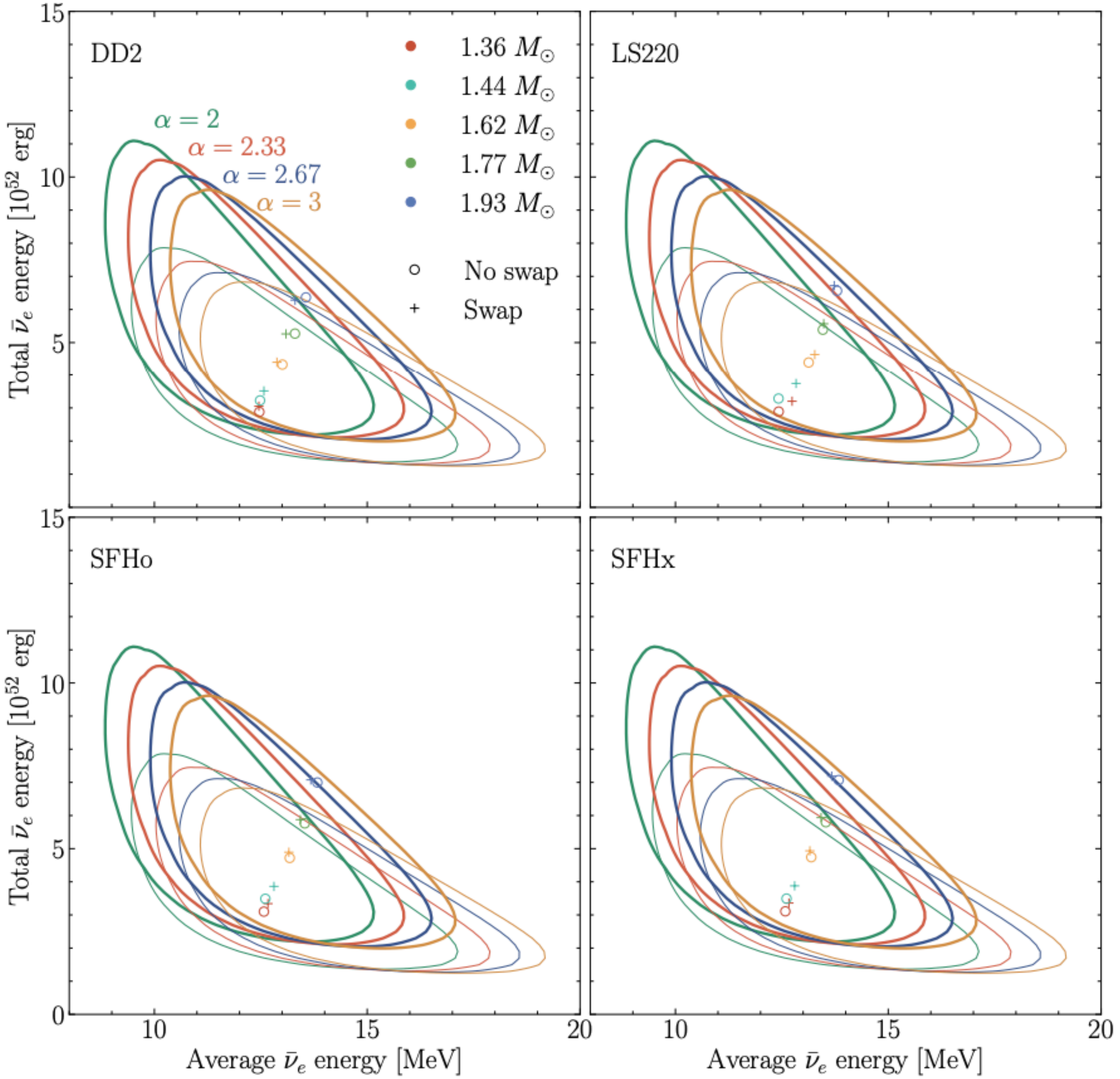
Flavor dependence of neutrino signal



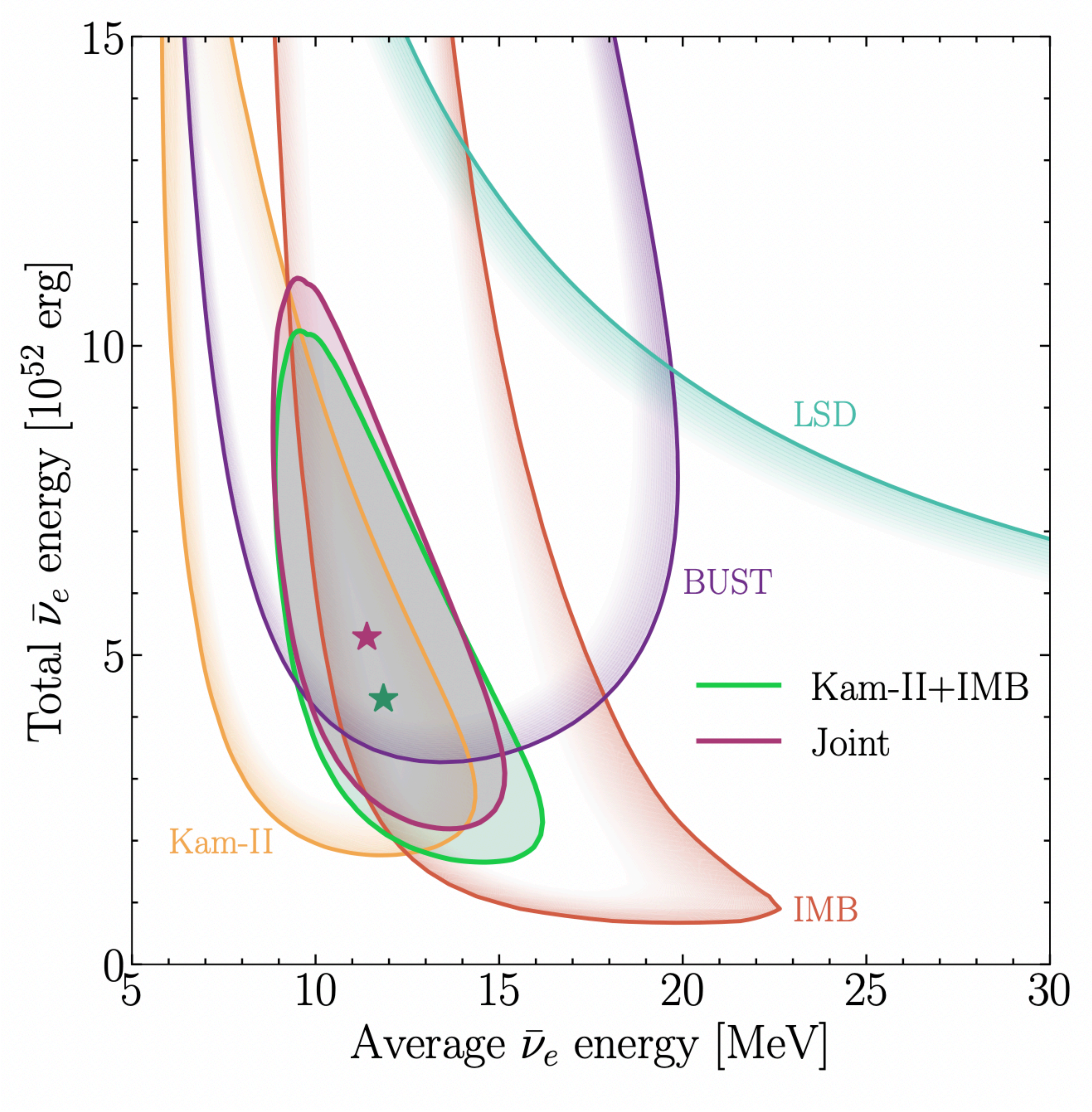
Convection vs. no convection



Impact of flavor conversion

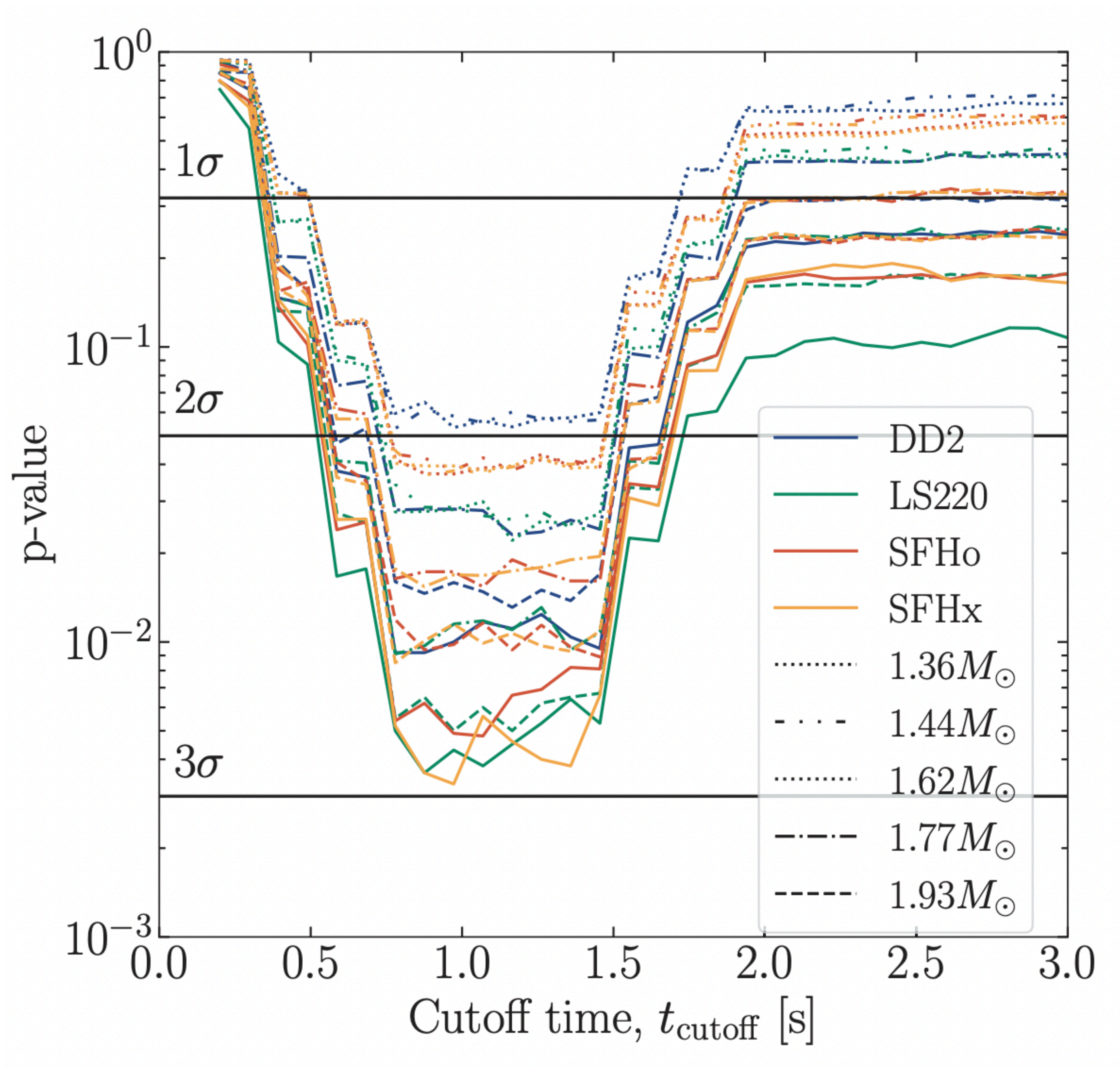


Time-integrated signal



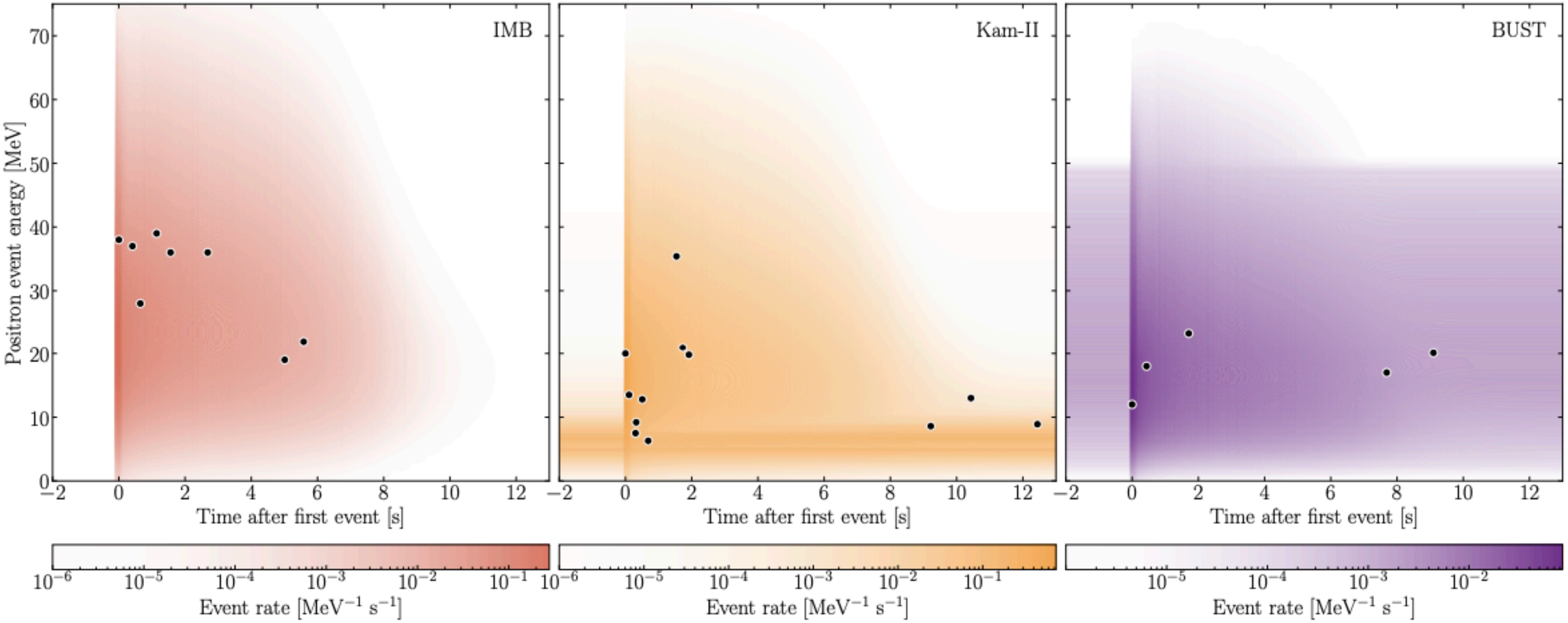
- ✦ Tension between Kam-II and IMB — slightly relieved, less than 2σ
- ✦ First combined analysis including all experiments!
- ✦ Assuming neutrino blackbody spectrum

First second of emission

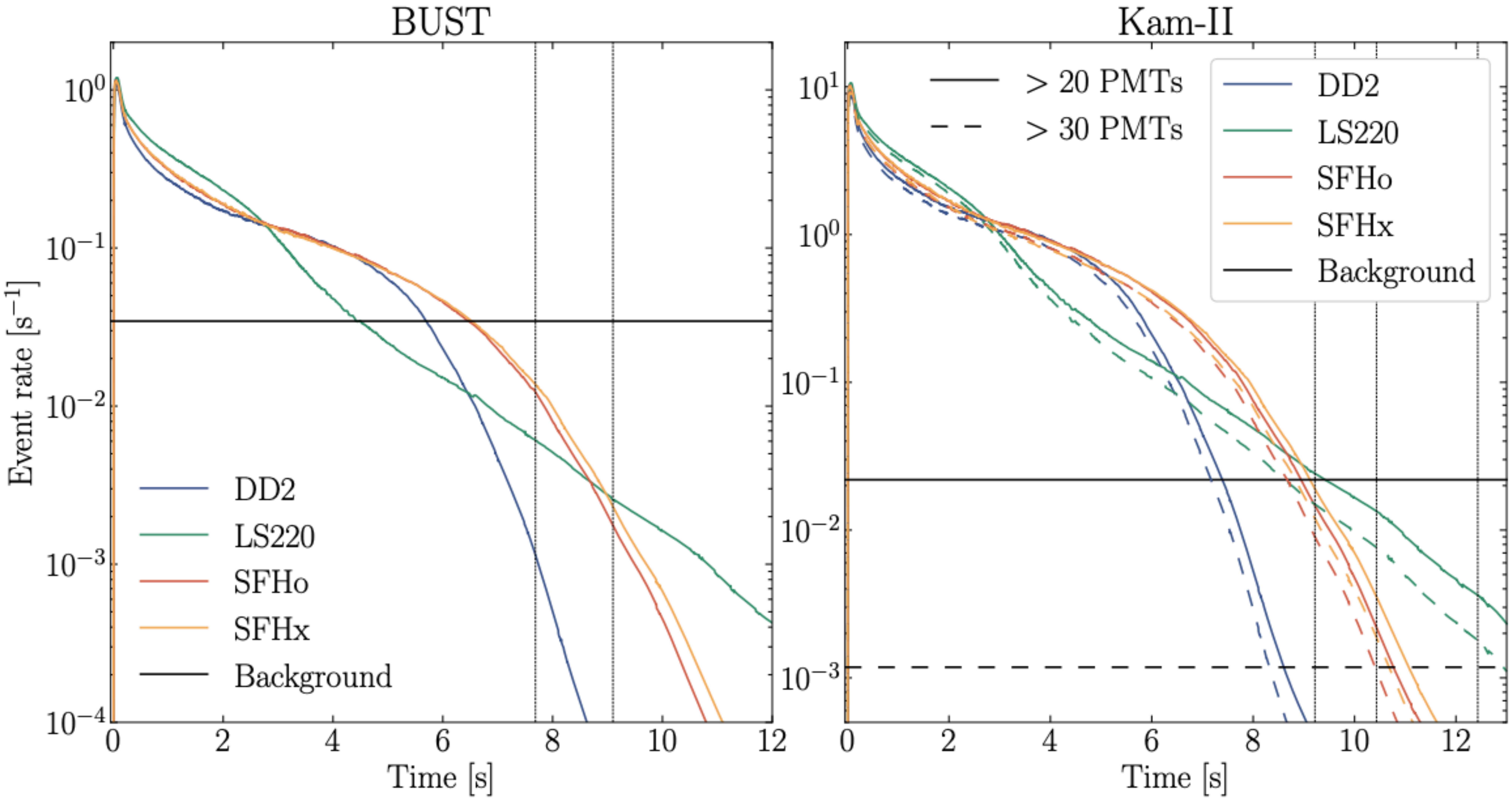


- Kolmogorov-Smirnov on first-second events to compare with Li et al., 2306.08024
- Cutting at 1 s maximizes tension (events 3 and 4 have low energy), but globally insignificant
- Models with low PNS less than 2 sigma even cutting the events

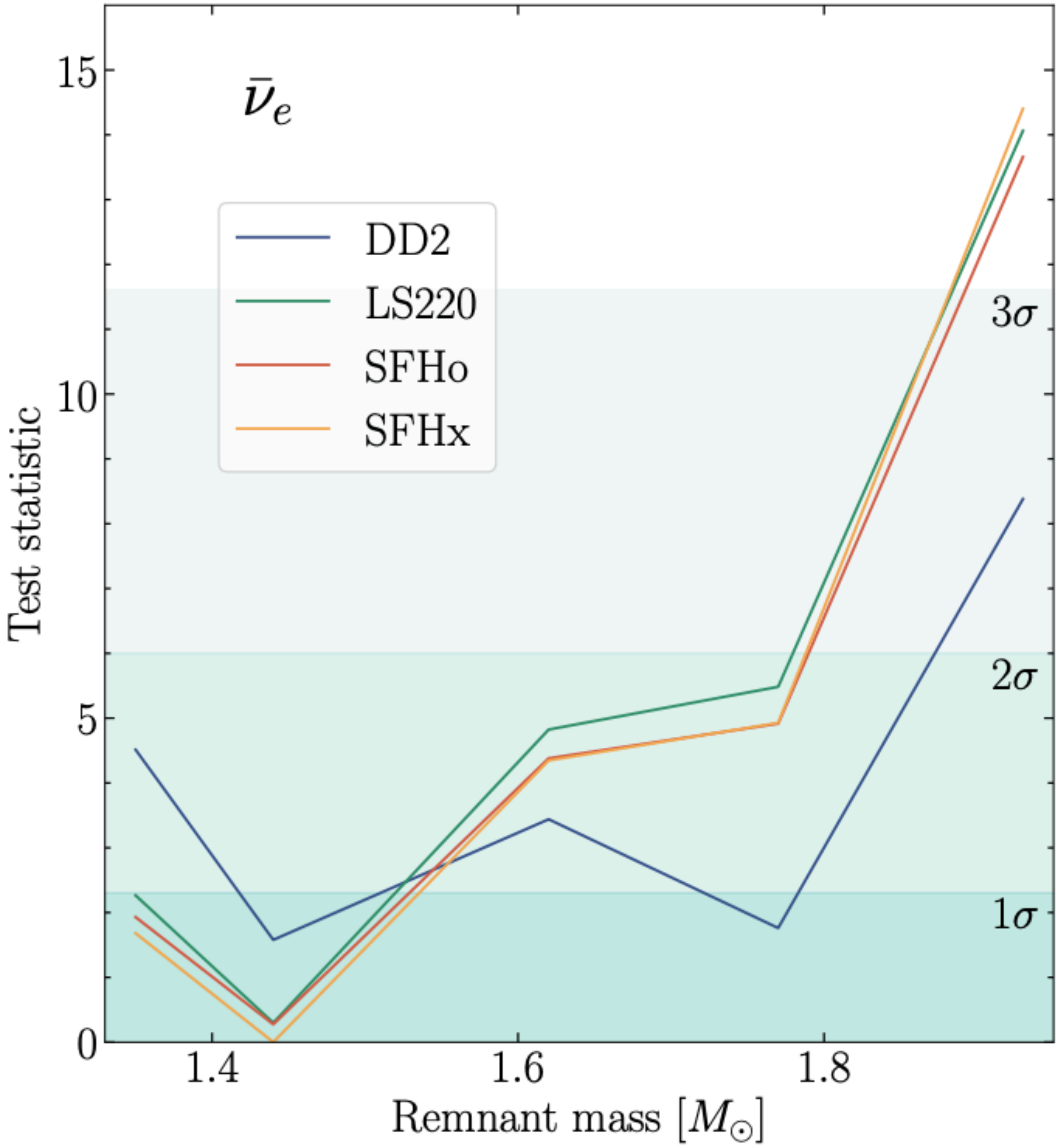
Event rates


IMB
Kam-II
BUST
Positron event energy [MeV]
0
10
20
30
40
50
60
70
−2
0
2
4
6
8
10
12
Time after first event [s]
Event rate [MeV^{-1} s^{-1}]
10^{-6}
10^{-5}
10^{-4}
10^{-3}
10^{-2}
10^{-1}

Late-time events

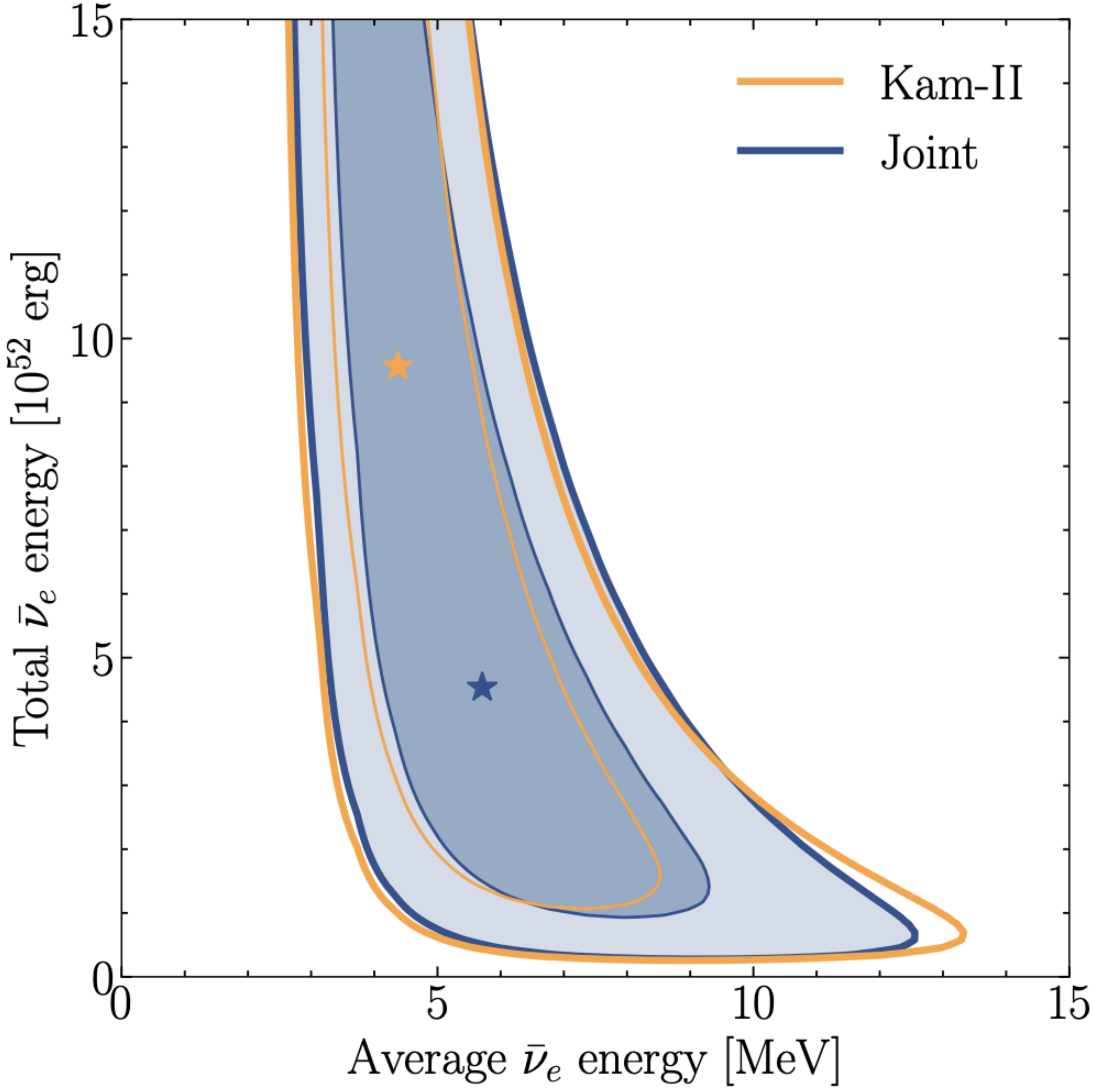


Full time and energy analysis



- ✦ Bimodal tendency — Kam-II and LSD point to light PNS, IMB and BUST to heavy PNS
- ✦ PNS mass of $1.93\ M_\odot$ excluded
- ✦ Weak sensitivity to EoS

Time structure of the signal



- Origin of late-time events is an open question
  - Background?
  - Late-time fallback accretion?

Why supernova?

- Endpoint of massive stars
- Internal densities reach up to nuclear densities (10^{14} g cm^{-3})
- Internal temperatures reach up to $30-40$ MeV
- Extreme conditions make even rare processes possible

Testing for new physics

New particles can be produced in supernova core…

Coupled to photons

- Axion-like particles
- Dark photons

Coupled to nucleons

- QCD axion
- Nucleophilic dark matter

Coupled to neutrinos

- Gauge bosons ($B - L$, $L_\mu - L_\tau$)
- Secret interactions
- Pseudo-majorons

Testing for new physics

New particles can be produced in supernova core…

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New particles can be produced in supernova core…

… but how do we probe them?

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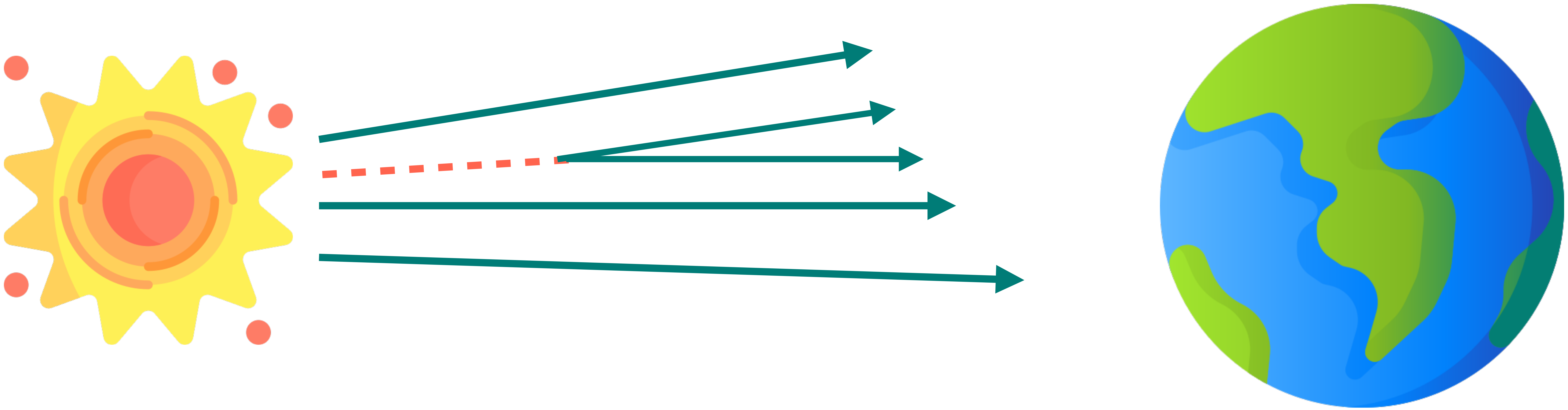
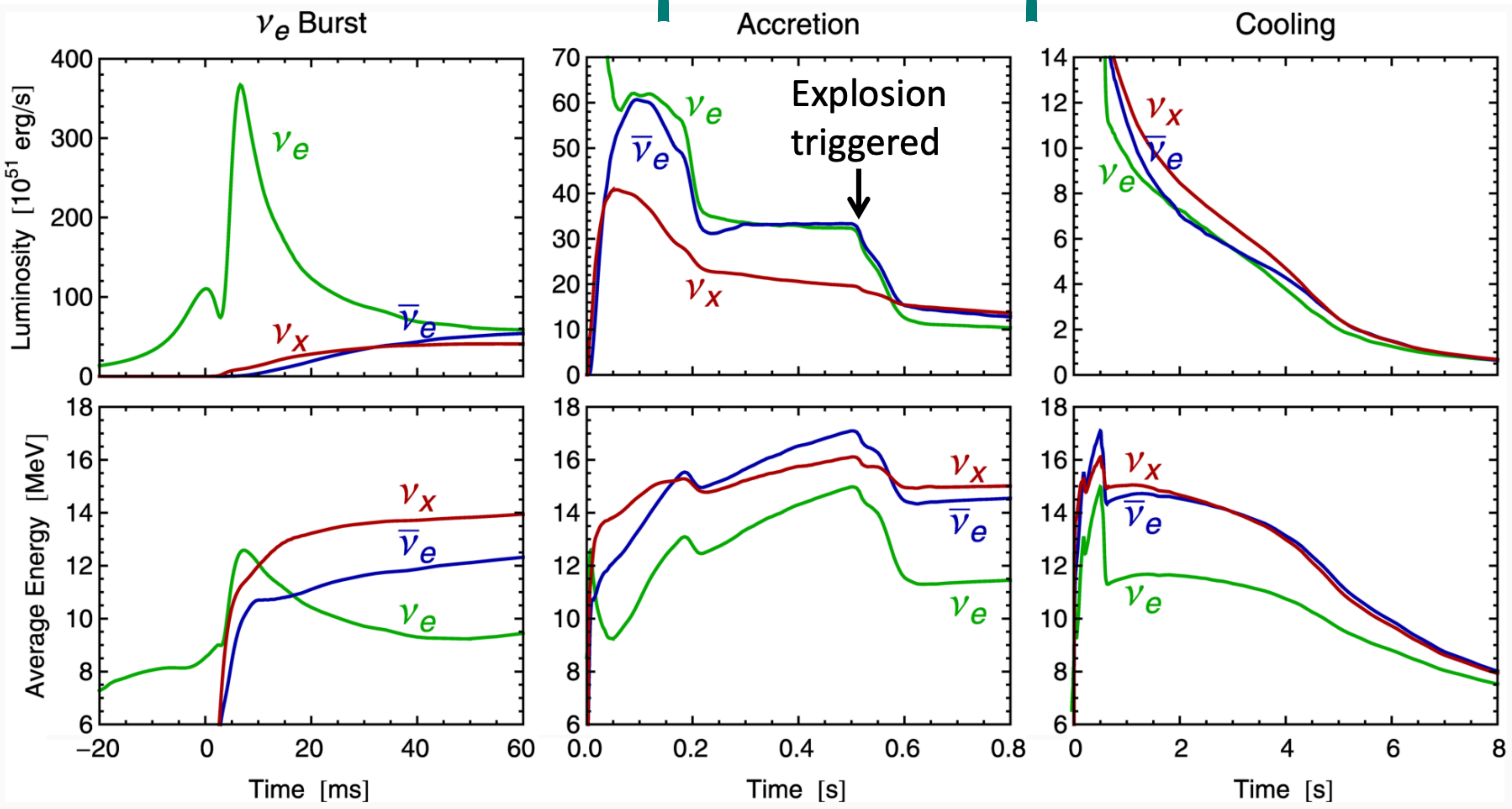


Core-Collapse Supernovae



- Shock breakout
- De-leptonization of outer core layers

- Shock stalls ~ 150 km
- Neutrinos powered by infalling matter

Cooling on neutrino diffusion time scale

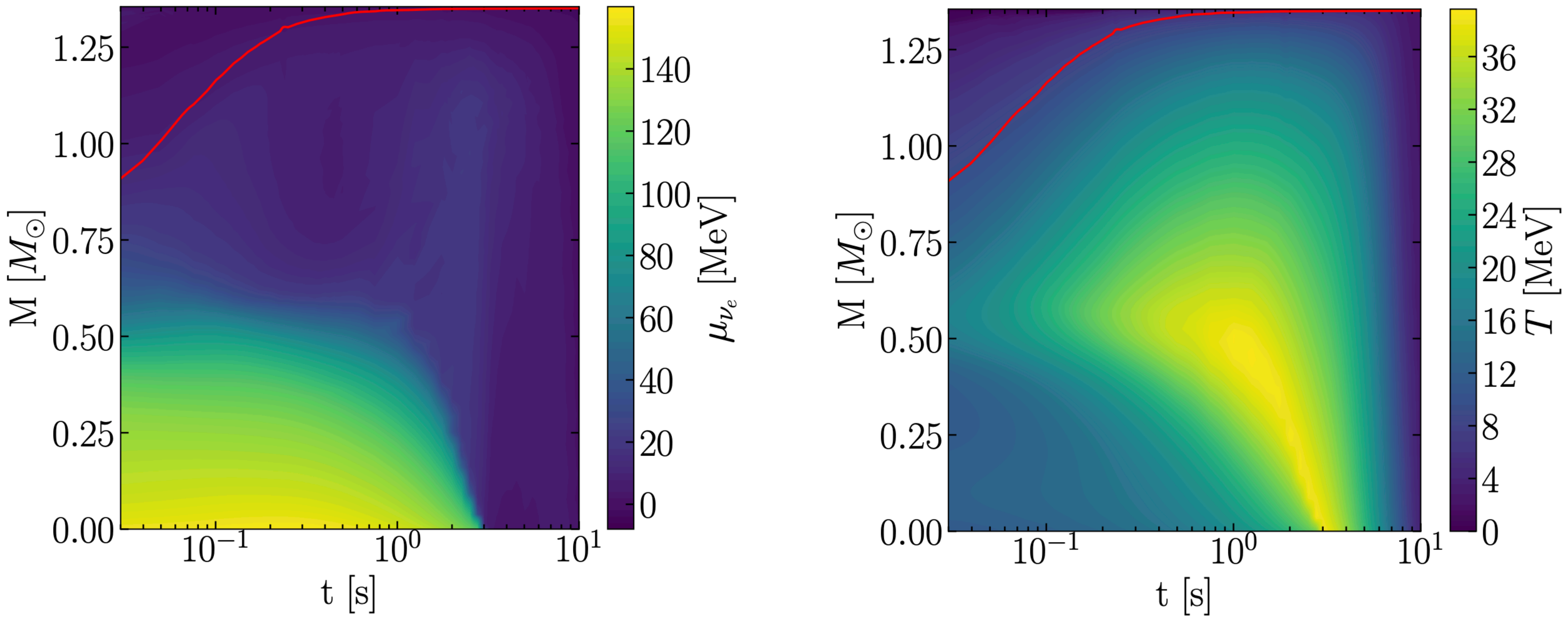
Credits to G. Raffelt

Spherically symmetric Garching model (25 $M_\odot$) with Boltzmann neutrino transport

Testing for new physics

New particles can be produced in supernova core…

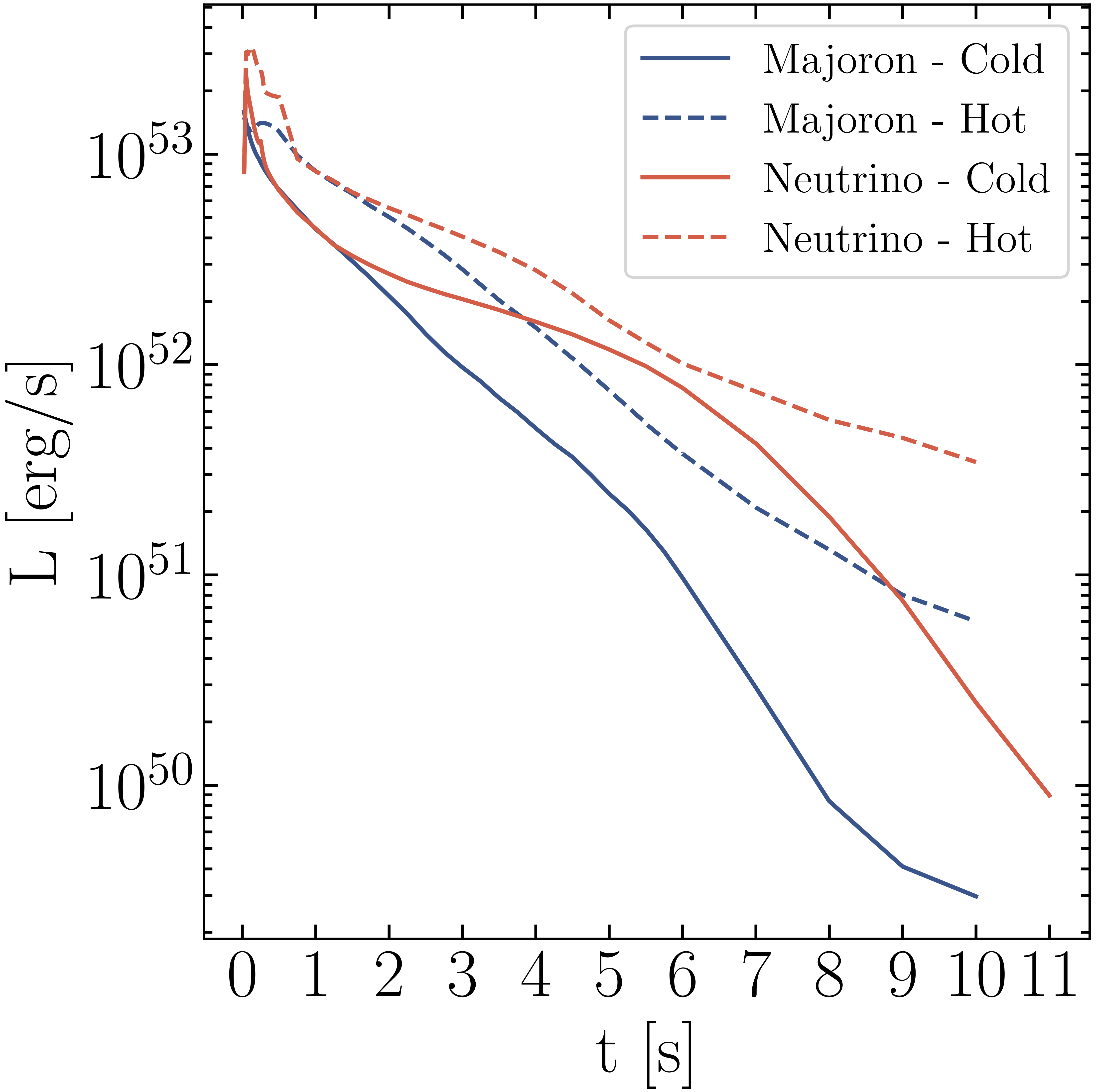
… but how do we probe them?



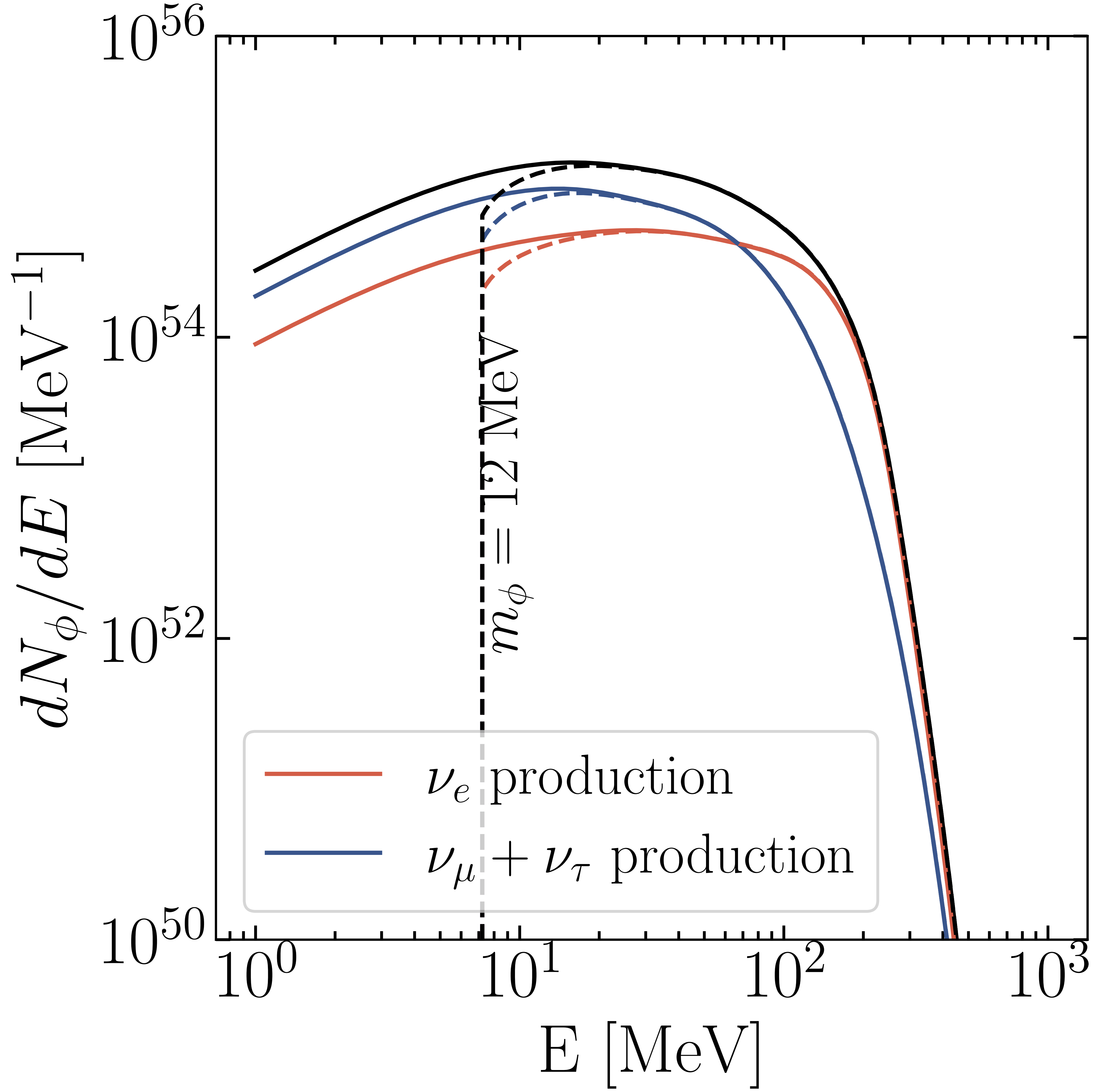
Majoron production

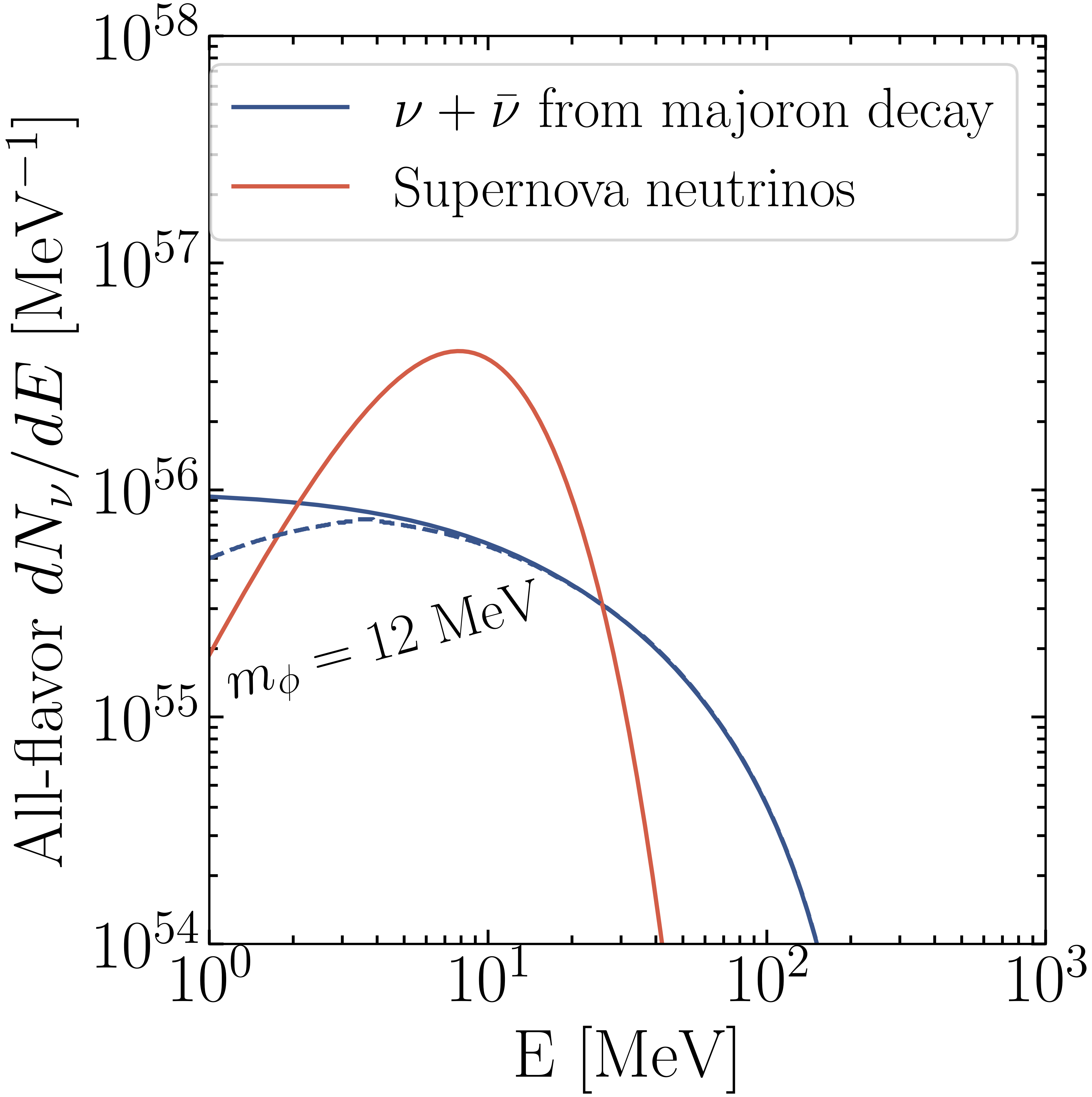
For small masses, signal depends only on gm_ϕ

$$\frac{dN_\phi}{dt} = \frac{(gm_\phi)^2 \mu_\nu^2}{192\pi^3}$$

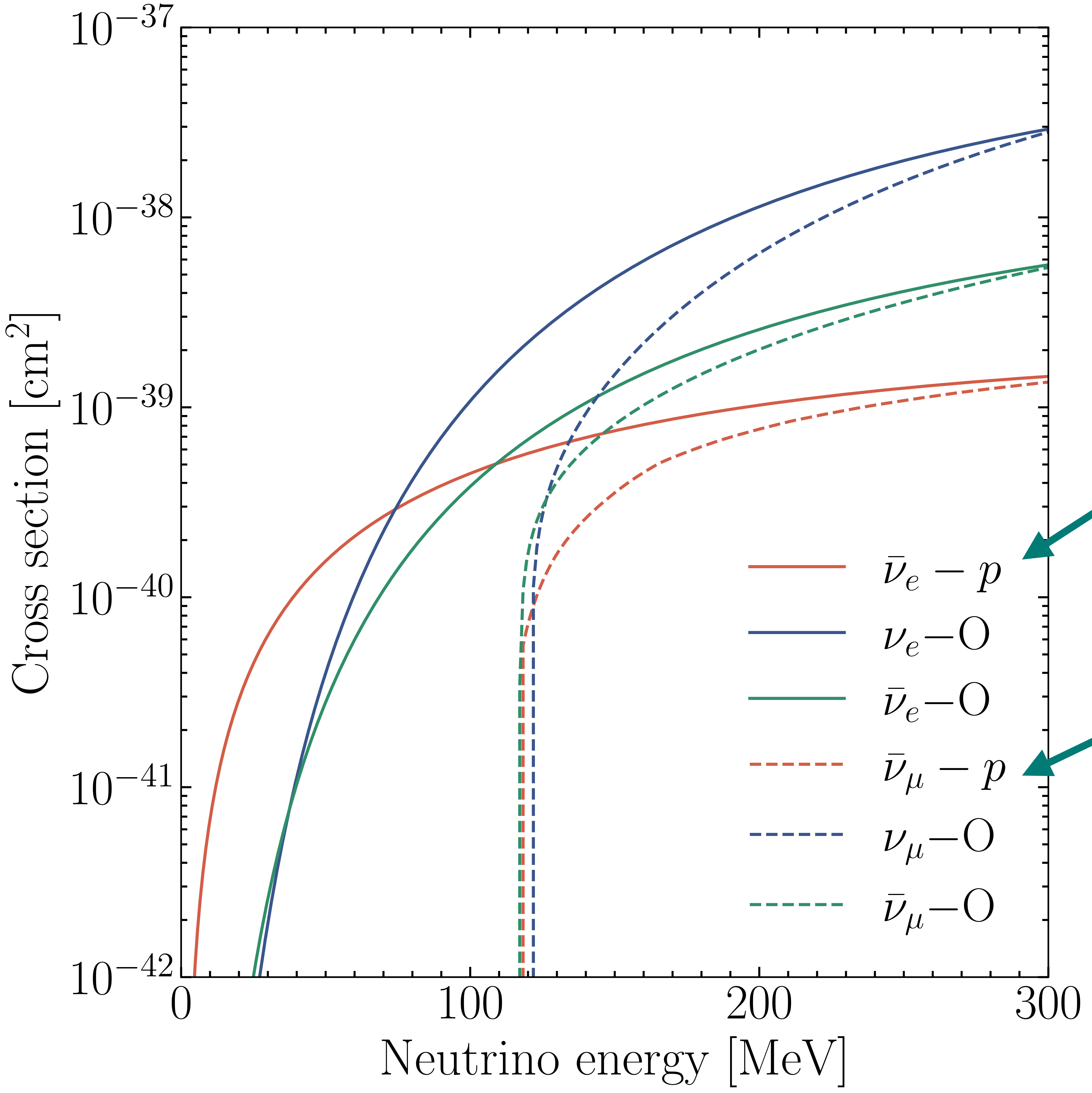


Majoron production





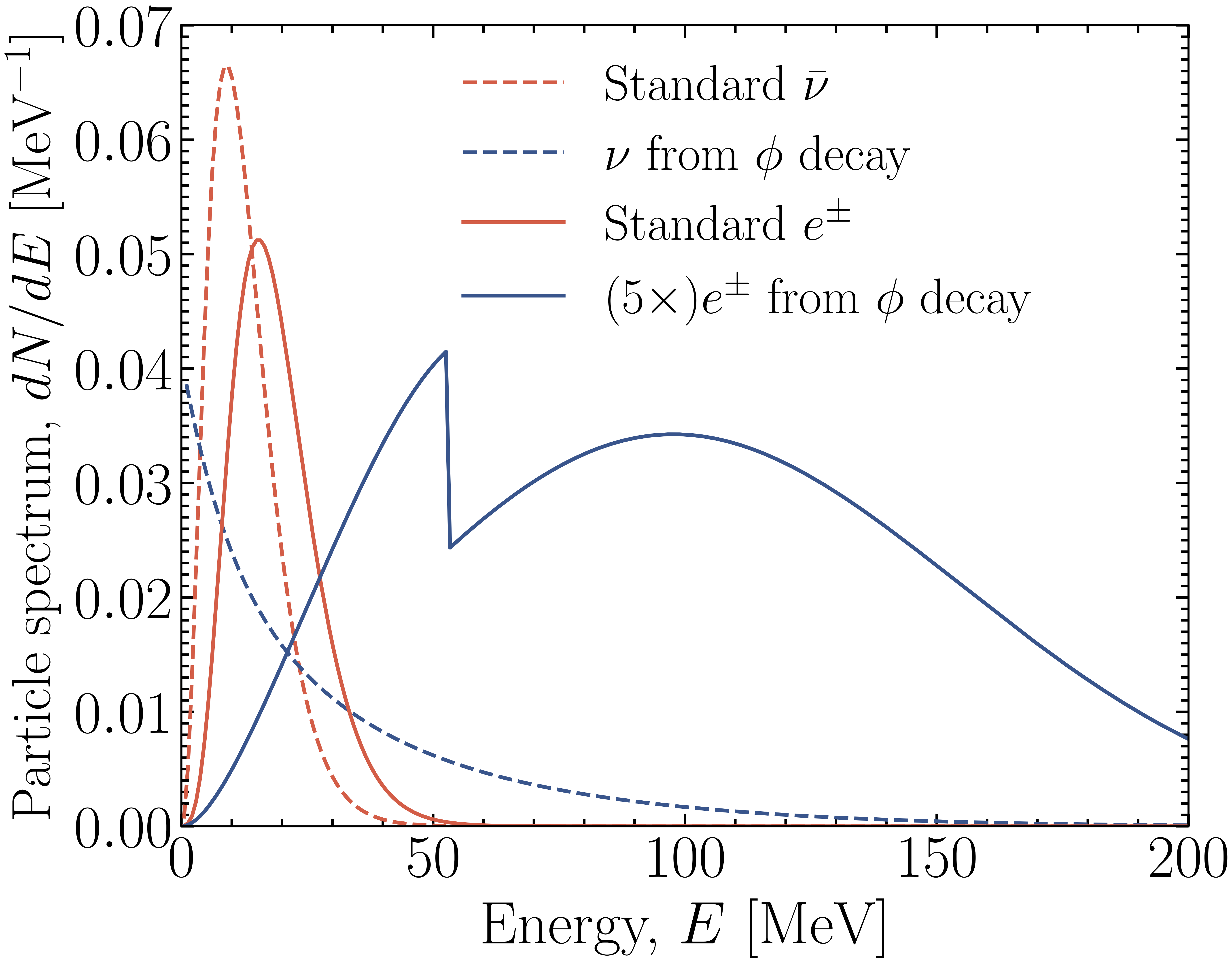
Neutrino detection



Appearing as $e^\pm$ Cherenkov signal

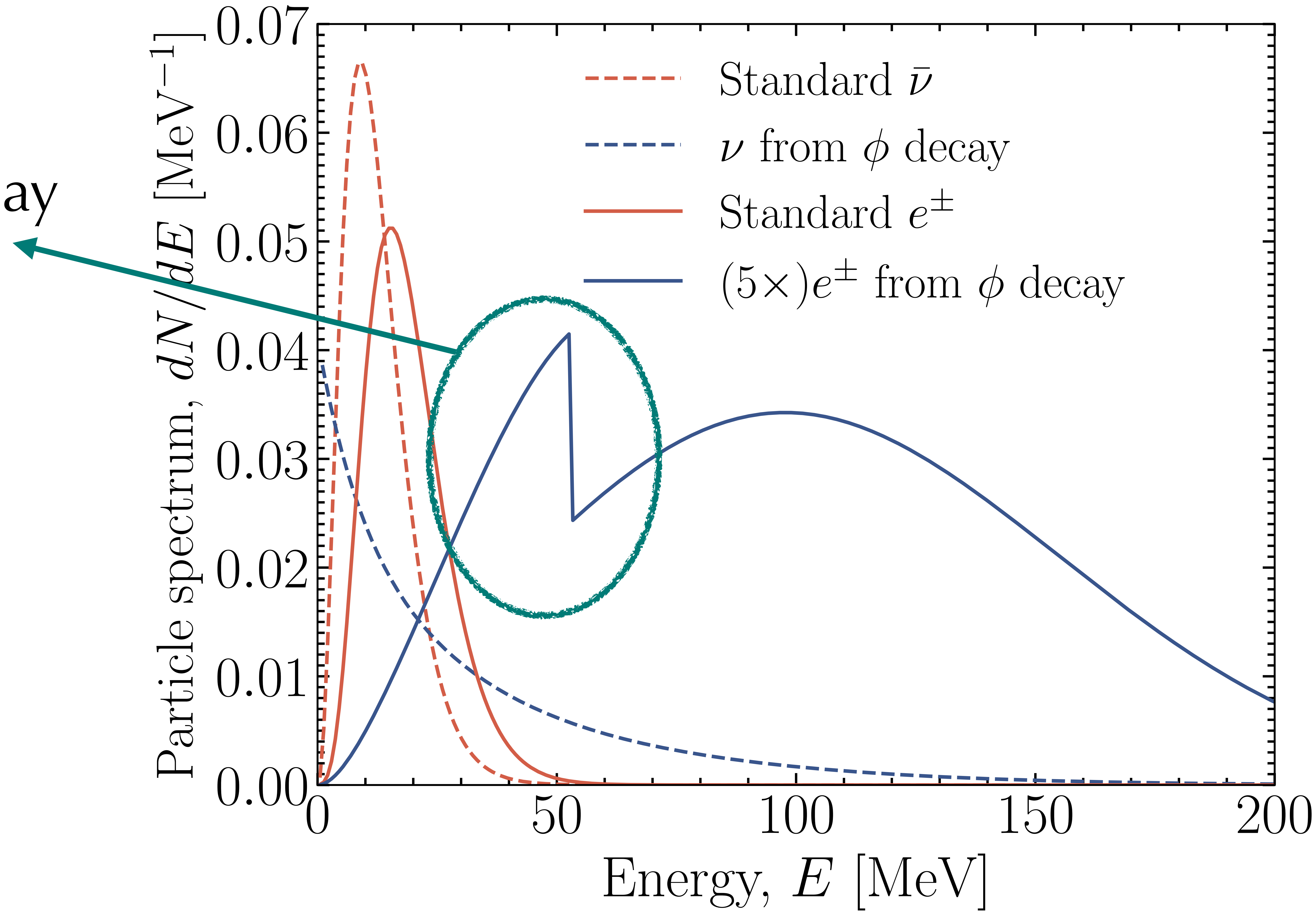
$\mu^\pm$ loses energy fast, appearing as $e^\pm$ from μ decay

Neutrino signal



Neutrino signal

From μ decay



No event observed above 40 MeV!

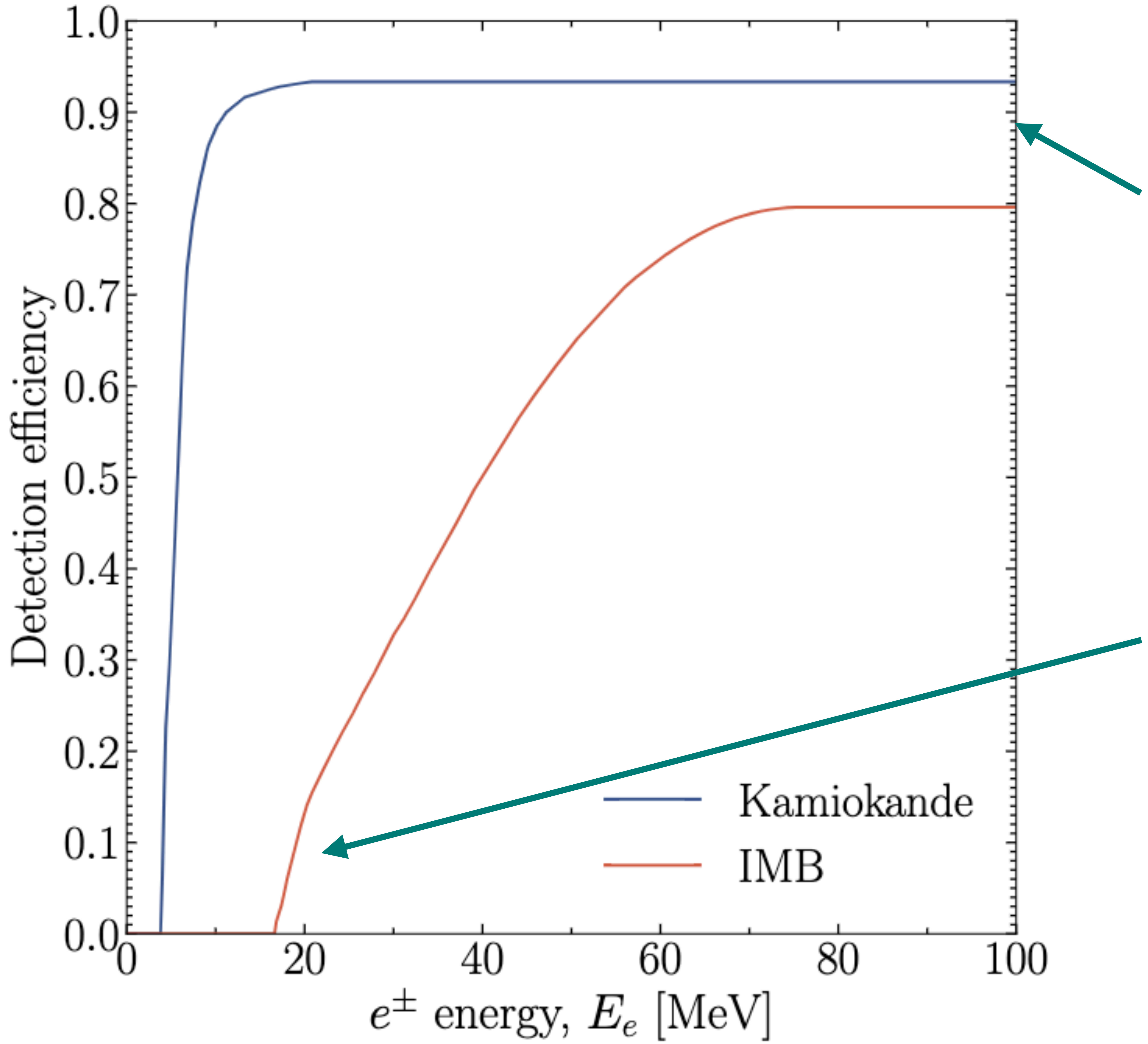
Statistical analysis

- Combined analysis of Kamiokande and IMB
- Supernova neutrino spectrum left as a fit parameter

$$\frac{dN_{\bar{\nu}_e}}{dE_\nu} = \frac{E_{\rm tot}}{6E_0^2}\frac{(1+\alpha)^{1+\alpha}}{\Gamma(1+\alpha)}\left(\frac{E_\nu}{E_0}\right)^\alpha e^{-(1+\alpha)E_\nu/E_0}$$

- Fit performed both for cold ($\alpha = 2.39$) and hot ($\alpha = 2.07$) model

Detection efficiency

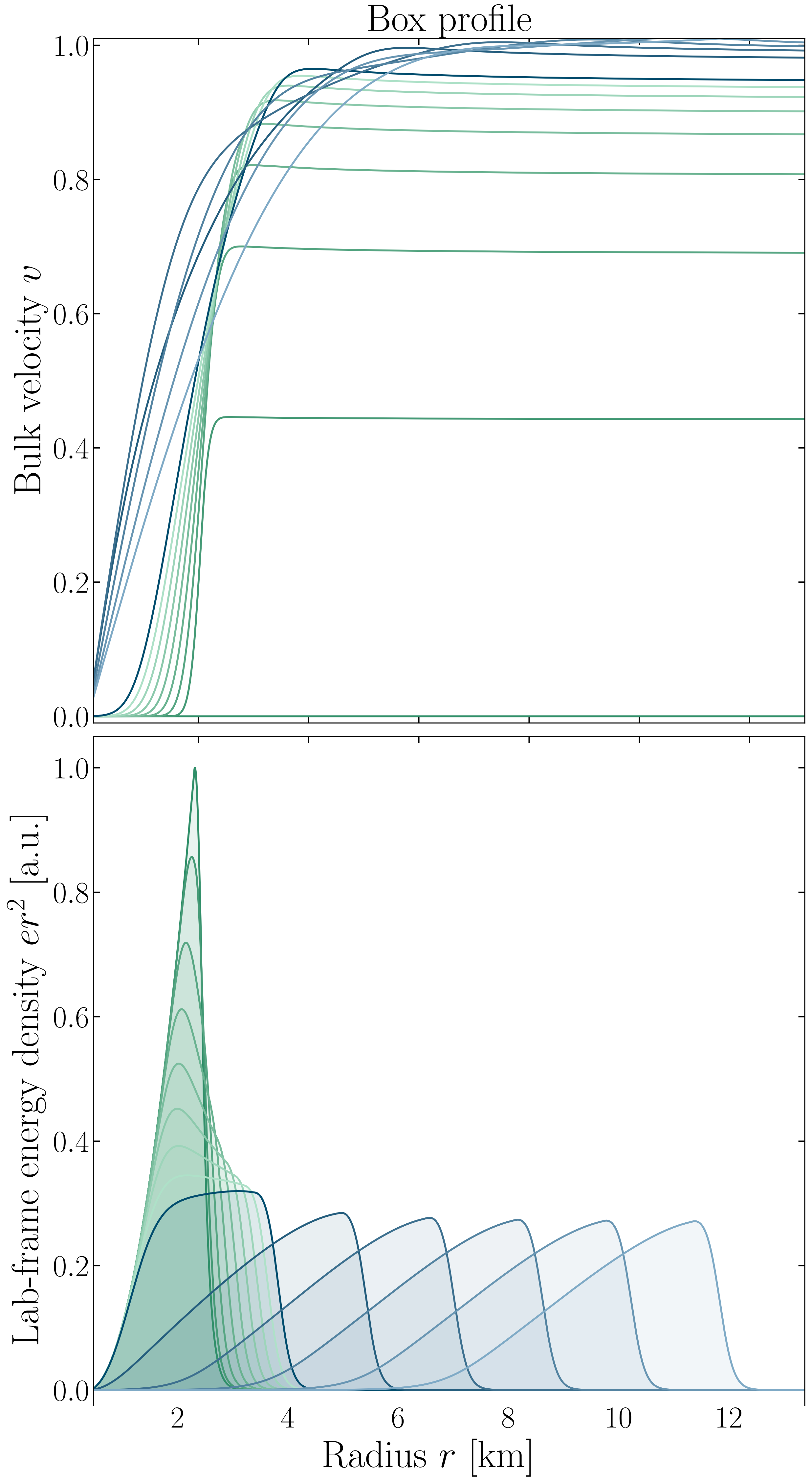


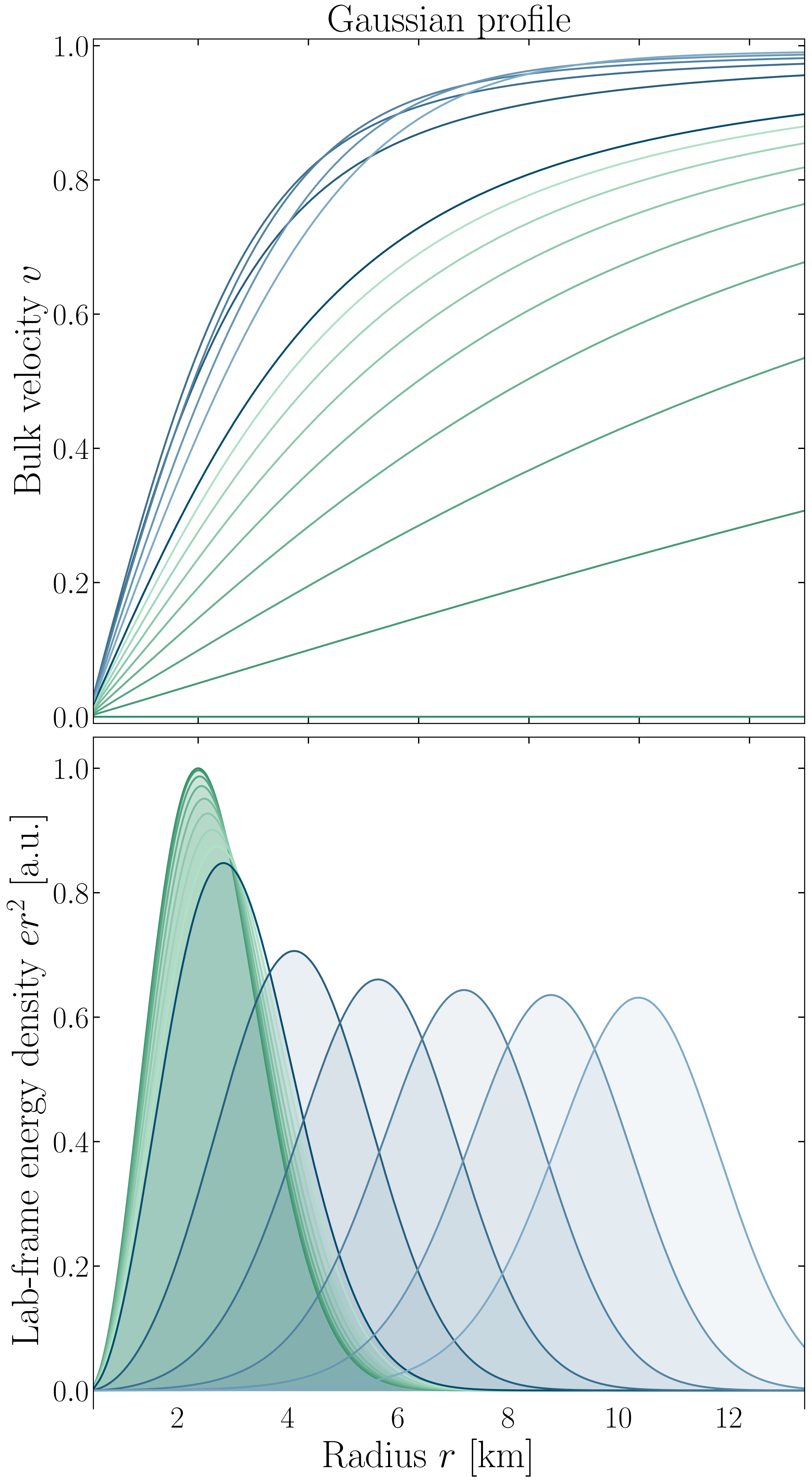
Comparable efficiencies at high energies, but IMB has larger volume

Few low-energy events at IMB

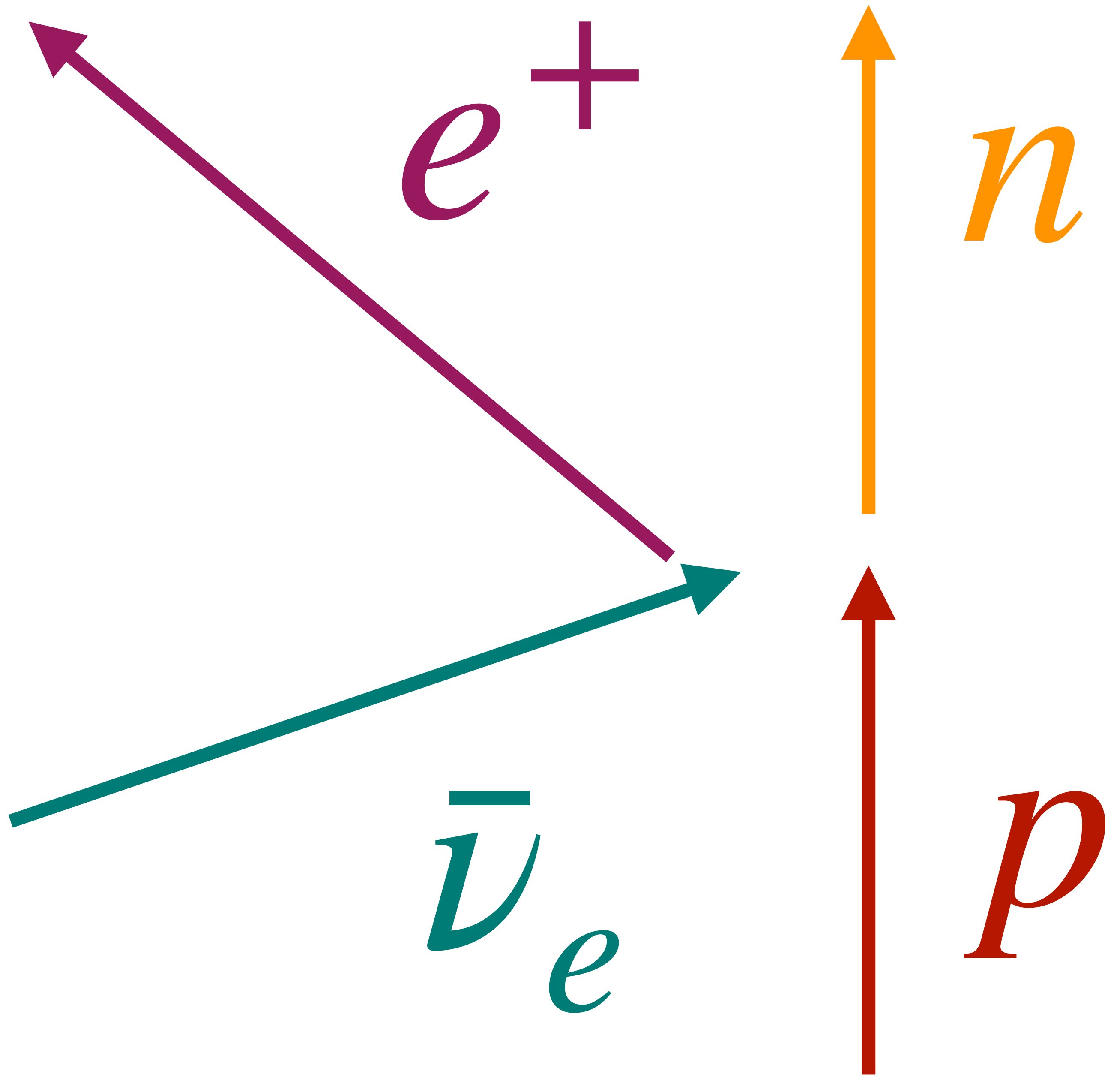
Above 75 MeV no events observed
(private communication)

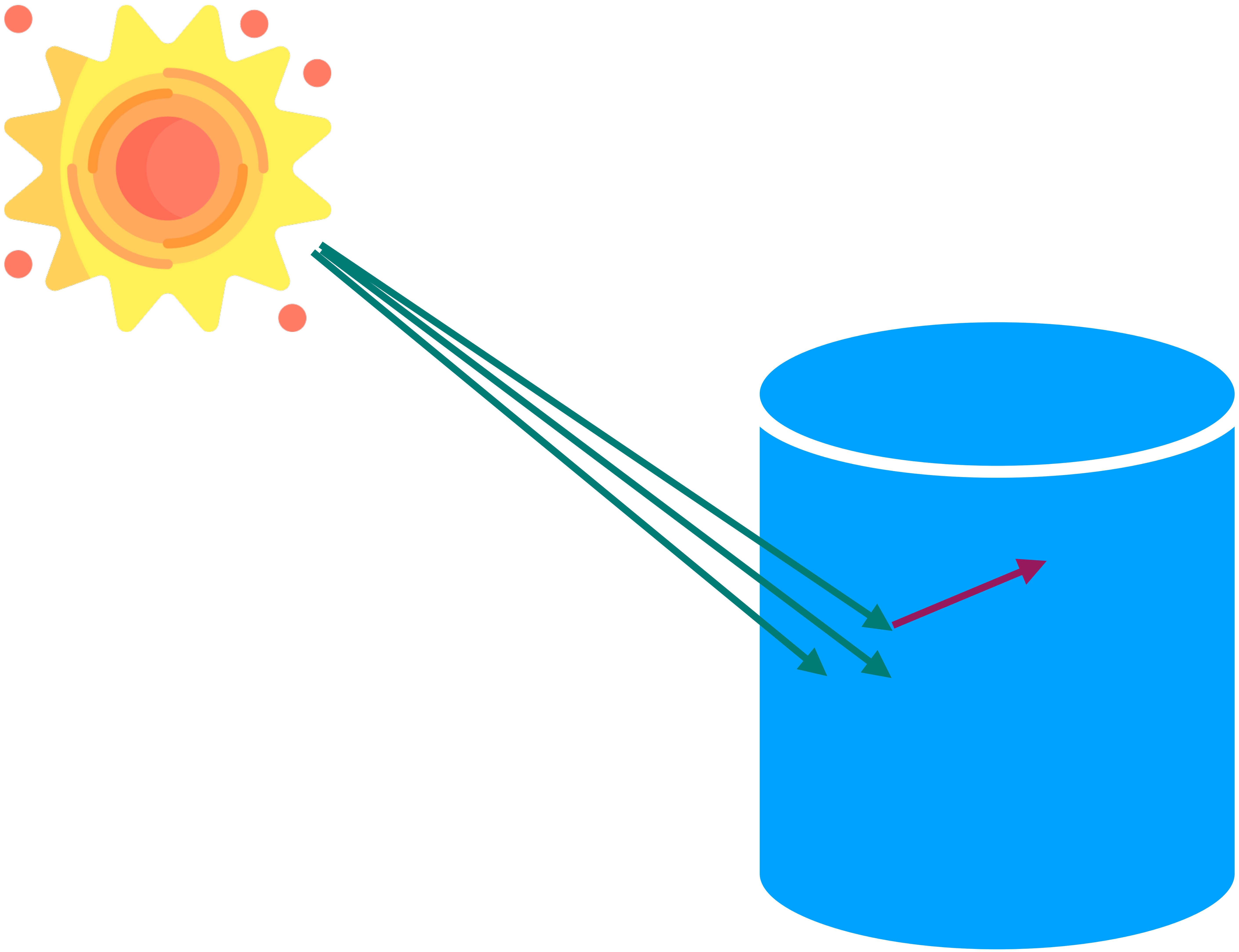
Fireball formation vs. burst outflow



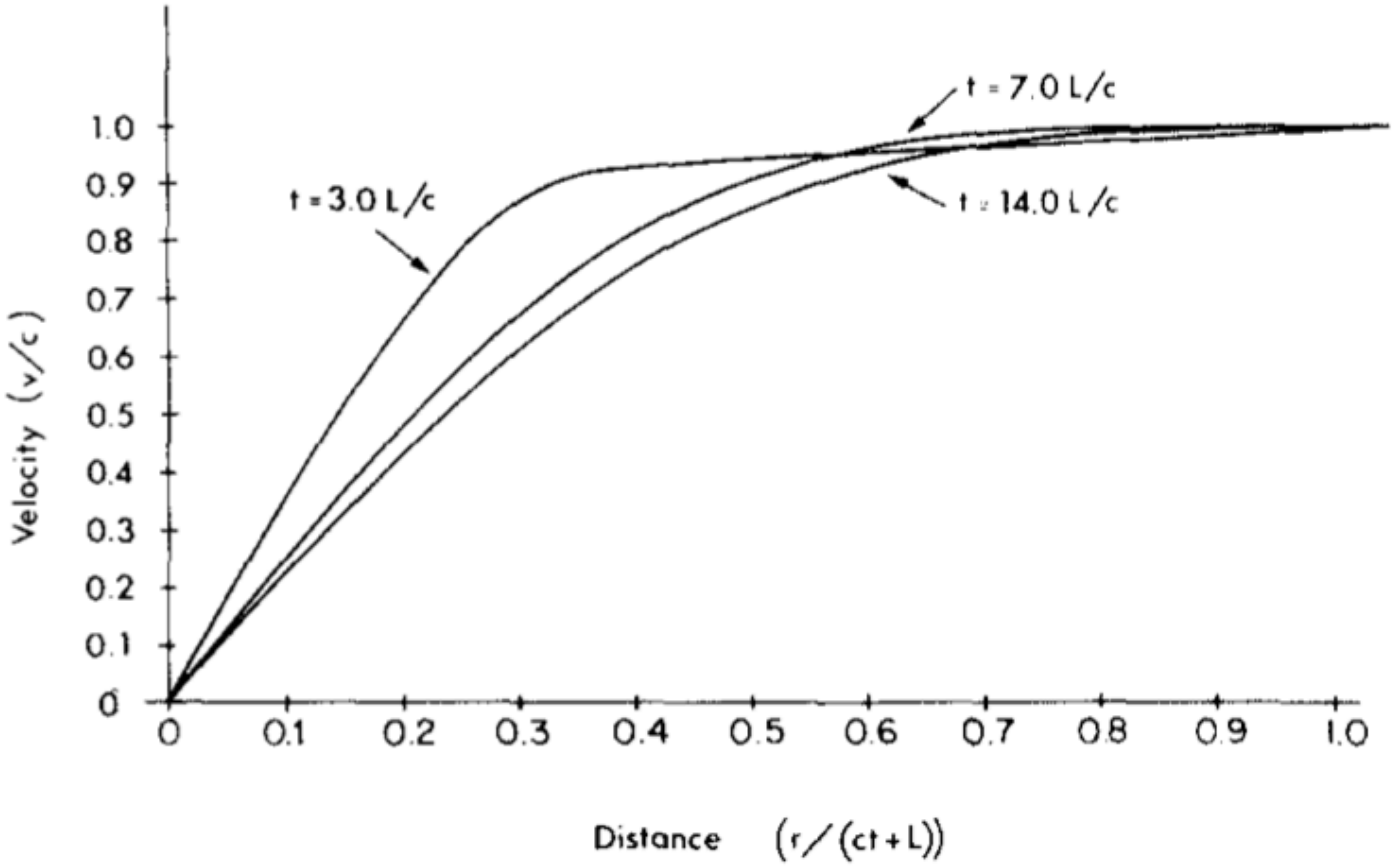


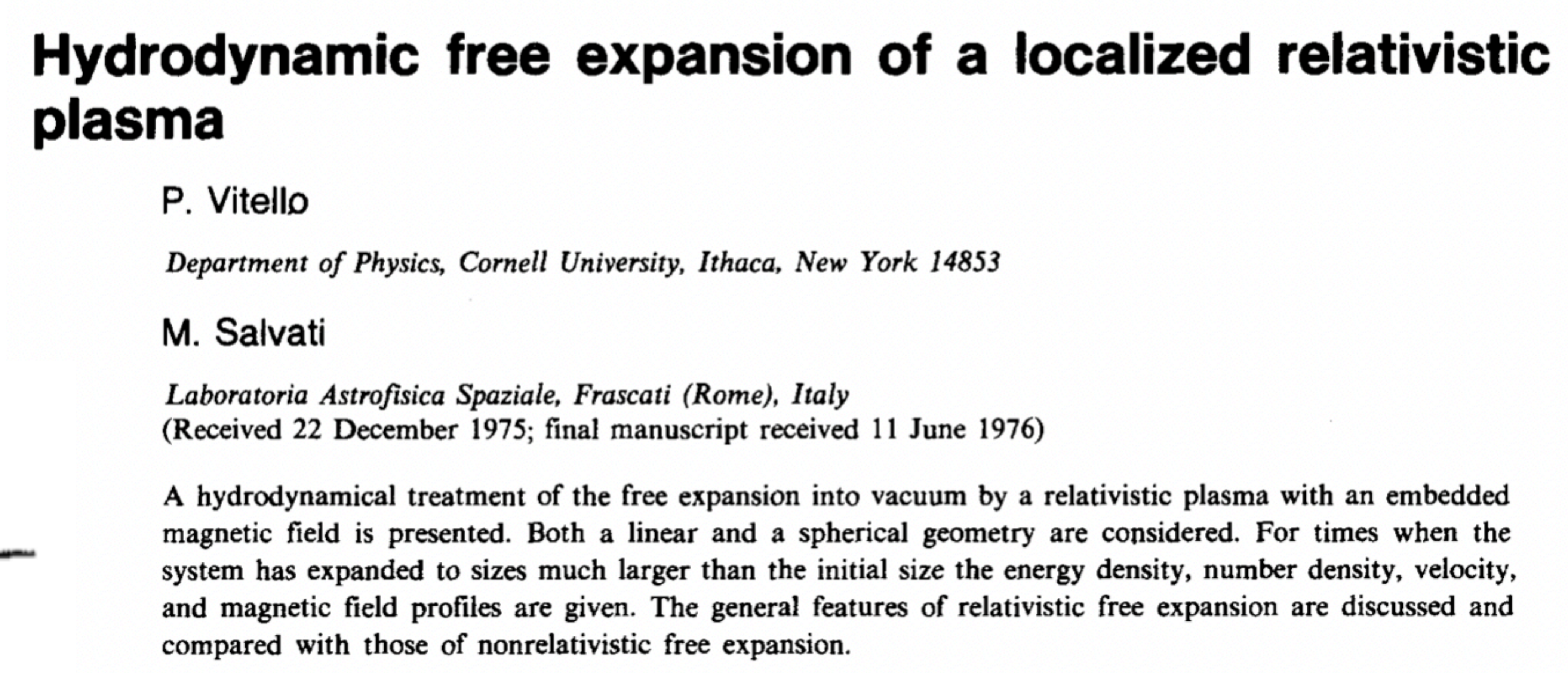
SN1987A neutrino observations





Fireball formation vs. burst outflow





Hydrodynamic free expansion of a localized relativistic plasma

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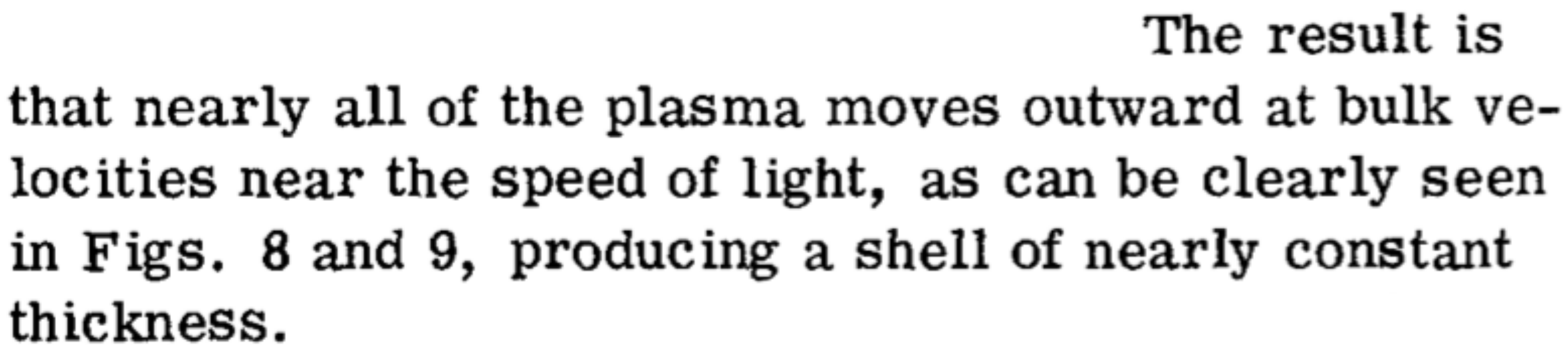
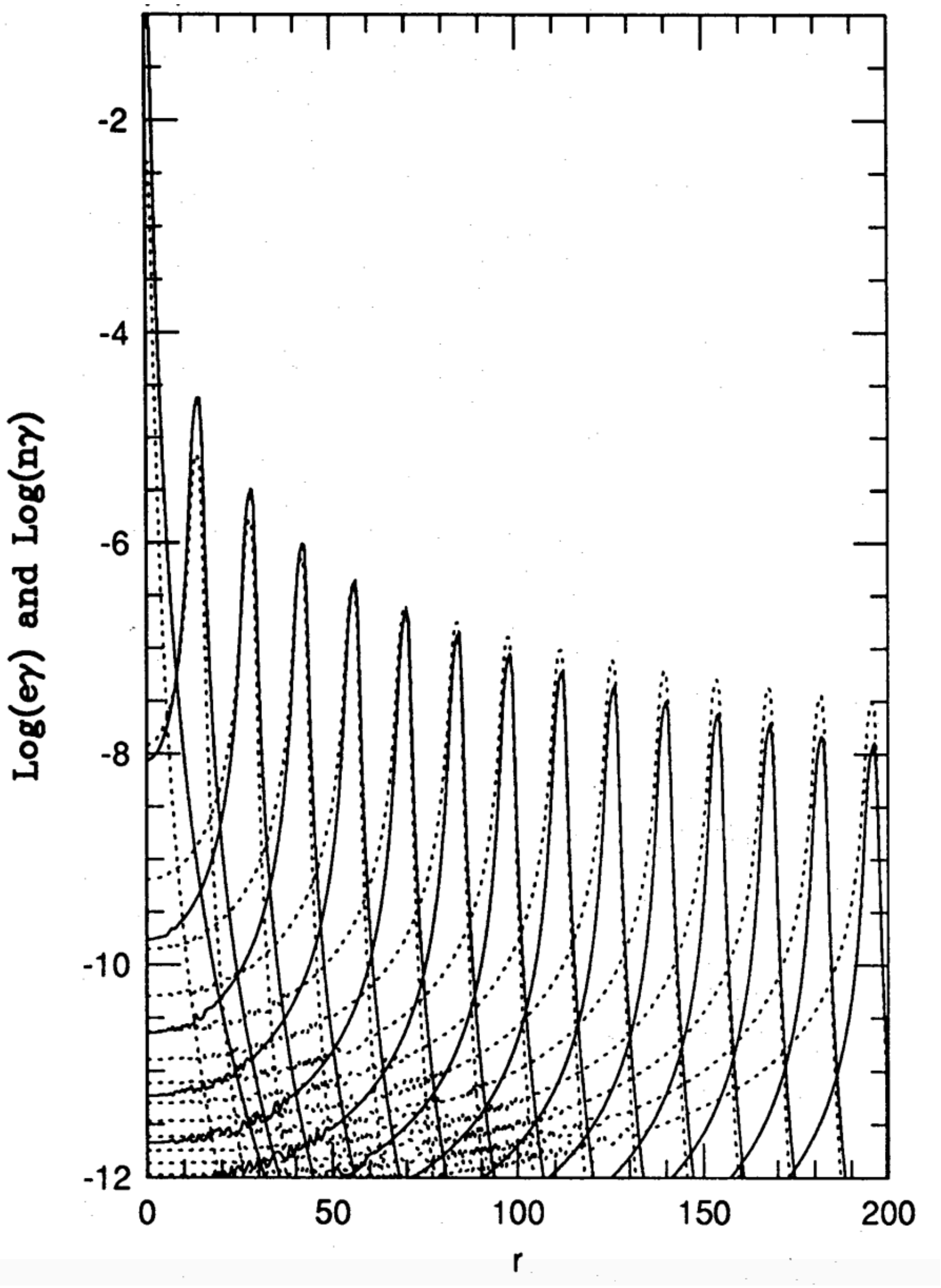
Laboratoria Astrofisica Spaziale, Frascati (Rome), Italy
(Received 22 December 1975; final manuscript received 11 June 1976)

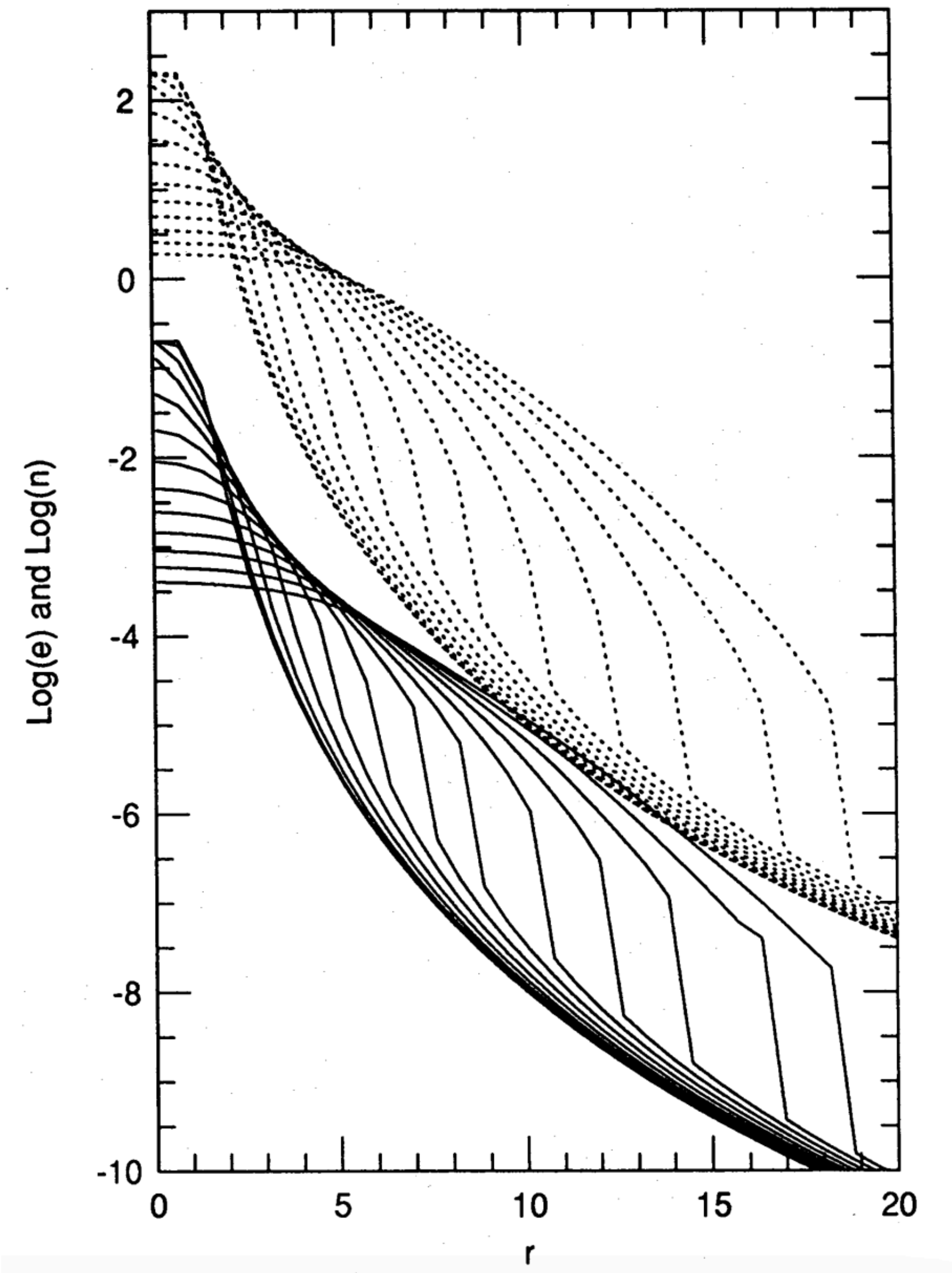
A hydrodynamical treatment of the free expansion into vacuum by a relativistic plasma with an embedded magnetic field is presented. Both a linear and a spherical geometry are considered. For times when the system has expanded to sizes much larger than the initial size the energy density, number density, velocity, and magnetic field profiles are given. The general features of relativistic free expansion are discussed and compared with those of nonrelativistic free expansion.



The result is that nearly all of the plasma moves outward at bulk velocities near the speed of light, as can be clearly seen in Figs. 8 and 9, producing a shell of nearly constant thickness.

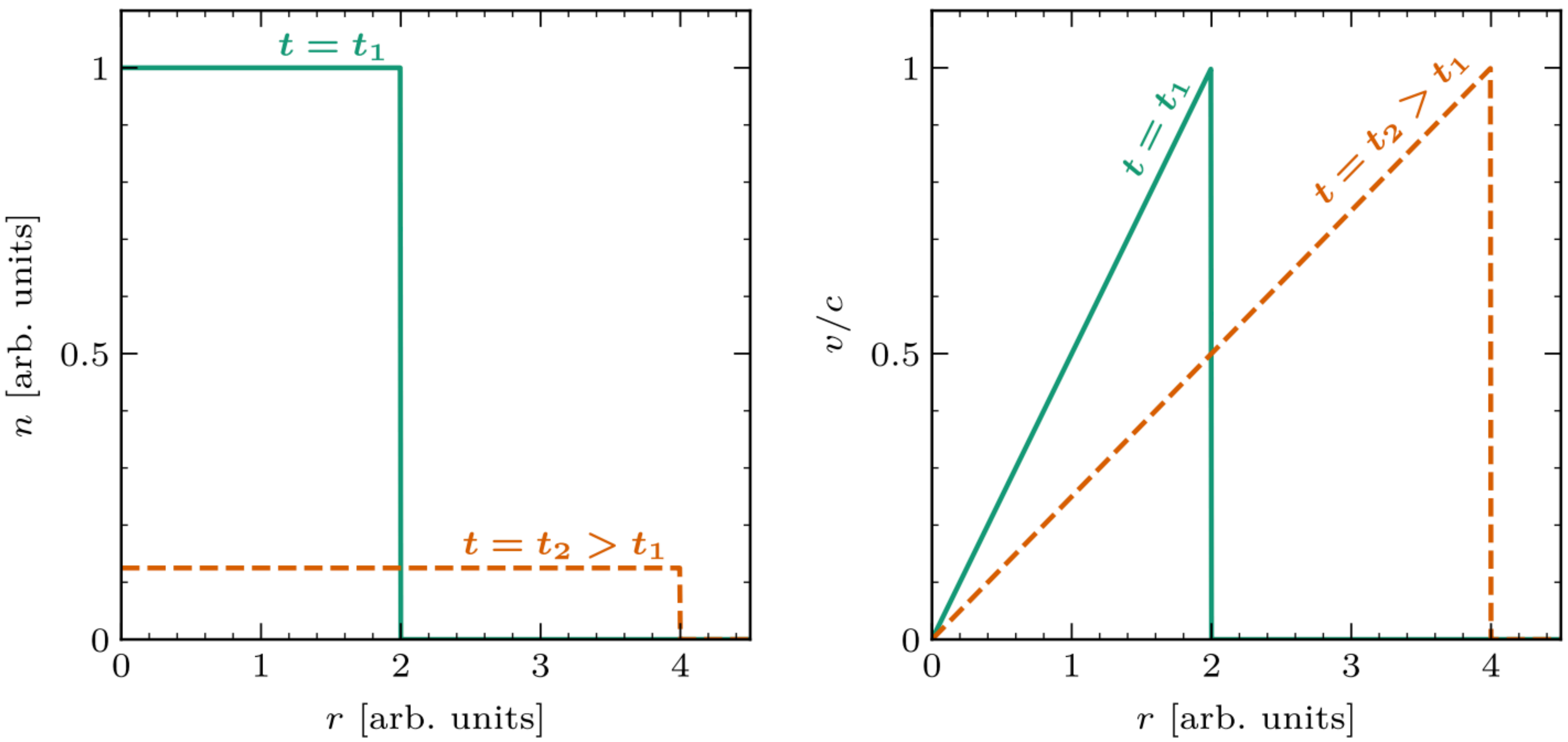
Fireball formation vs. burst outflow





Piran et al., arXiv:9301004

Burst outflow - analytical solution



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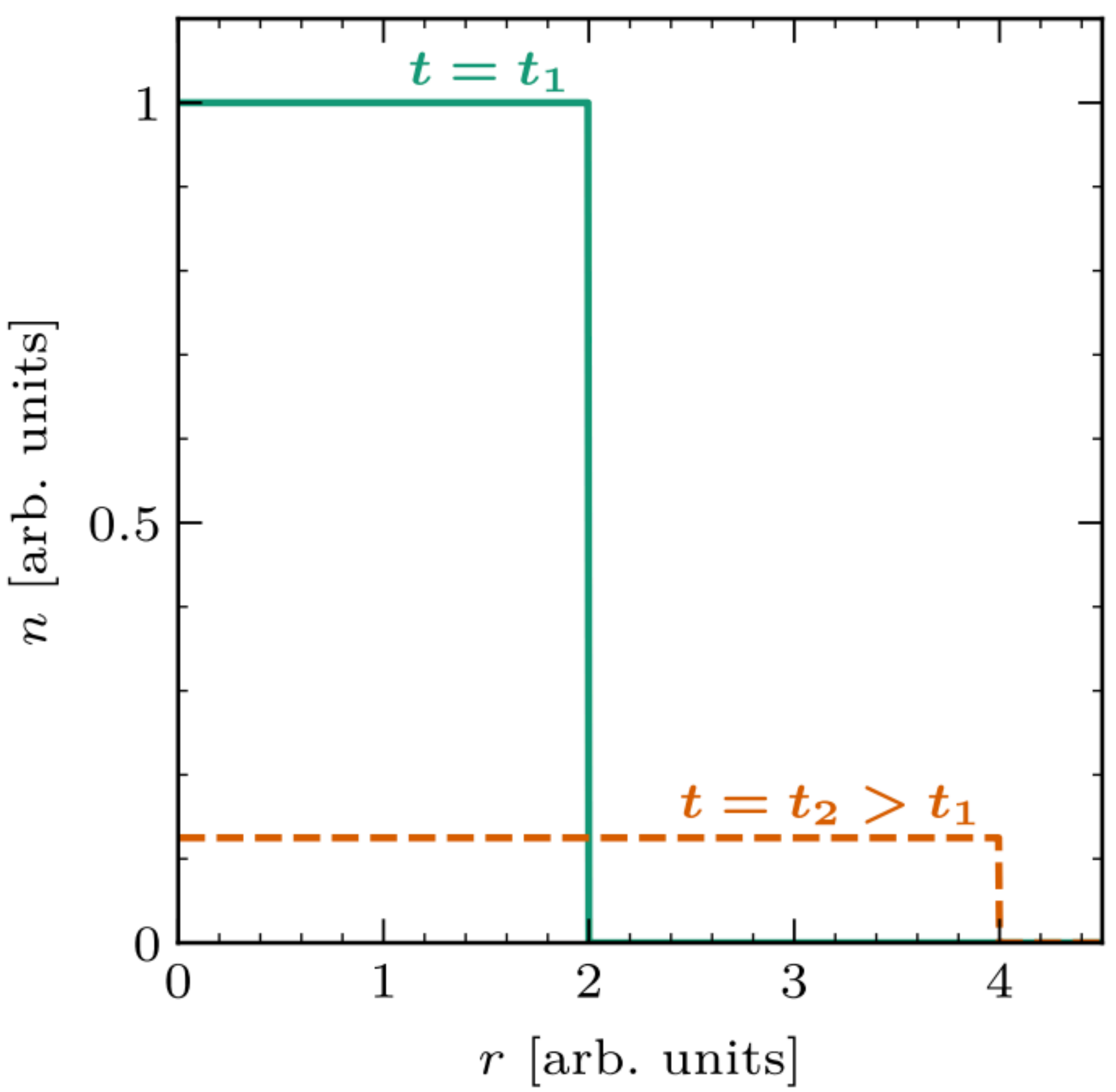
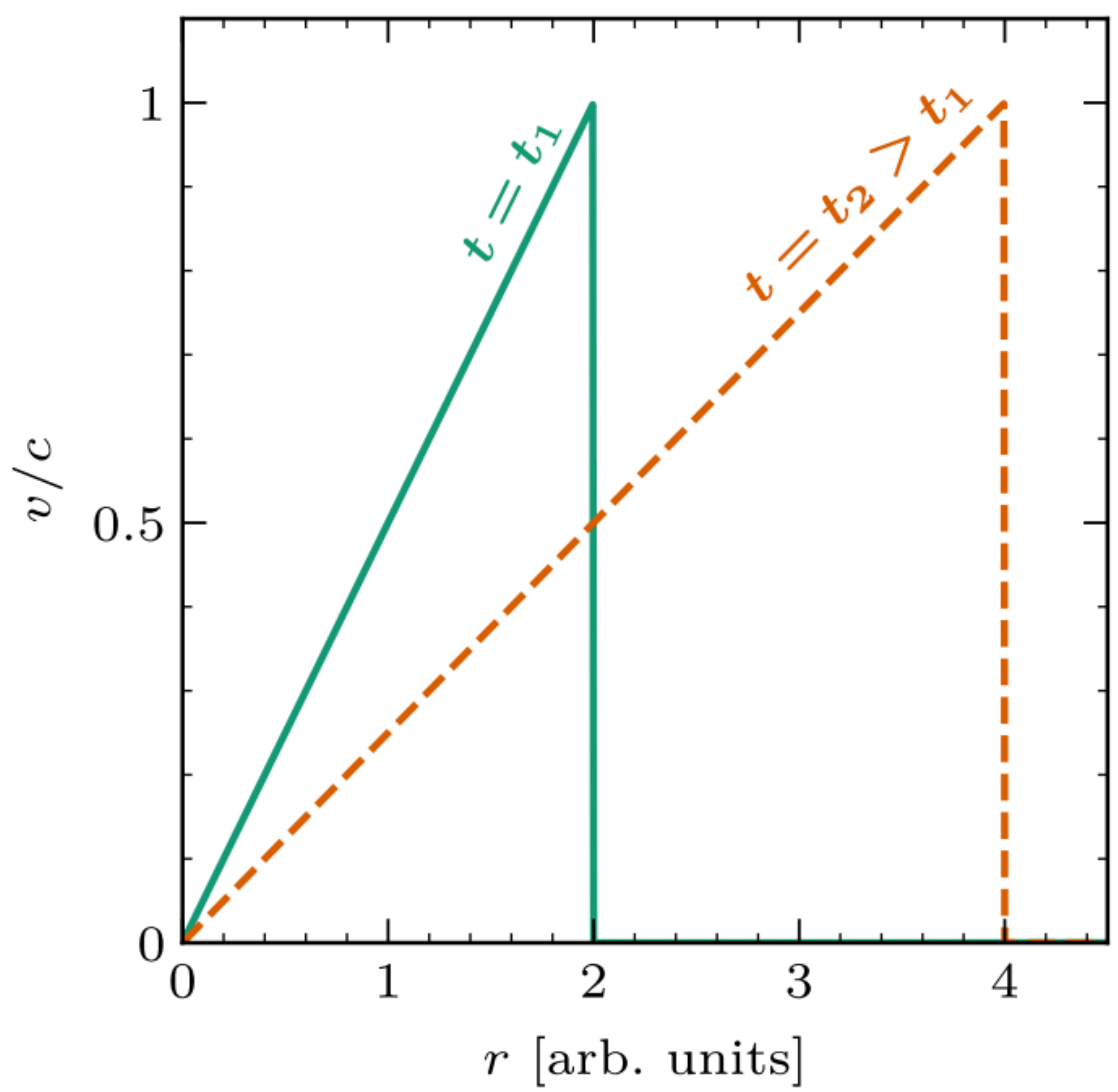
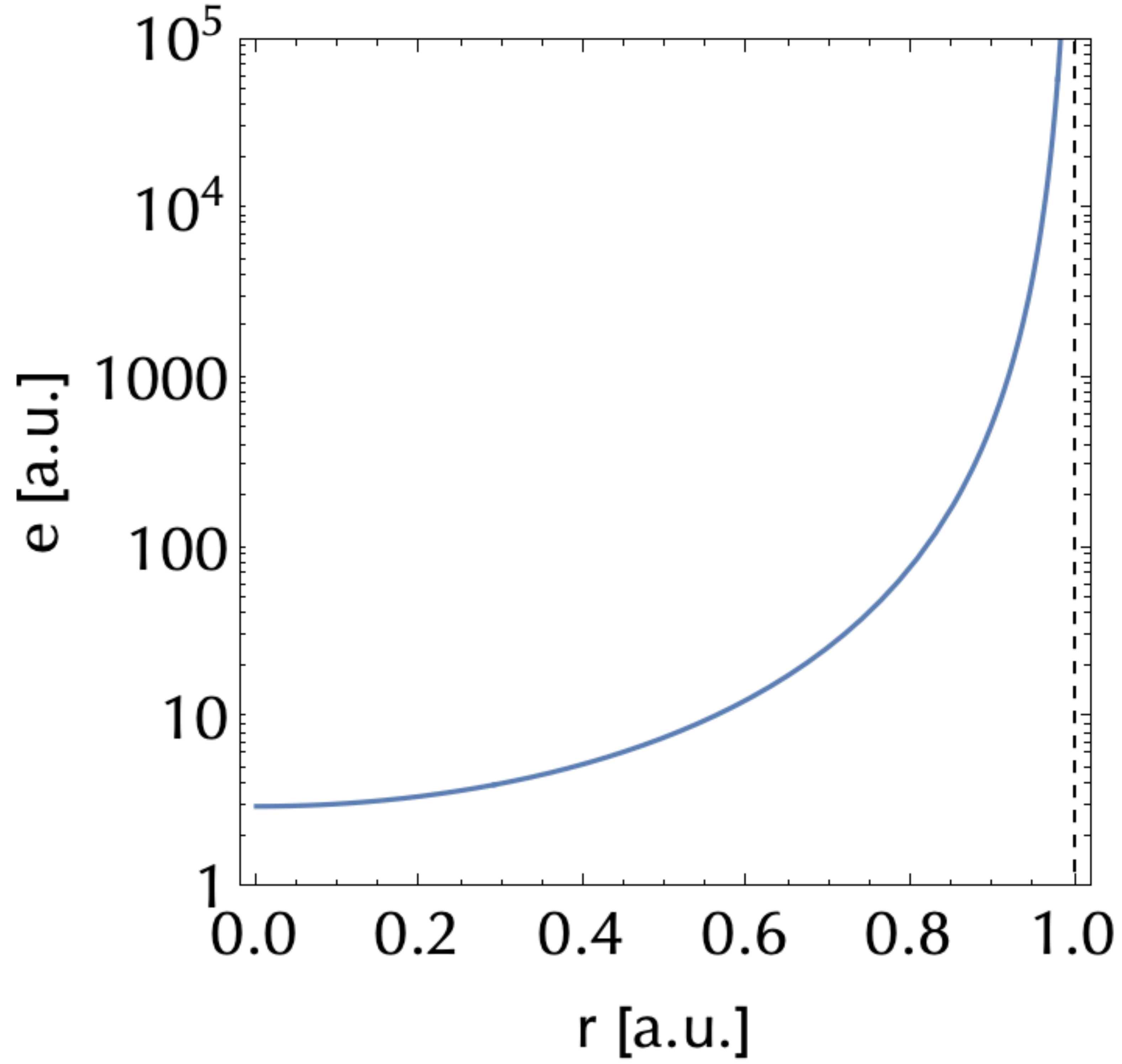
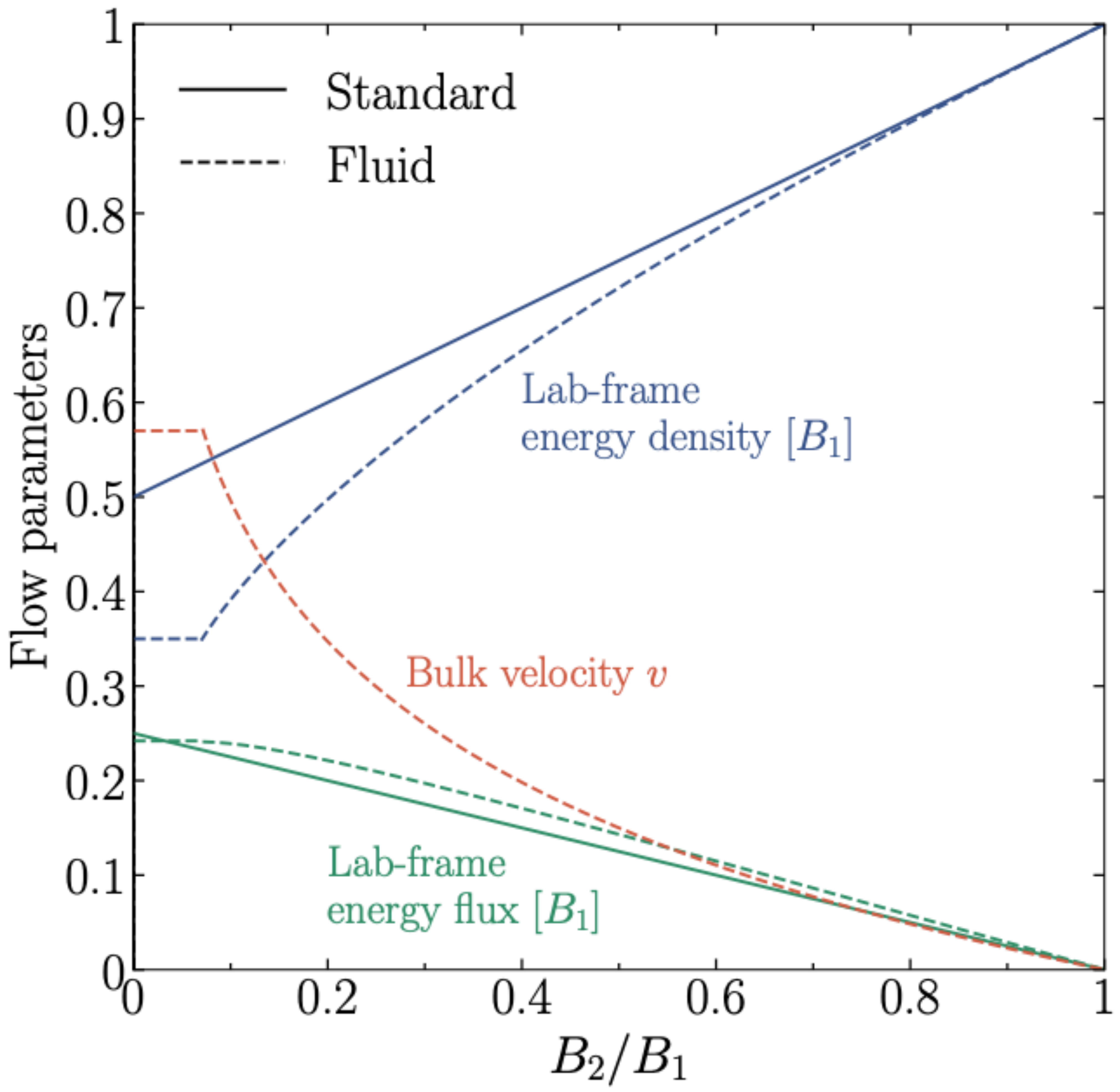




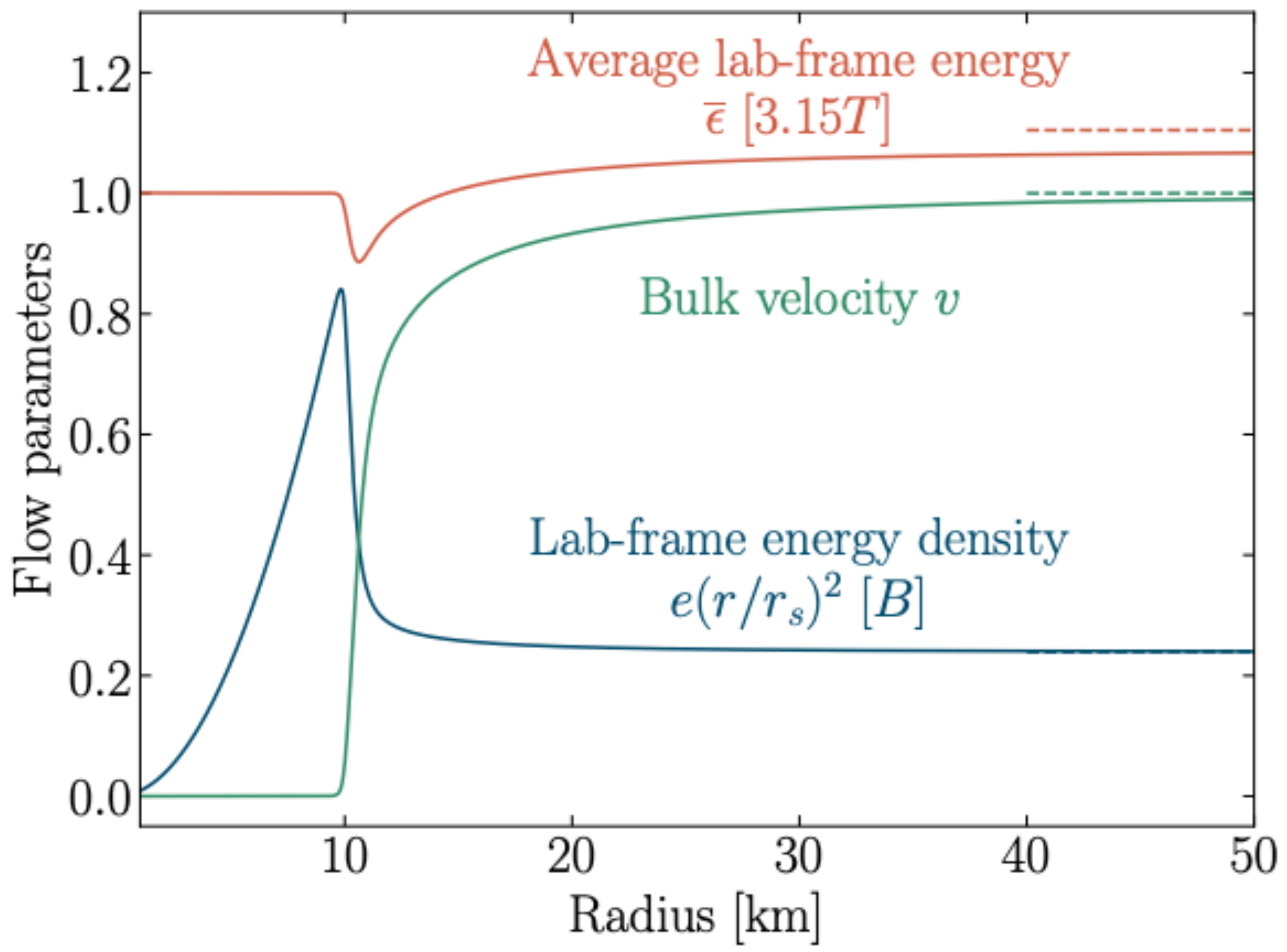


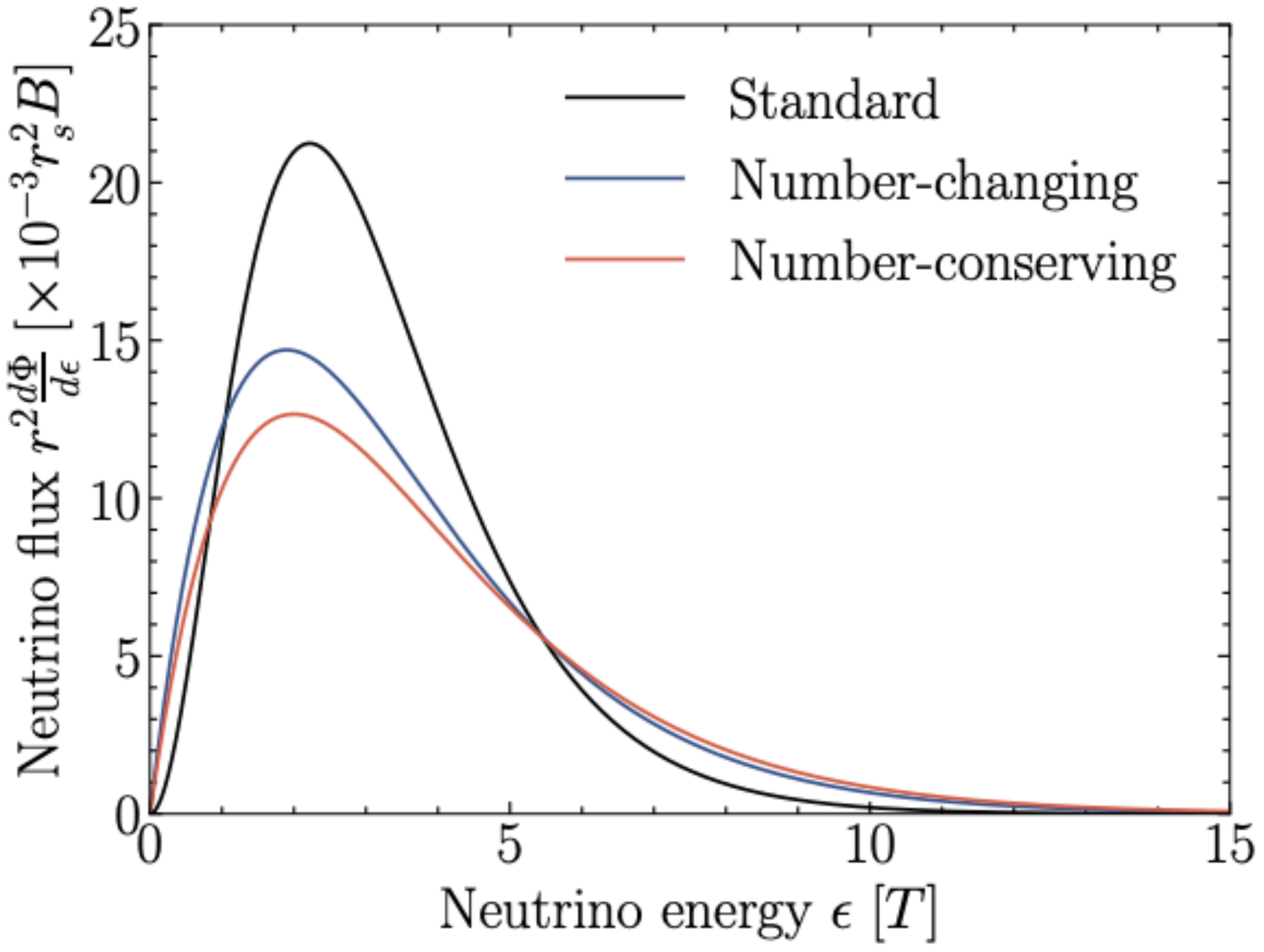
Cooling time with νSI



- Energy flux emitted is nearly identical to the standard case
- Fluid flows with speed of sound
- Energy density is somewhat higher than in the standard case

Spectrum with νSI





Novel bounds from SN1987A

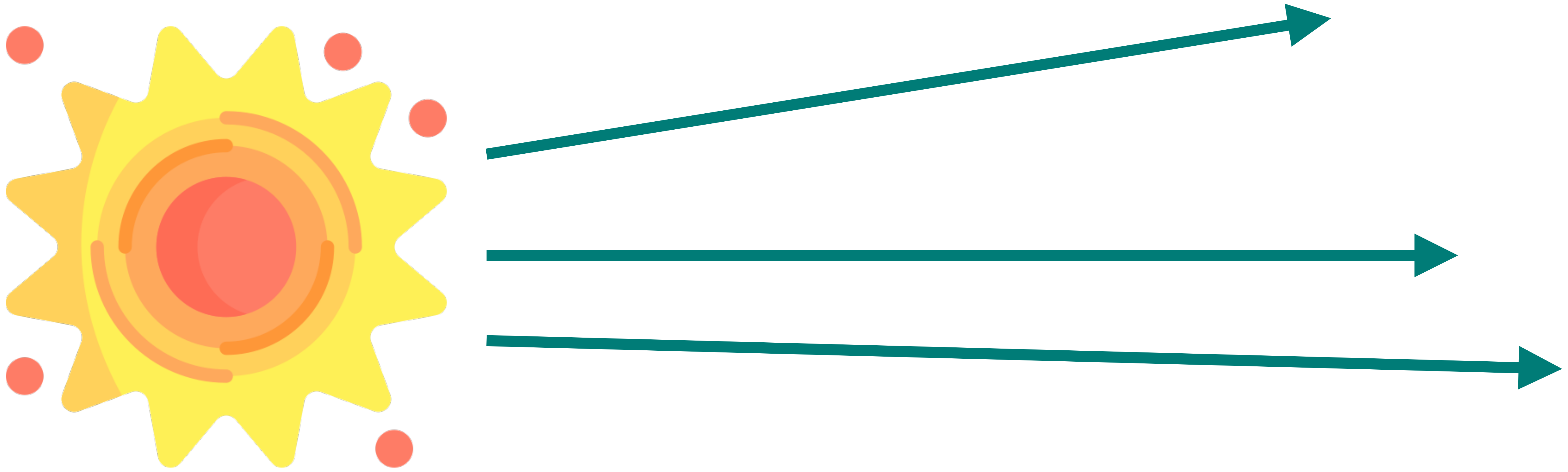
based on arXiv:2209.11773 (Phys. Rev. Lett. 131 2, 021001)

with G. Raffelt, E. Vitagliano

Testing for new physics

New particles can be produced in supernova core…

… but how do we probe them?

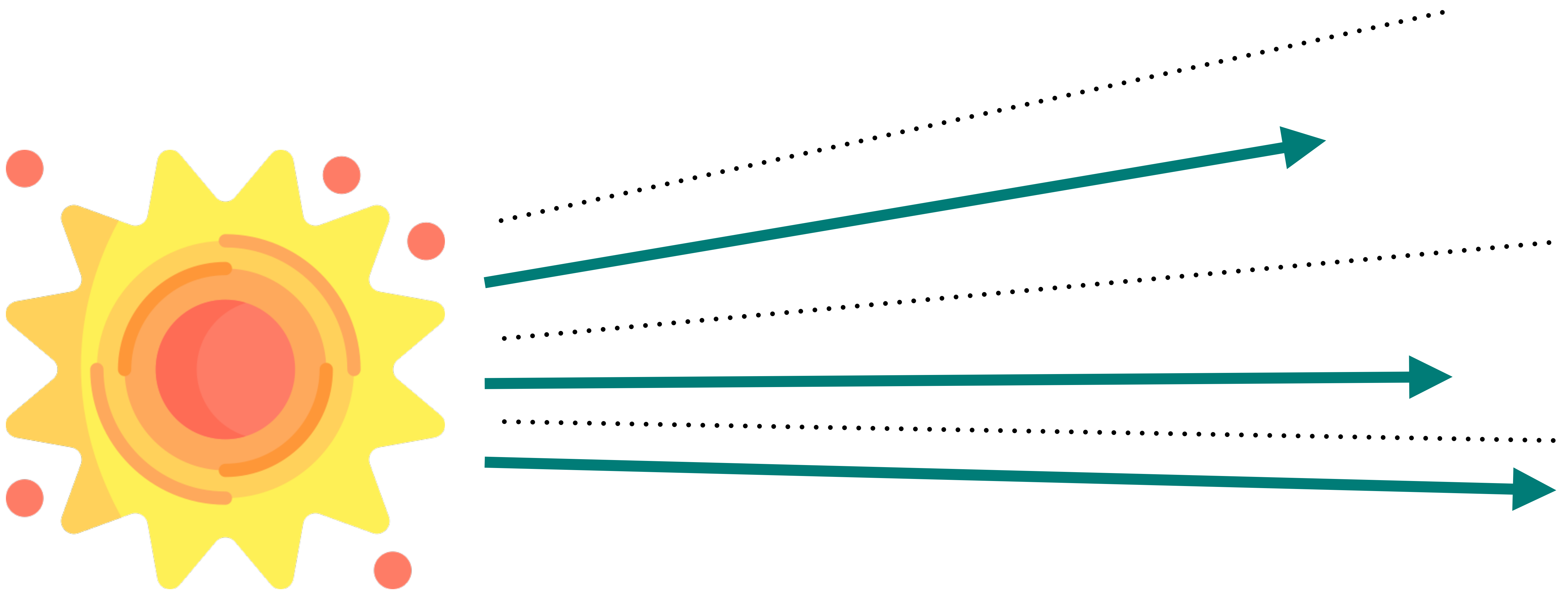


$$T \simeq \frac{R^2}{\lambda_{\nu N}} \simeq 1 - 10 \text{ s}$$

Testing for new physics

New particles can be produced in supernova core…

… but how do we probe them?



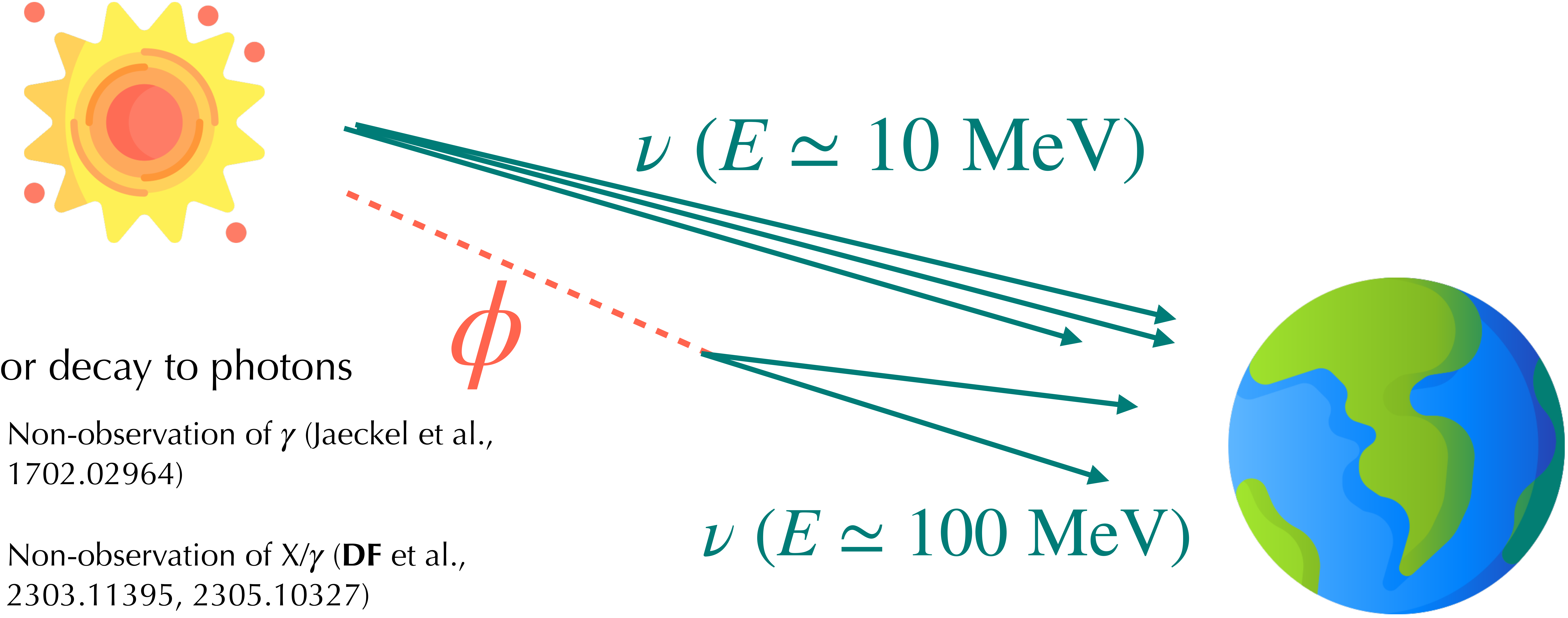
PNS cools faster,
shorter burst

Energy loss/cooling bound

$$L_\phi(1\ \mathrm{s}) < L_\nu(1\ \mathrm{s})$$

Burrows et al., 1989; Raffelt, 1996

Novel bounds from SN1987A



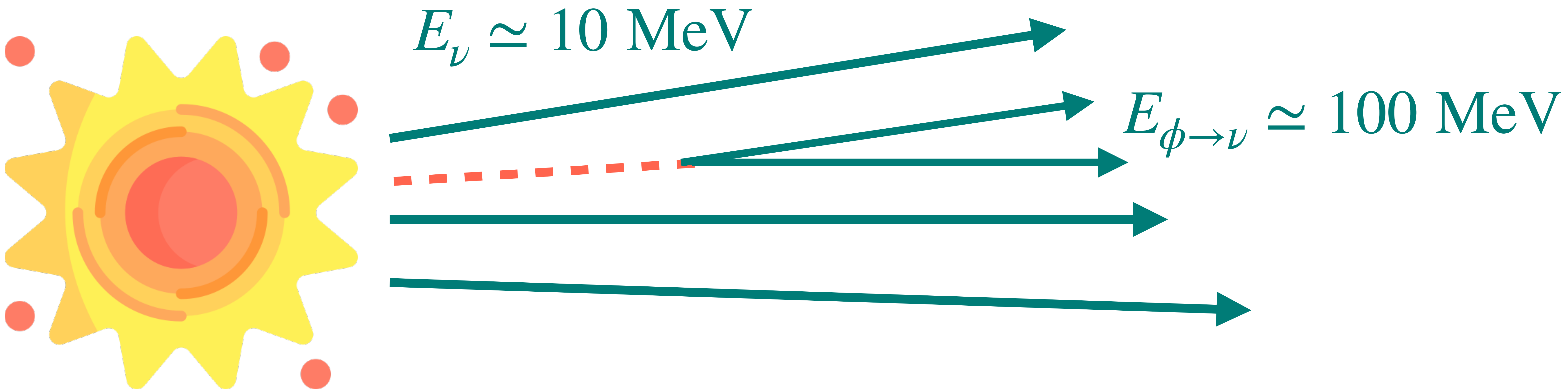
For decay to photons

- Non-observation of γ (Jaeckel et al., 1702.02964)
- Non-observation of X/γ (**DF** et al., 2303.11395, 2305.10327)
- Energy deposition in low-energy SNe (Caputo et al., 2201.09890)

Testing for new physics

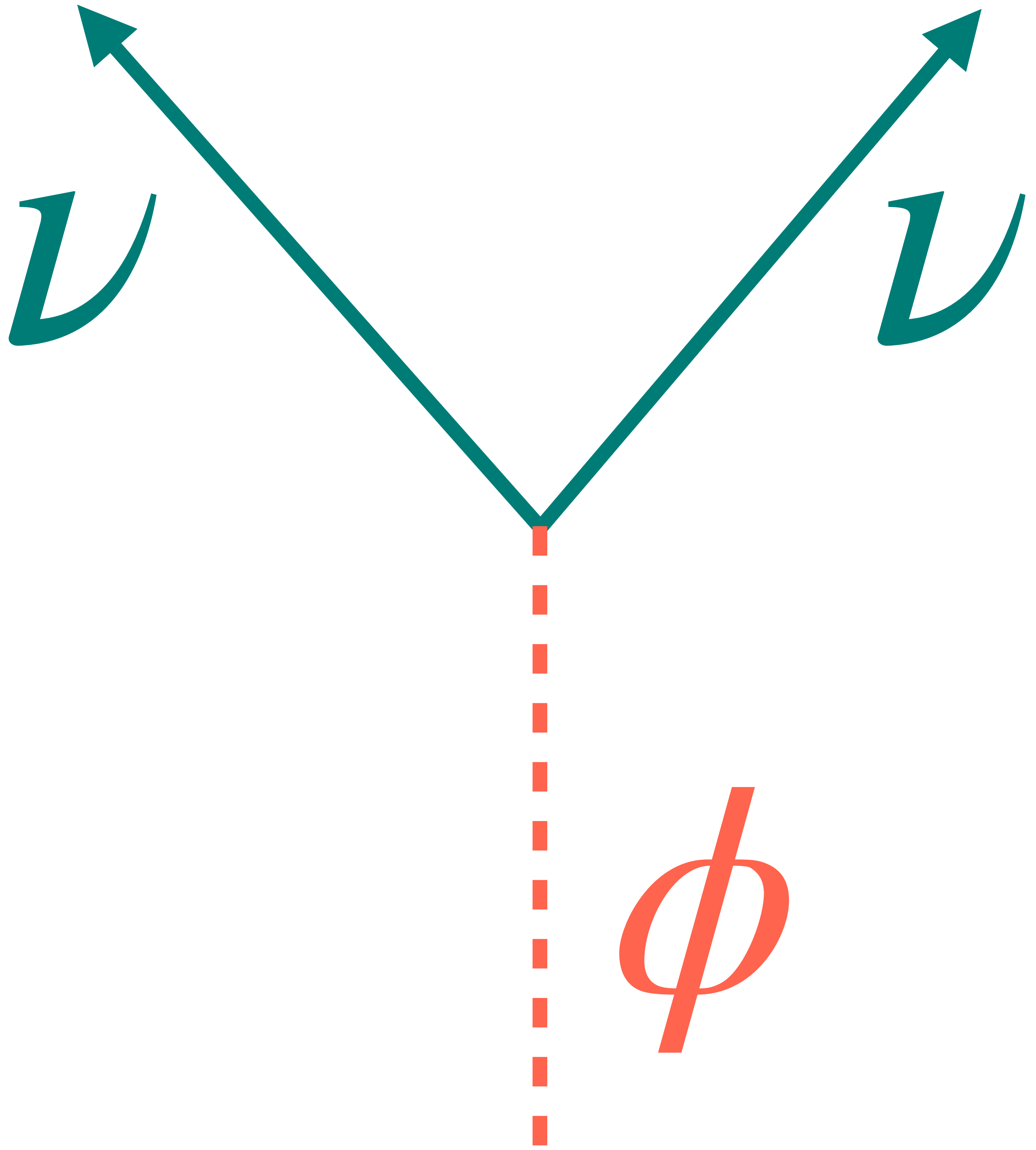
New particles can be produced in supernova core…

… but how do we probe them?



$$\frac{N^{\rm evts}_{\phi\to\nu}}{N^{\rm evts}_{\nu}} \simeq \frac{\Phi_{\phi\to\nu}}{\Phi_\nu}\frac{\sigma_{\nu N}(E_{\phi\to\nu})}{\sigma_{\nu N}(E_\nu)} \simeq \frac{L_\phi/E_{\phi\to\nu}}{L_\nu/E_\nu}\frac{E^2_{\phi\to\nu}}{E^2_\nu} \simeq 10\frac{L_\phi}{L_\nu}$$

(Pseudo)-majorons



$$\mathscr{L} = \frac{g}{2}\bar{\nu}^{\mathrm{c}}\nu\phi$$

In supernova, neutrino-neutrino and antineutrino-antineutrino coalescence

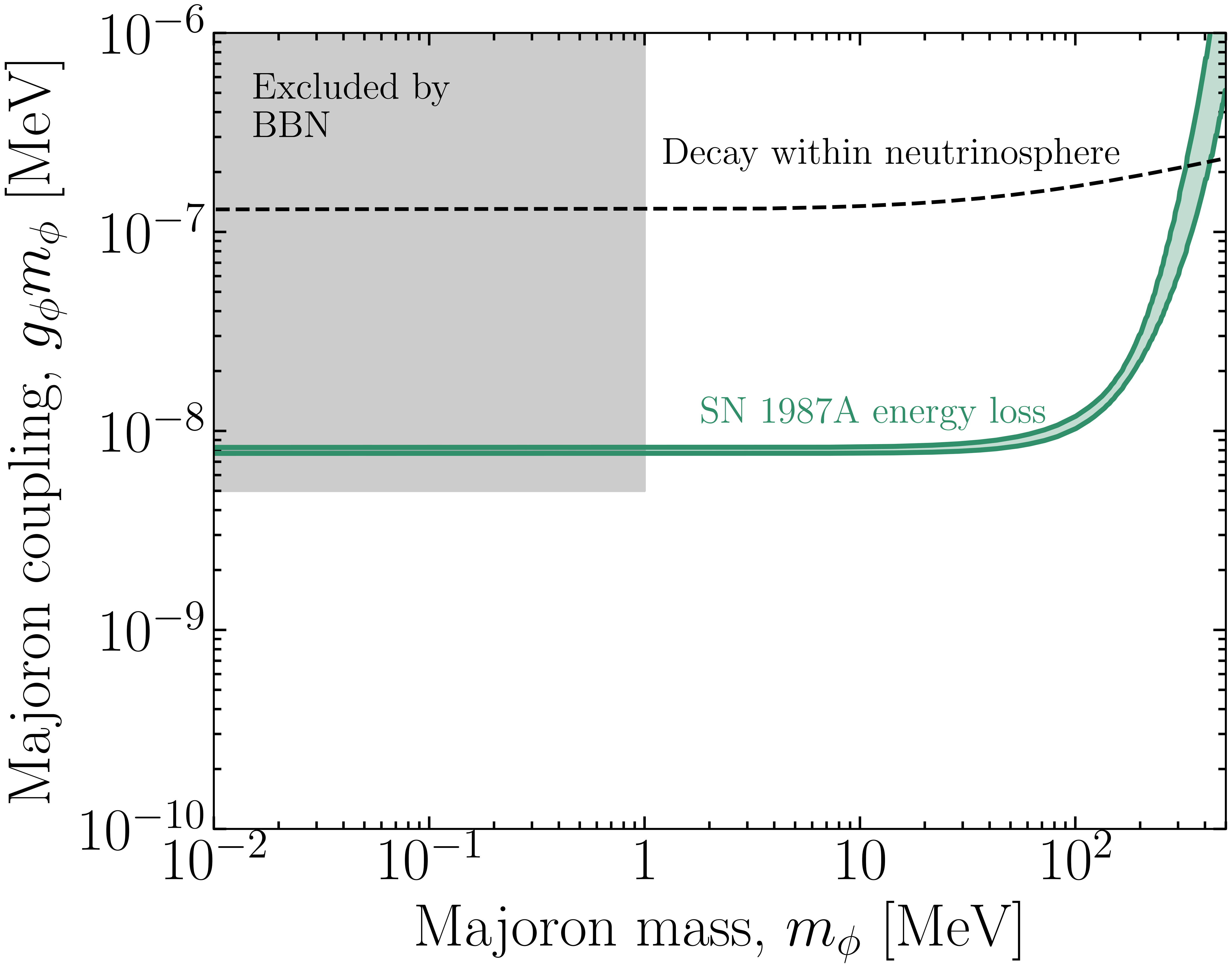
$$m_\phi \gtrsim 10^{-4}\ \mathrm{MeV}$$

Chicashige, Mohapatra, Peccei (1981);
Gelmini, Roncadelli (1981)

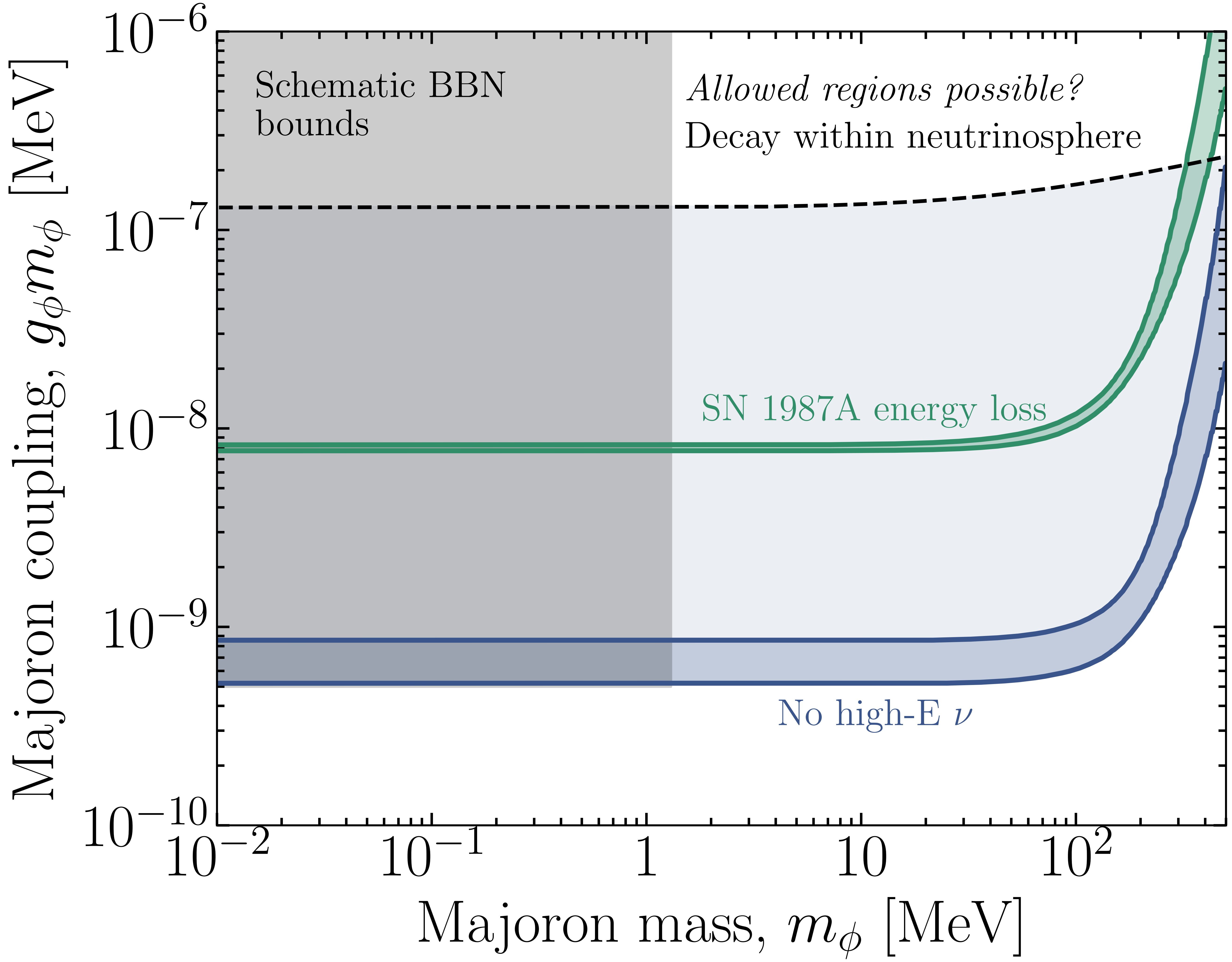
Majoron production

For small masses, signal depends only on gm_ϕ

$$\frac{dN_\phi}{dt} = \frac{(gm_\phi)^2\mu_\nu^2}{192\pi^3}$$



Novel bounds



$L_\phi < L_\nu/100$

Impact on supernova explosion ruled out

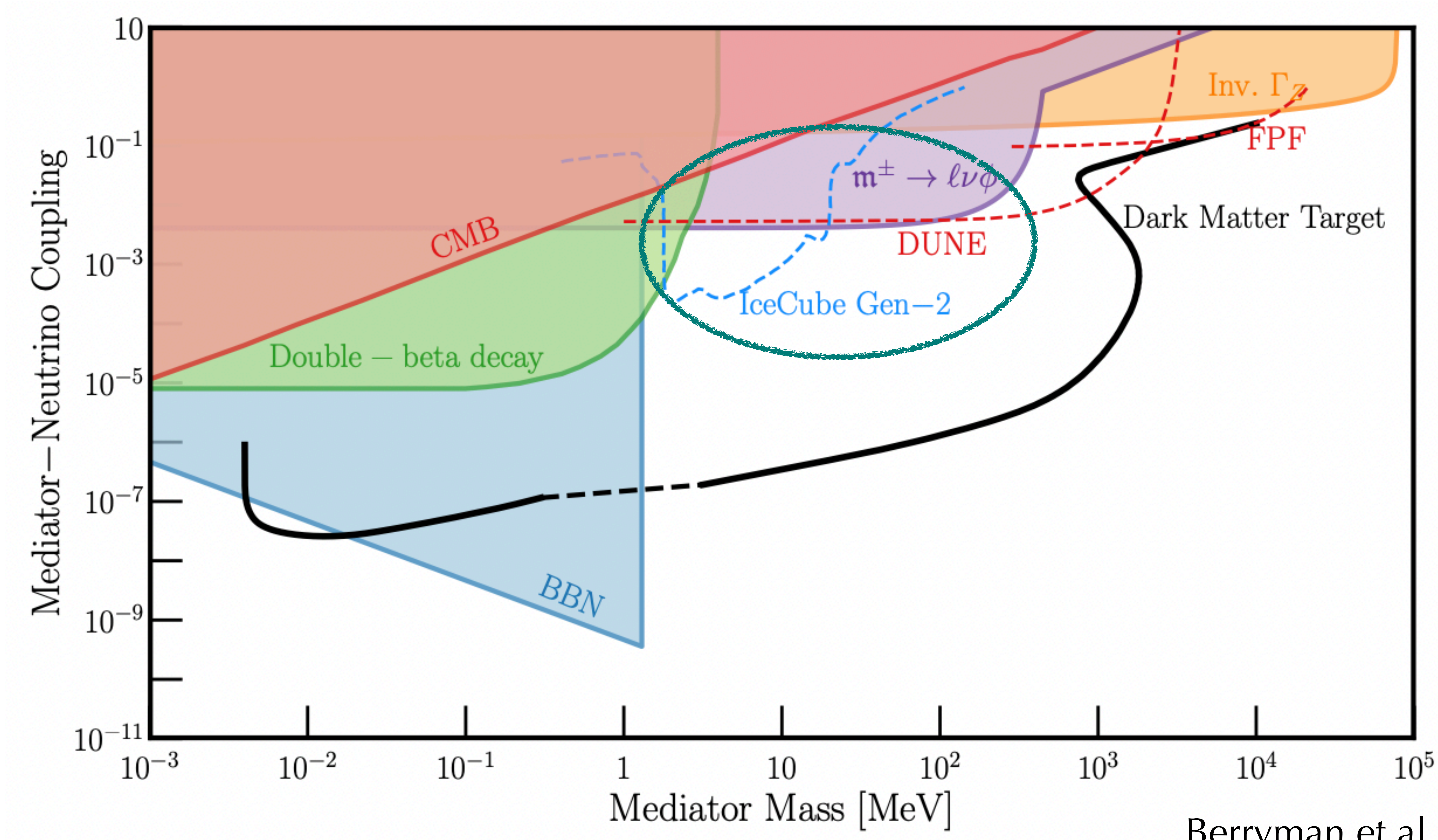
Next galactic supernova in Akita et al., arXiv:2206.06852

Application to sterile neutrino in Brdar et al., arXiv:2302.10965

Secret interactions in supernovae

based on arXiv:2307.15122 (accepted at Phys. Rev. D) and arXiv:2307.15115 (submitted to Phys. Rev. Lett.)
with G. Raffelt, E. Vitagliano

Secret interactions

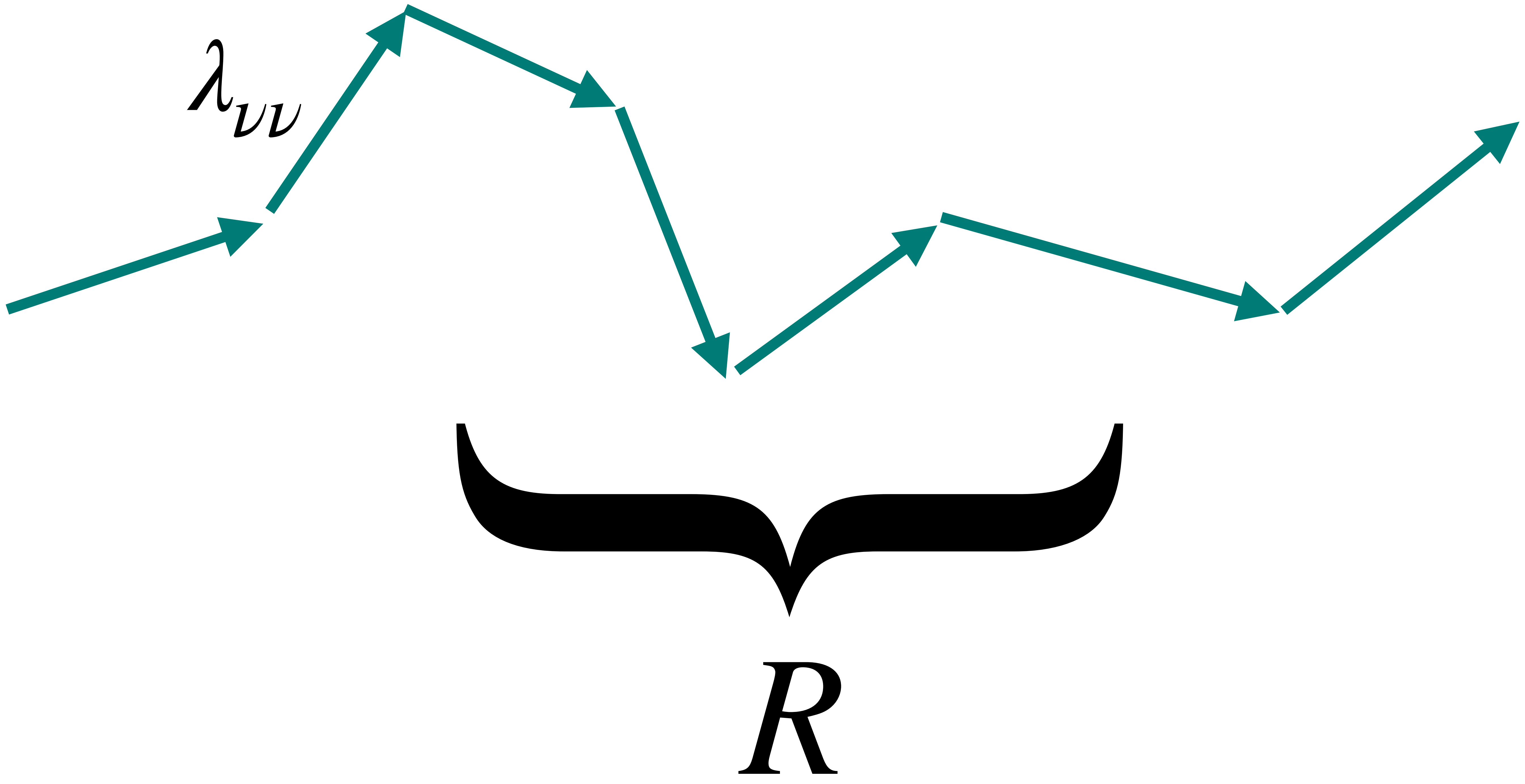


$$\mathscr{L} \propto \bar{\nu}\gamma^{\mu}\nu Z'_{\mu}$$

Berryman et al., 2203.01955

Secret interactions (νSIs) and SNe

- **Manohar (1987)**: νSI delay ν



$$\delta t \sim R\frac{R}{\lambda_{\nu\nu}}$$

Secret interactions (νSIs) and SNe

- **Manohar (1987)**: νSI delay ν
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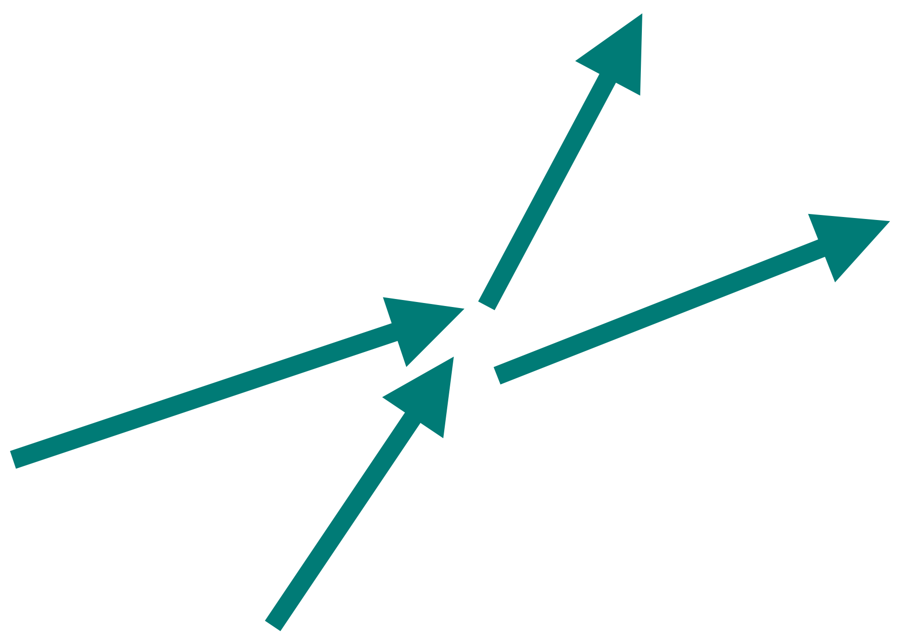
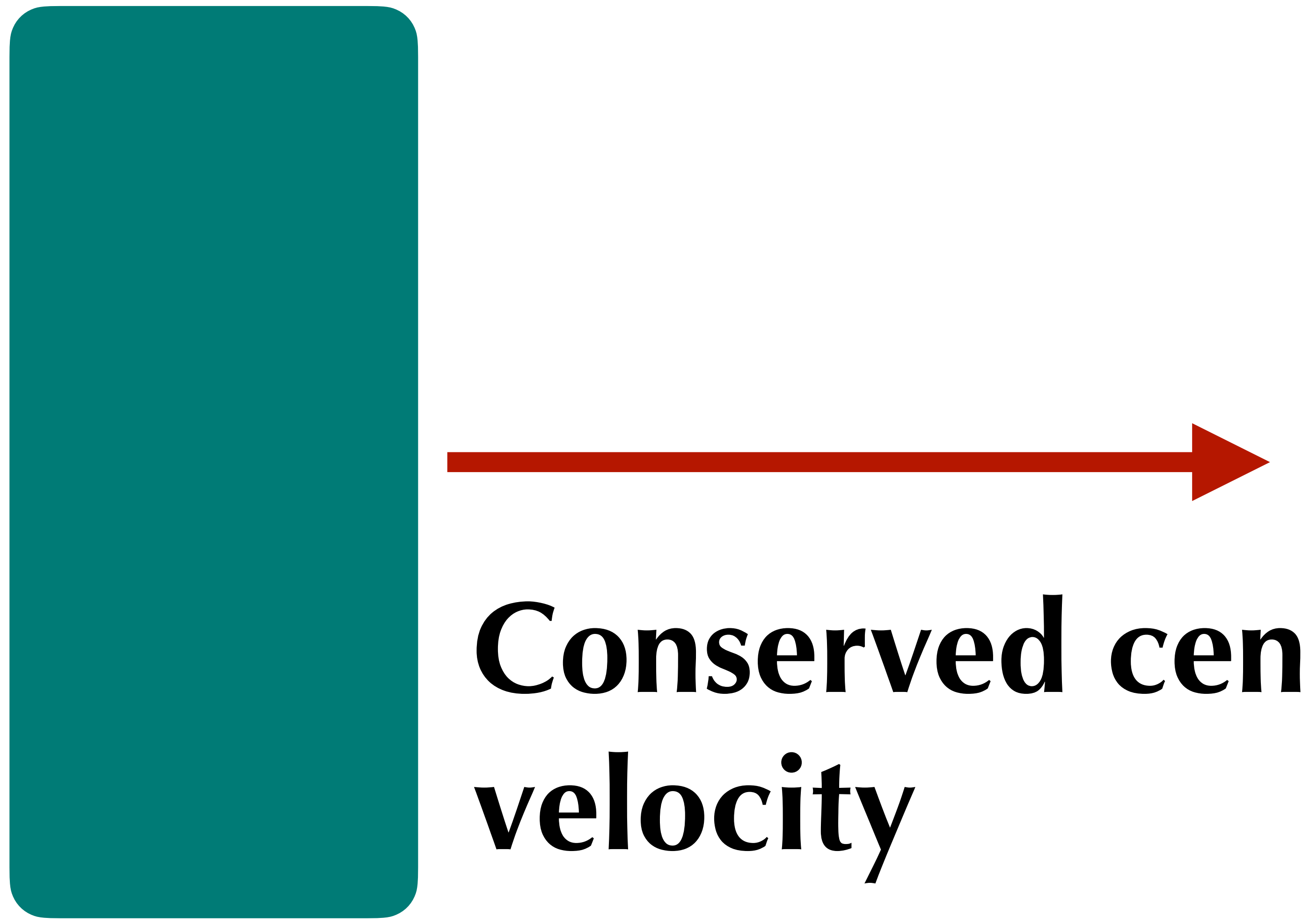


Total momentum conserved



Conserved center-of-mass velocity

Secret interactions (νSIs) and SNe

- **Manohar (1987)**: νSI delay ν
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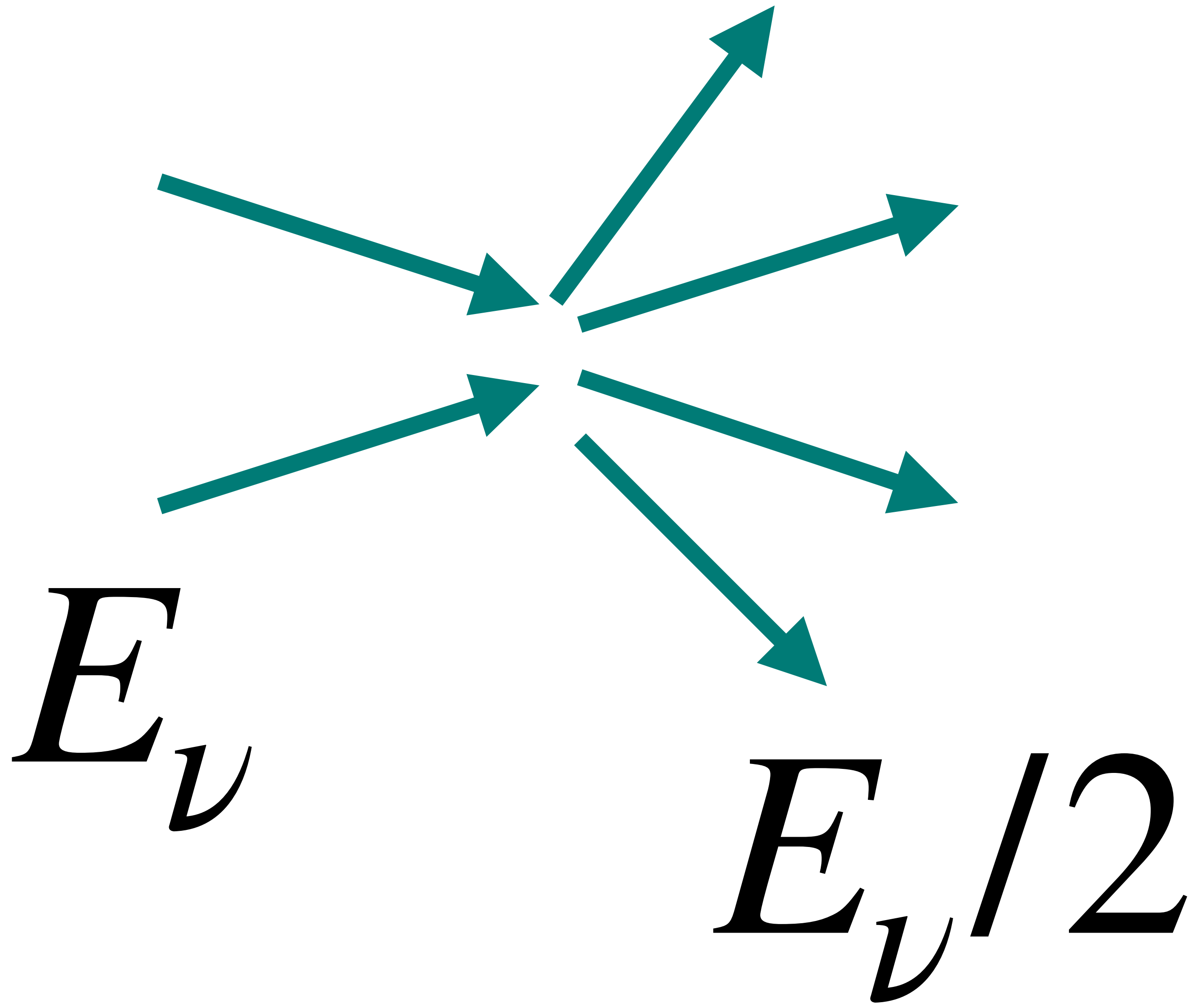
- **Shalgar, Tamborra, Bustamante (2019)**: $\nu\bar{\nu} \rightarrow \nu\bar{\nu}\nu\bar{\nu}$ reduce ν energy



$$\frac{E_\nu}{2^{R/\lambda_{\nu\nu}}}$$

Secret interactions (νSIs) and SNe

- **Manohar (1987)**: νSI delay ν
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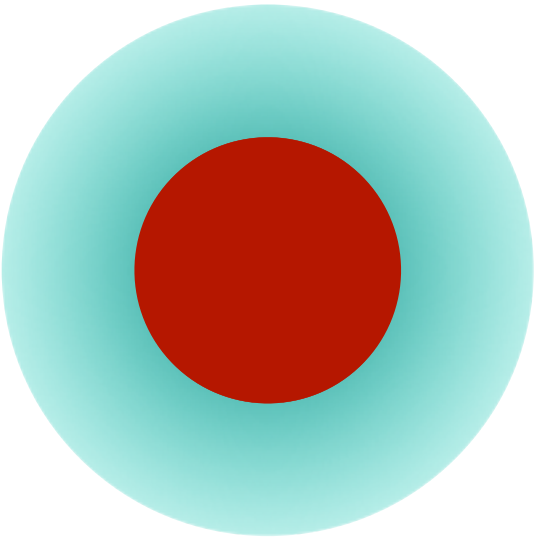
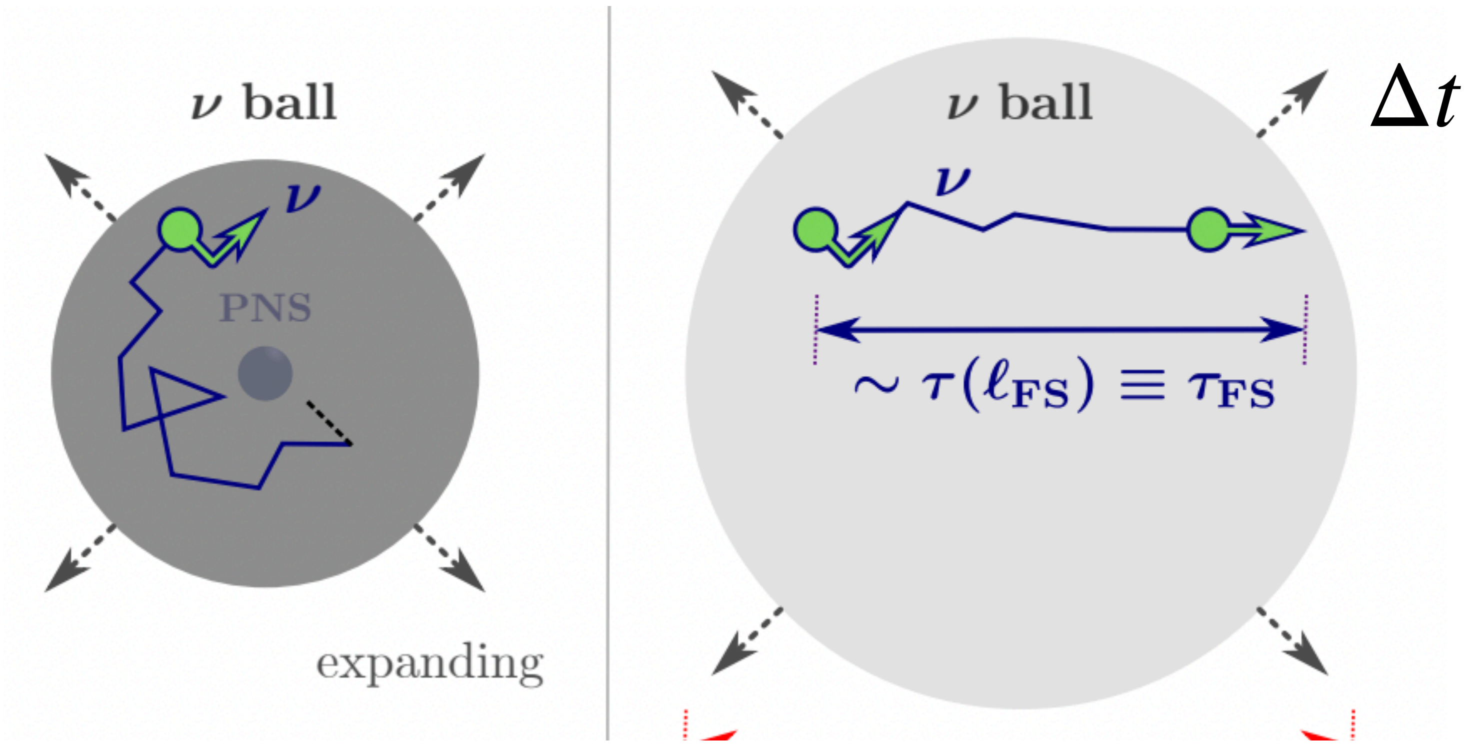
- **Chang, Esteban, Beacom, Thompson, Hirata (2022)**: how does ν fluid escape?

Wind outflow

Steady flow - incomplete picture

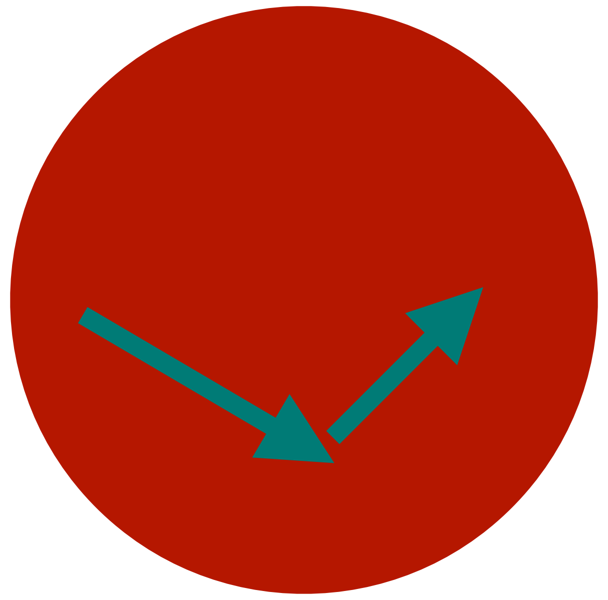




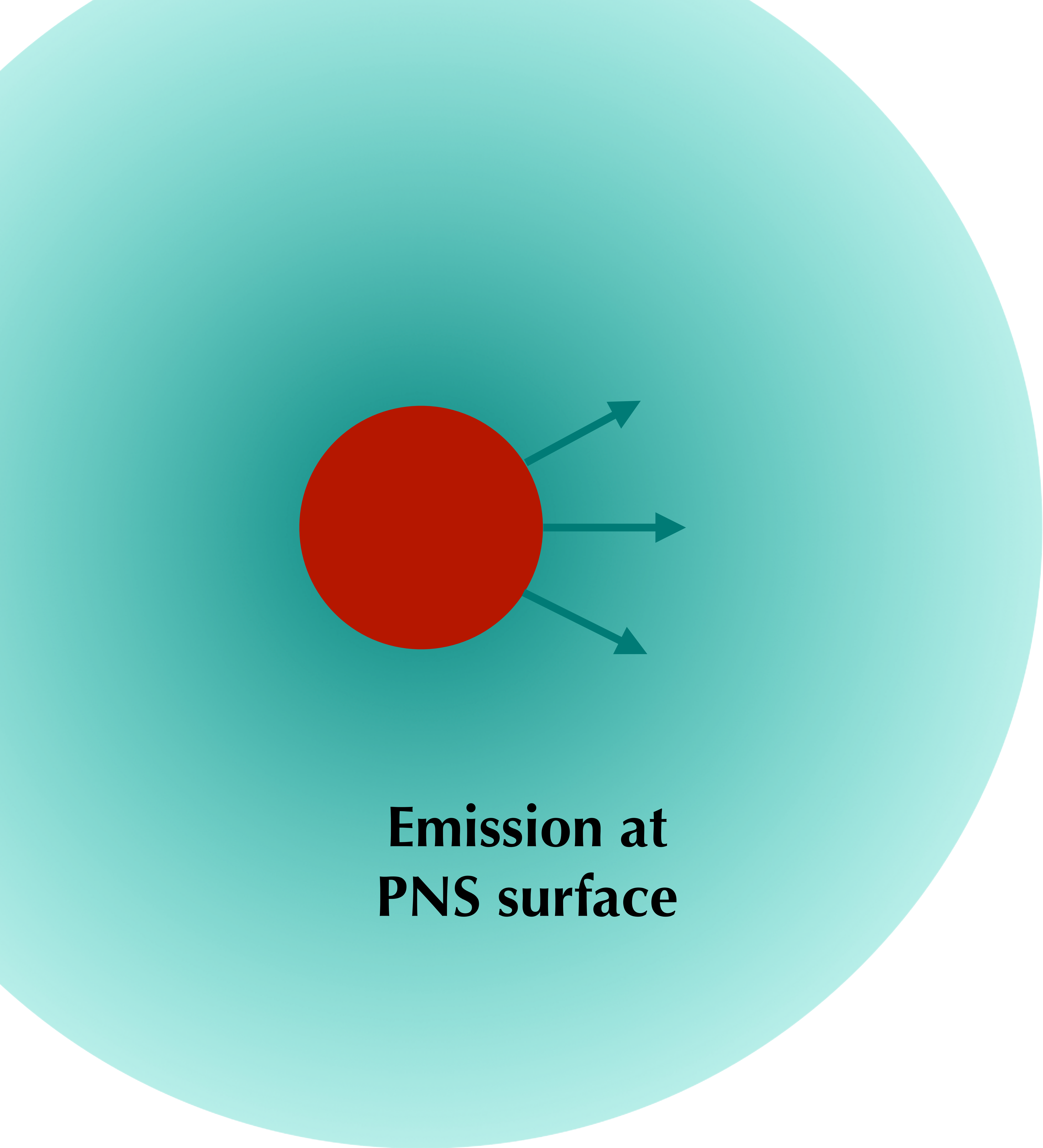
$$\Delta t \sim R\sqrt{\frac{R}{\lambda_{\nu\nu}}}$$

Burst outflow

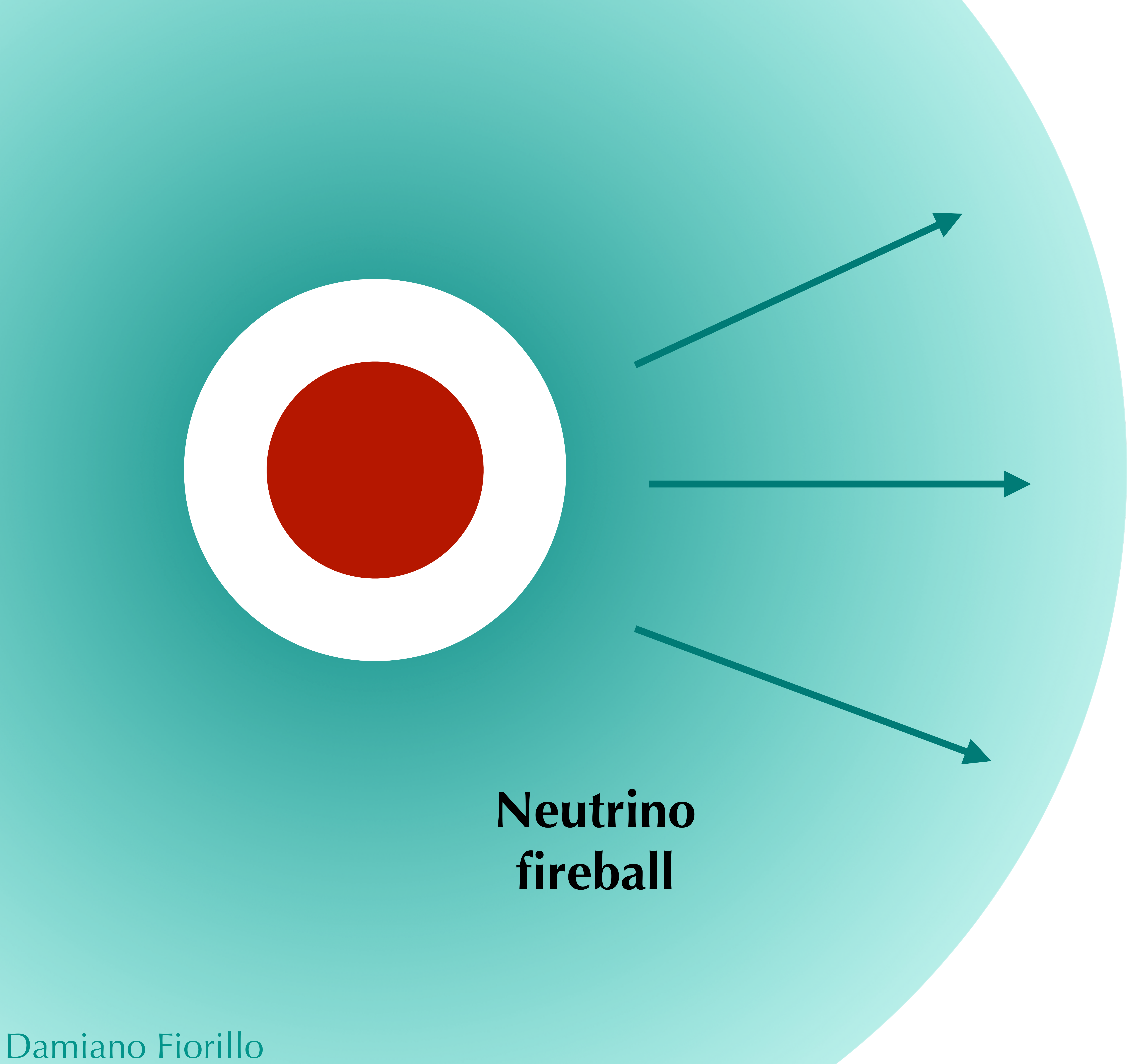
Numerical simulations in conflict with previous literature



Heat transport inside PNS



Emission at PNS surface

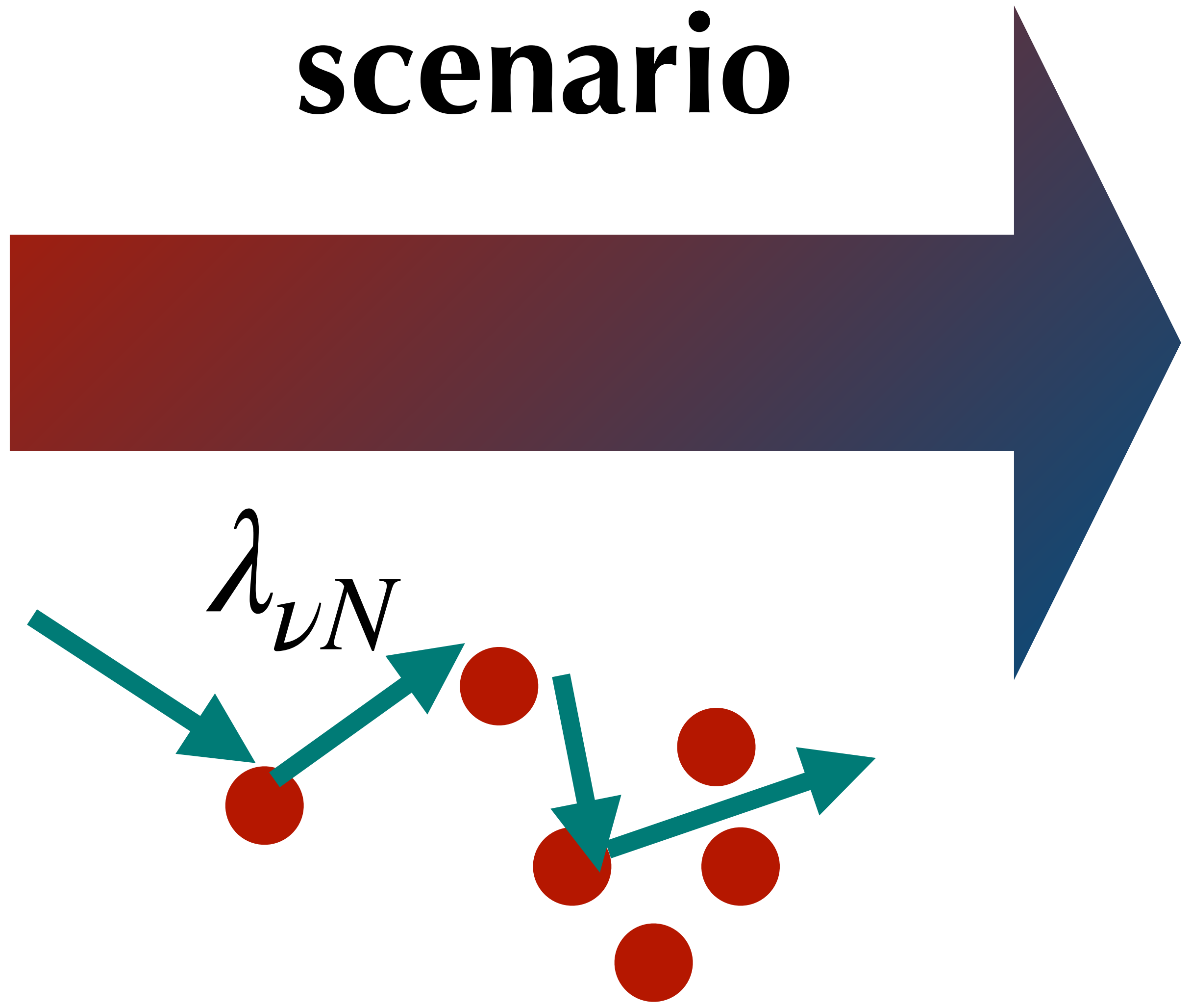

Neutrino
fireball

Decoupling

Damiano Fiorillo

Heat transport inside PNS

Standard scenario

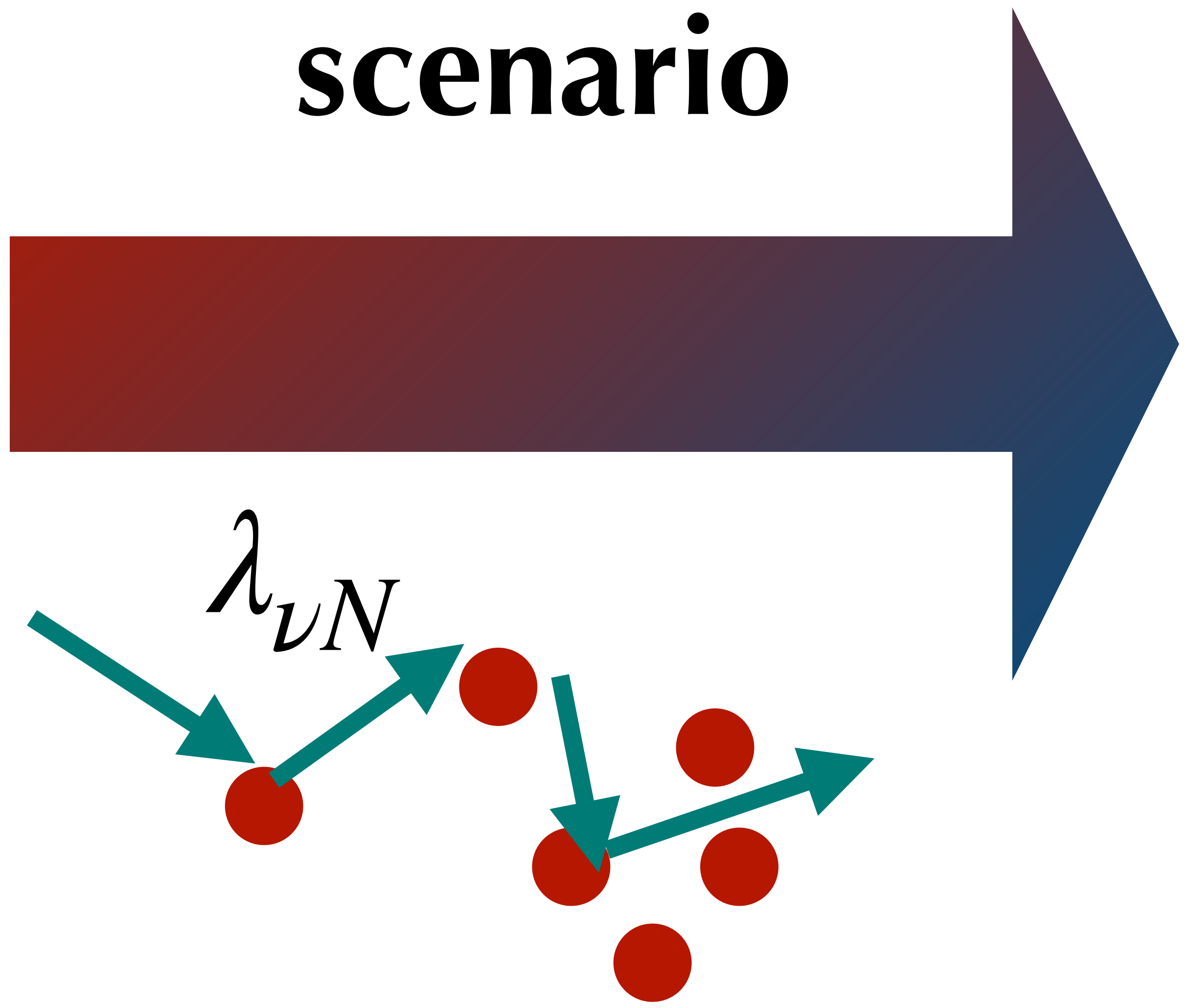


$$F \propto \lambda_{\nu N} \nabla \rho_\nu$$

Temperature gradients
induce anisotropy

Heat transport inside PNS

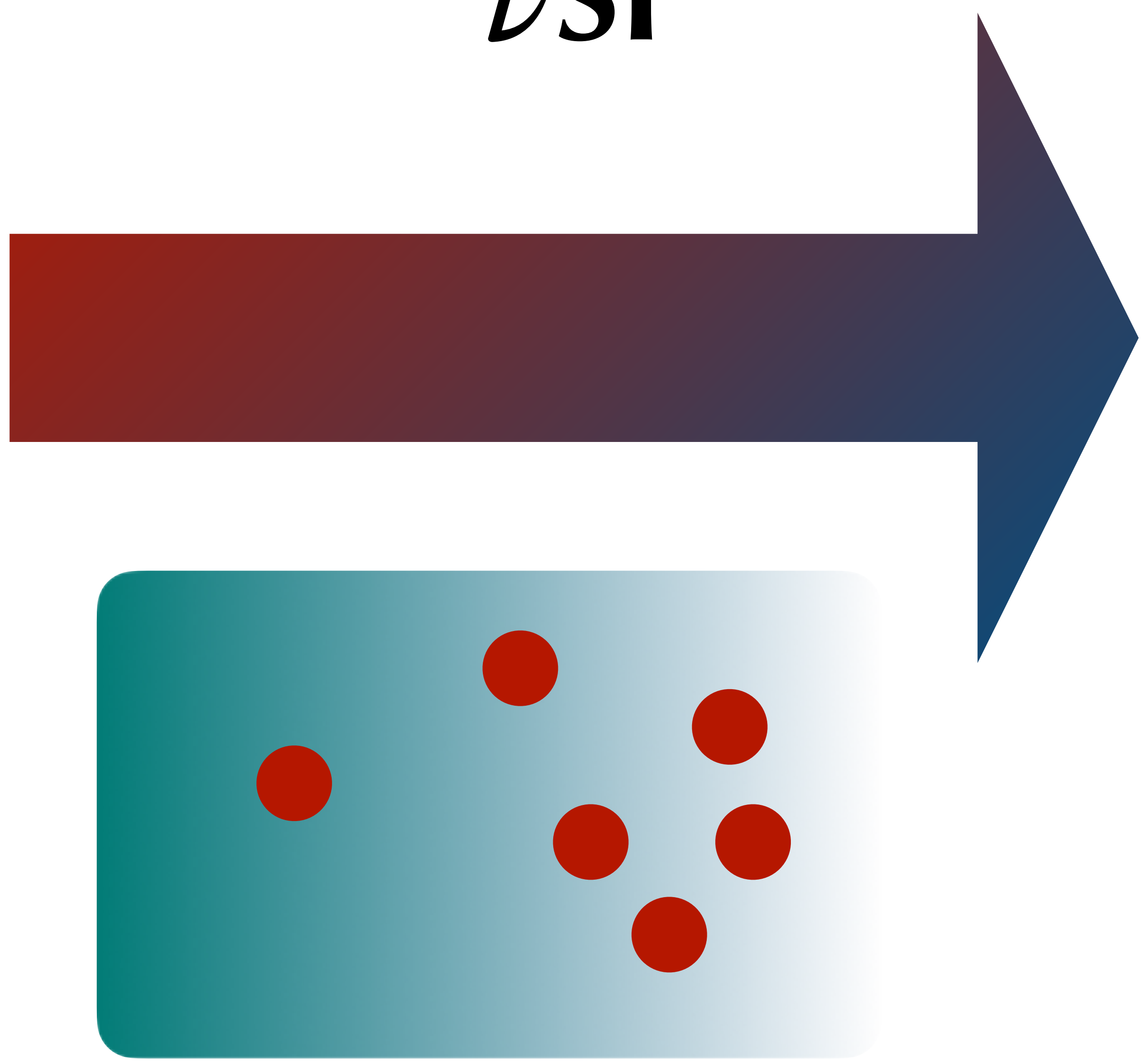
Standard scenario



$$F \propto \lambda_{\nu N} \nabla \rho_\nu$$

Temperature gradients induce anisotropy

νSI

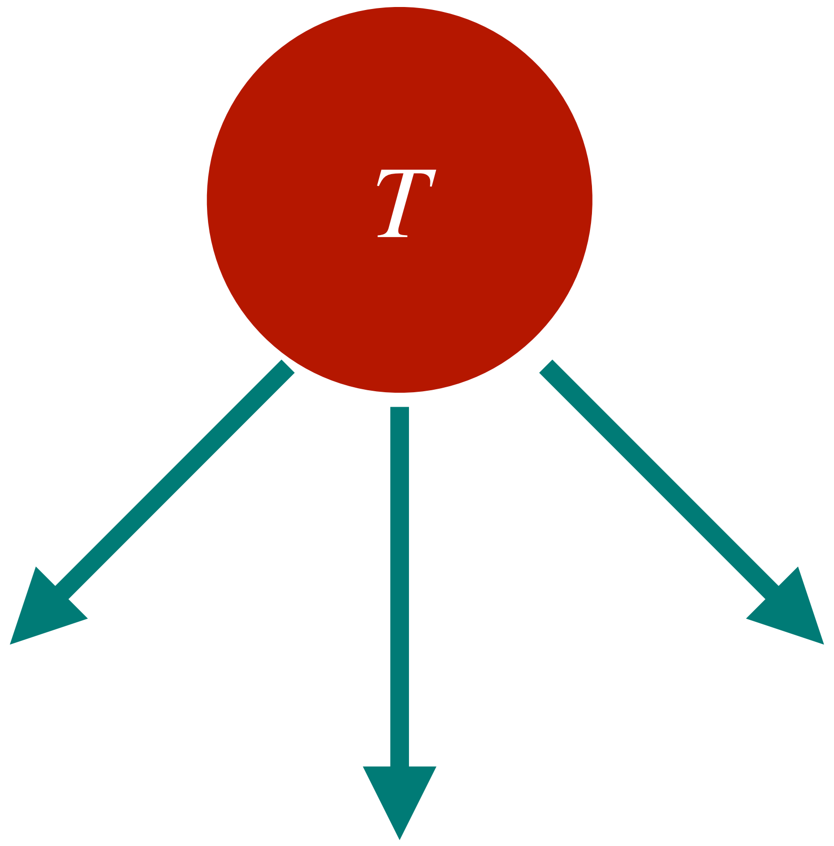


$$\nabla p \sim \frac{\rho_\nu v}{\lambda_{\nu N}} \longrightarrow F = \rho_\nu v \propto \lambda_{\nu N} \nabla \rho_\nu$$

νSI convert anisotropy in bulk motion

Emission at the PNS surface

Blackbody emission of a fluid

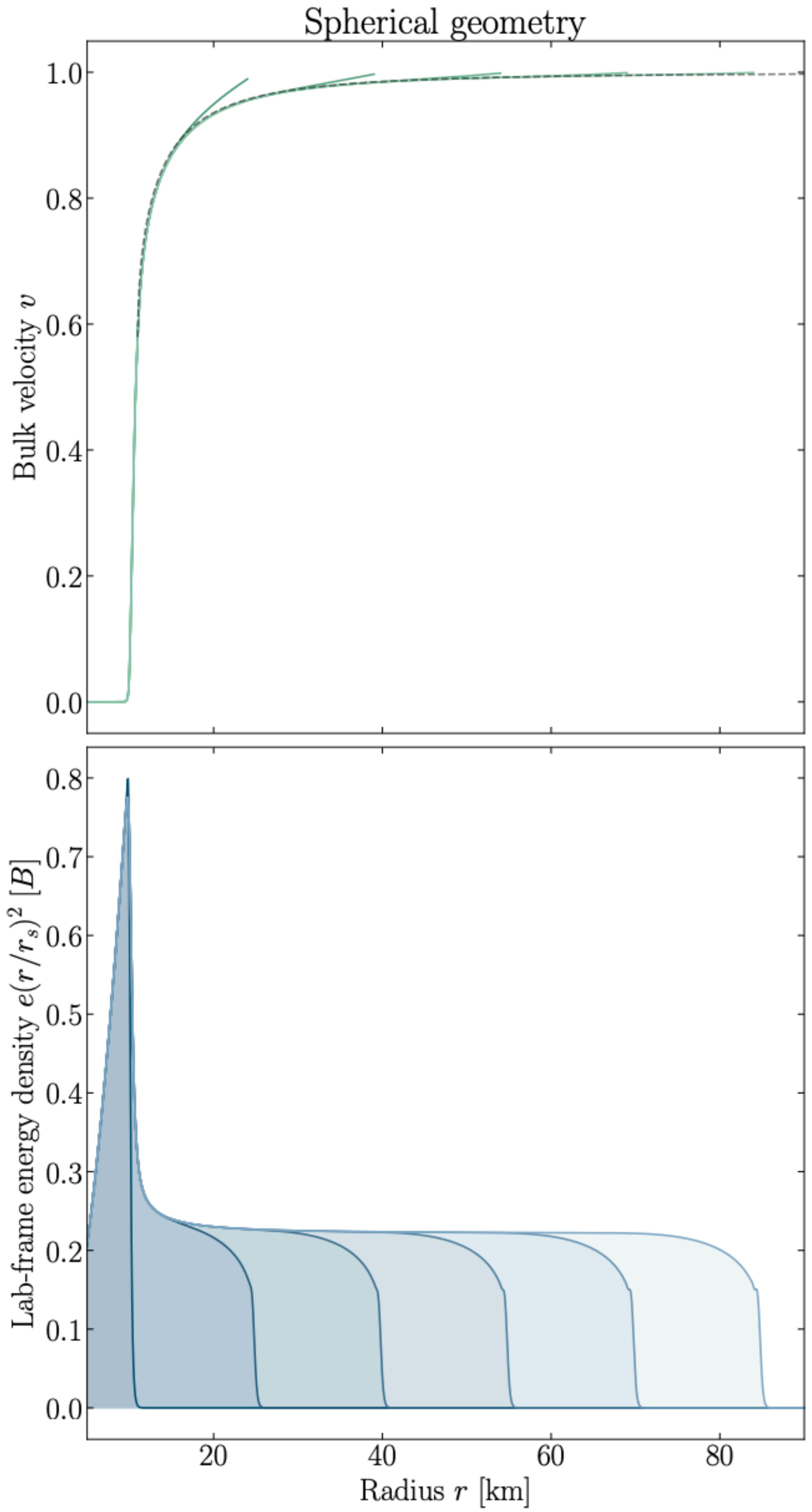


$$T_\nu \propto r^{-1}, \gamma \propto r$$

$$\bar{\epsilon} = 3.48T$$

Compare with standard

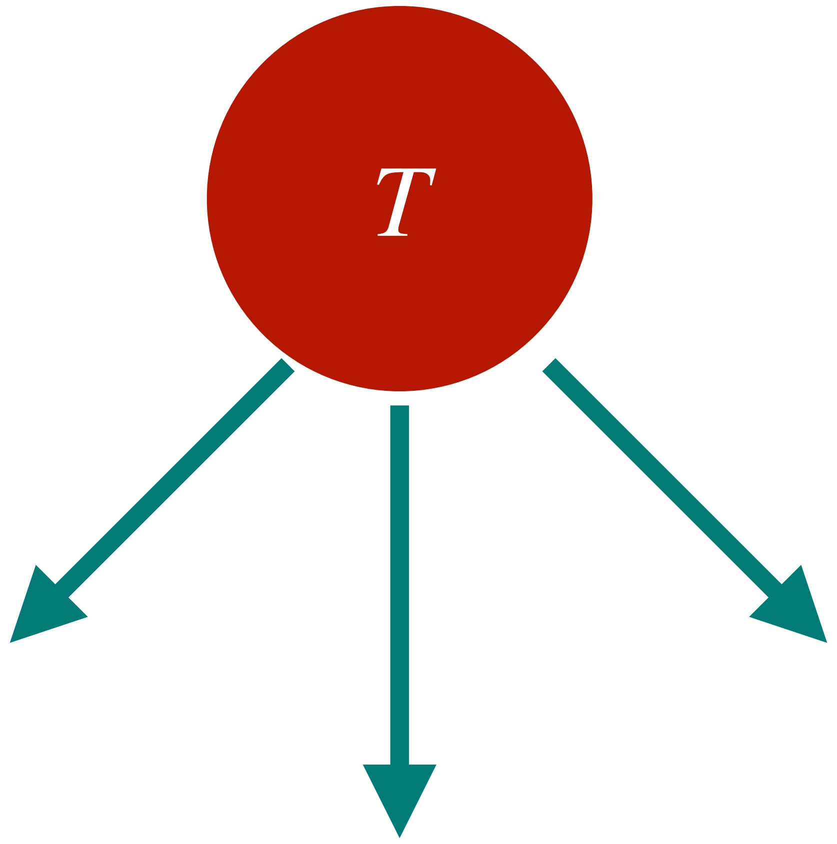
$$\bar{\epsilon} = 3.15T$$

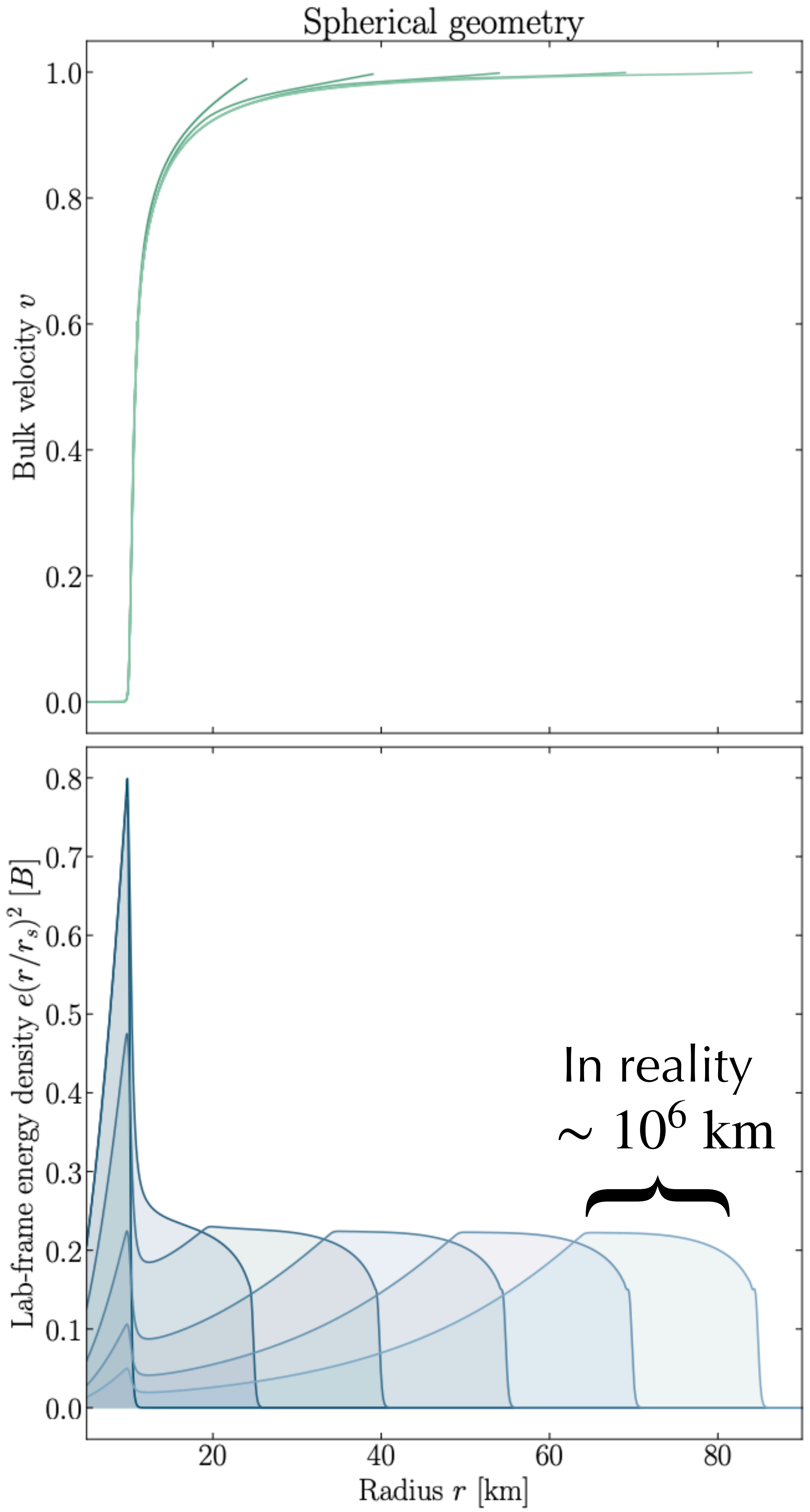


- ✦ Front emission moves with the speed of light
- ✦ Quasi-steady emission for a time $\sim 5\ \mathrm{s} \sim 10^6\ \mathrm{km} \gg R_{\mathrm{PNS}}$
- ✦ Bulk of fluid moves with speed of light

Neutrino fireball

Blackbody emission of a fluid



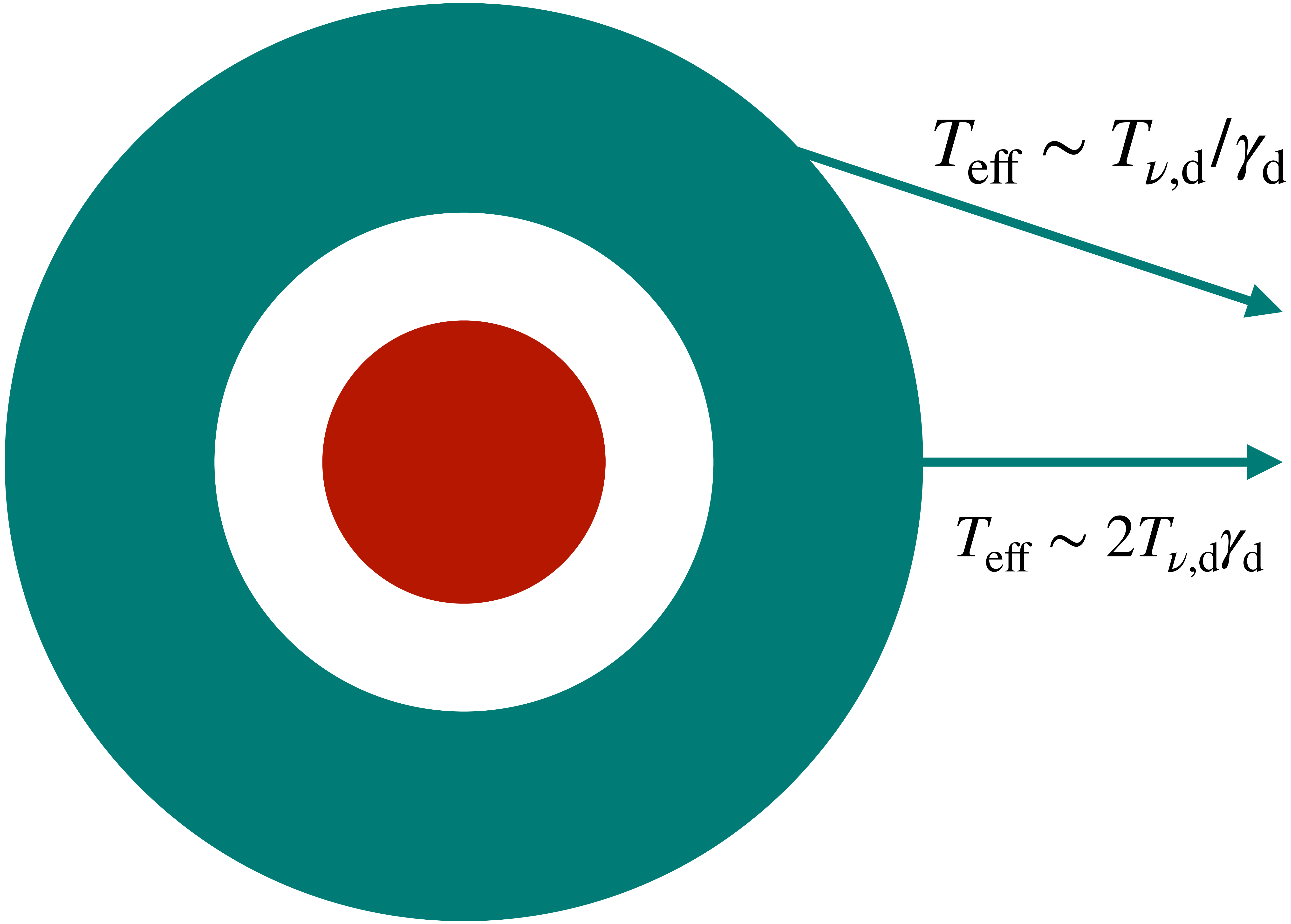


In reality $\sim 10^6$ km $\overbrace{}$

- If PNS turns off, bulk of the fluid moves so fast it does not perceive it (sound horizon)
- No change in ν thickness, since most fluid is moving with $v \sim 1$
- $\nu\bar{\nu} \rightarrow \nu\bar{\nu}\nu\bar{\nu}$ balanced by inverse reaction — detailed balance wants $\mu_\nu = -\mu_{\bar{\nu}} = 0$, already satisfied

No observable affected by large ratio $R/\lambda_{\nu\nu}$!

Decoupling



- Strongly boosted blackbody
- Superposition of different directions lead to

$$\frac{d\Phi}{d\epsilon} \propto \epsilon \log\left[1 + e^{\eta-\epsilon/2\overline{T}}\right]$$

- Obtained in **DF** et al., 2303.11395, 2305.10327

BSM conclusions

- SN 1987A ideal laboratory for BSM physics in neutrino sector
- Cooling bound complemented by decay bounds — stronger by 1 order of magnitude, more robust
- Trapped mediators produce νSI, but only $O(1)$ changes to observables